Centre Canadien de Recherche en Physique des Astroparticules
Arthur B. McDonald
Canadian Astroparticle Physics Research Institute

Neutrinos & Dark Matter & Dark Matter & Neutrinos

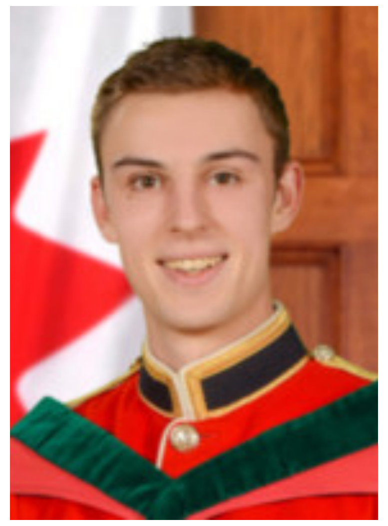
Aaron C. Vincent

KCL TPPC Weekly Seminar 21/04/2021

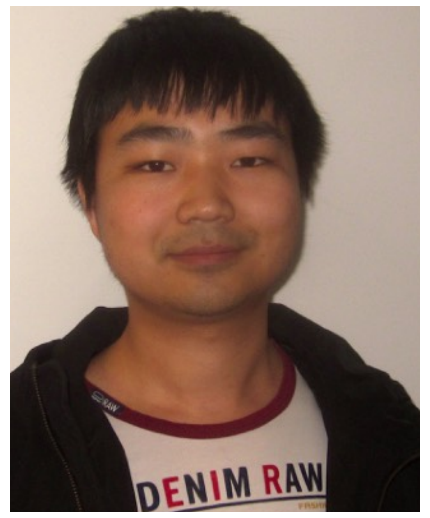
Featuring

Carlos Argüelles
Harvard

Ali Kheirandish
Penn State



Adam
McMullen



Ningqiang
Song

Mauricio Bustamante
Niels Bohr Institute

Shirley Weishi Li
Fermilab

Andrés Olivares-del-Campo
IPPP Durham

Moving to Liverpool in September you should talk to him

Alejandro Diaz
MIT

Katie Mack
NCSU

Ibrahim Safa
Harvard

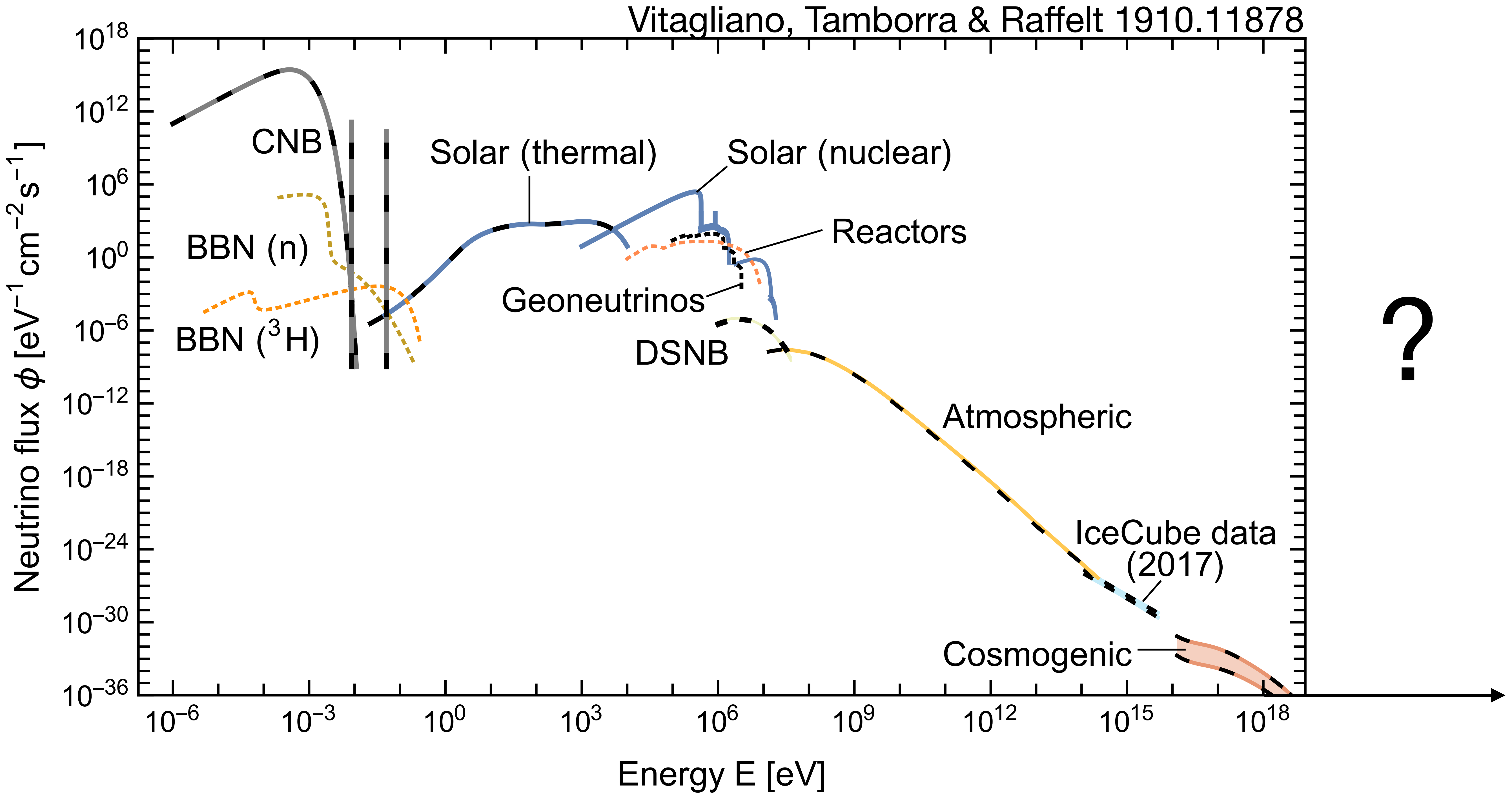
Overview

1. The neutrino sky today & in the future
2. New physics
    i. Neutrino decay
    ii. Dark matter
    iii. Large extra dimensions
3. Conclusions

The neutrino sky

Neutrinos from space carry:

- Directional information
- Timing information
- Energy
- **Flavour:** ν_e, ν_μ, ν_τ



Now have experiments that cover this entire range.
What can we learn from extraterrestrial neutrinos?
Where do we go?

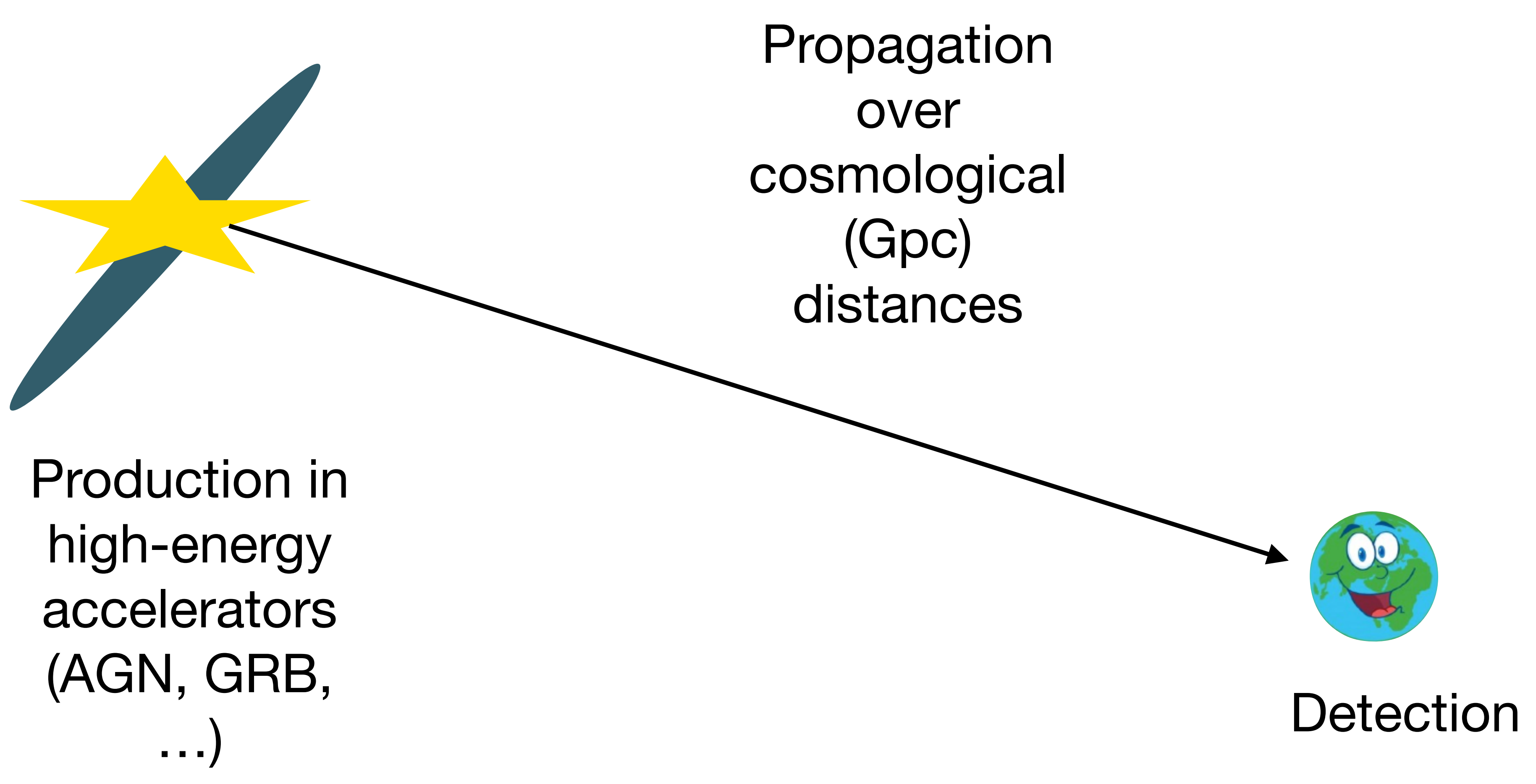
High energies & Flavour

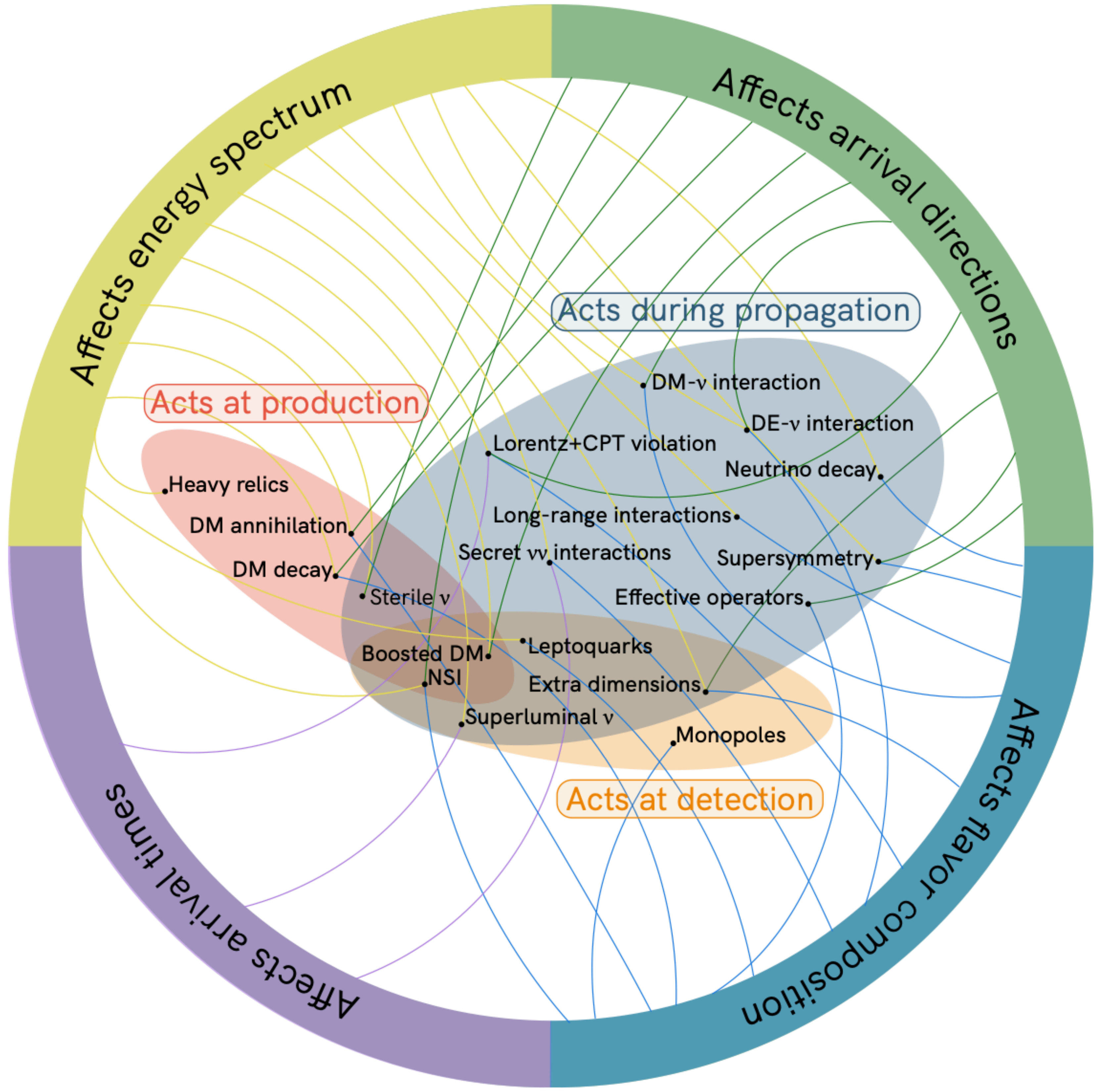
The Future of High-Energy Astrophysical Neutrino Flavor Measurements

Ningqiang Song, Shirley Weishi Li, Carlos A. Argüelles, Mauricio Bustamante, Aaron C. Vincent

[Accepted/JCAP] https://arxiv.org/abs/2012.12893

High-energy neutrinos





Neutrinos can tell us about "standard model" physics:

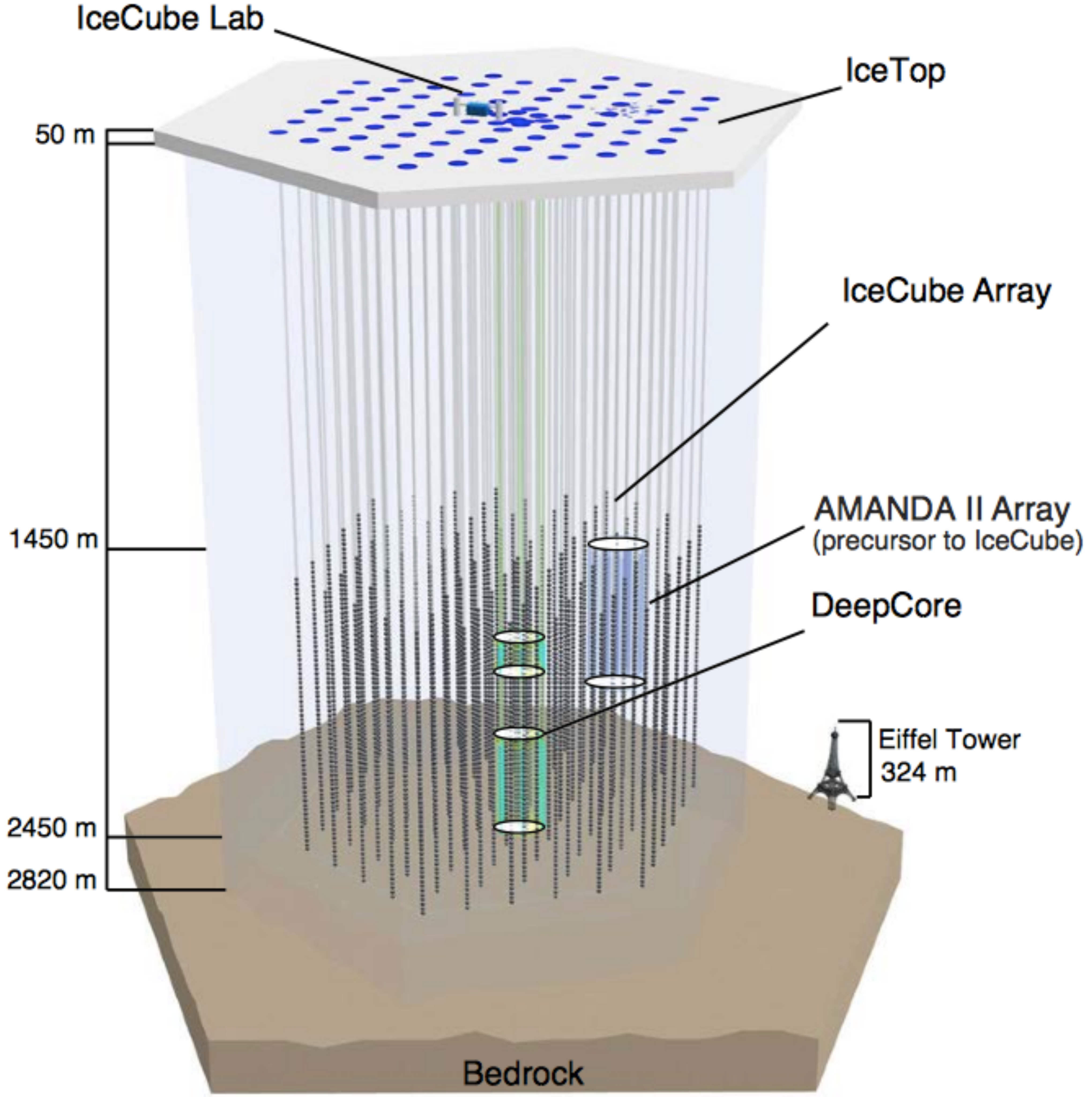
- Nature of these accelerators
- Oscillation, interaction with intergalactic medium
- Detection: high-energy neutrino-nucleus cross sections

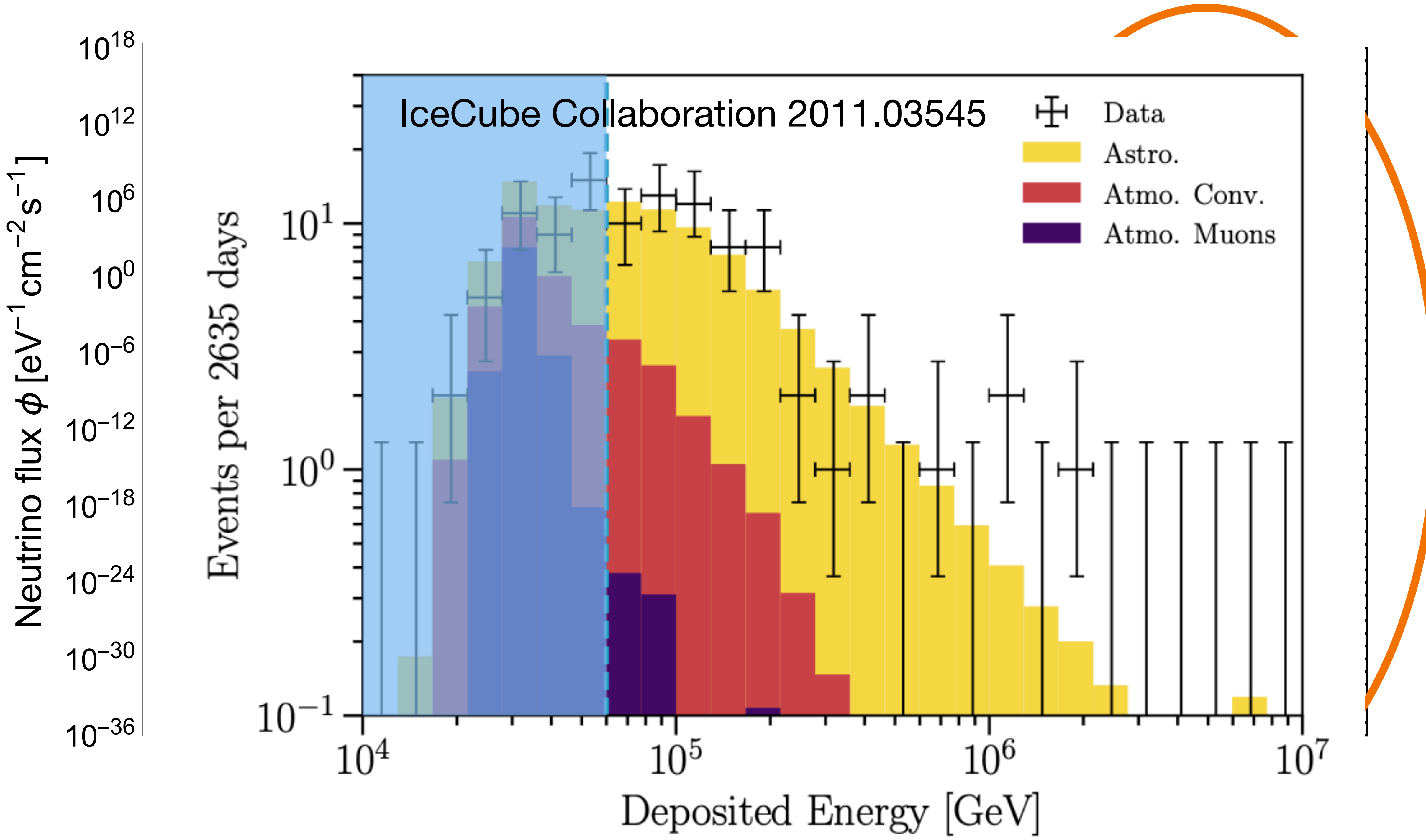
New Physics?

1907.08690 Argüelles,Bustamante,Kheirandish,Palomares-Ruiz,Salvado, ACV

Current observations: IceCube (south pole)

Effective volume $\sim 1\,\mathrm{km}^3$

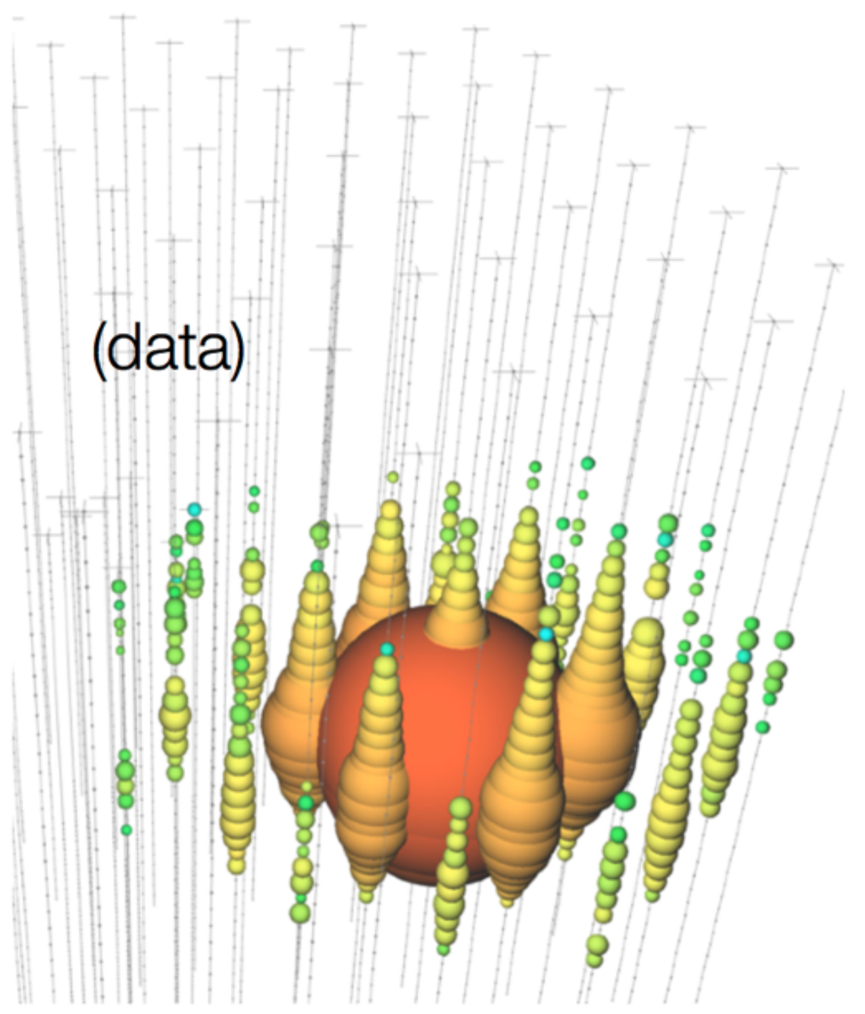




large exposures necessary due to low fluxes

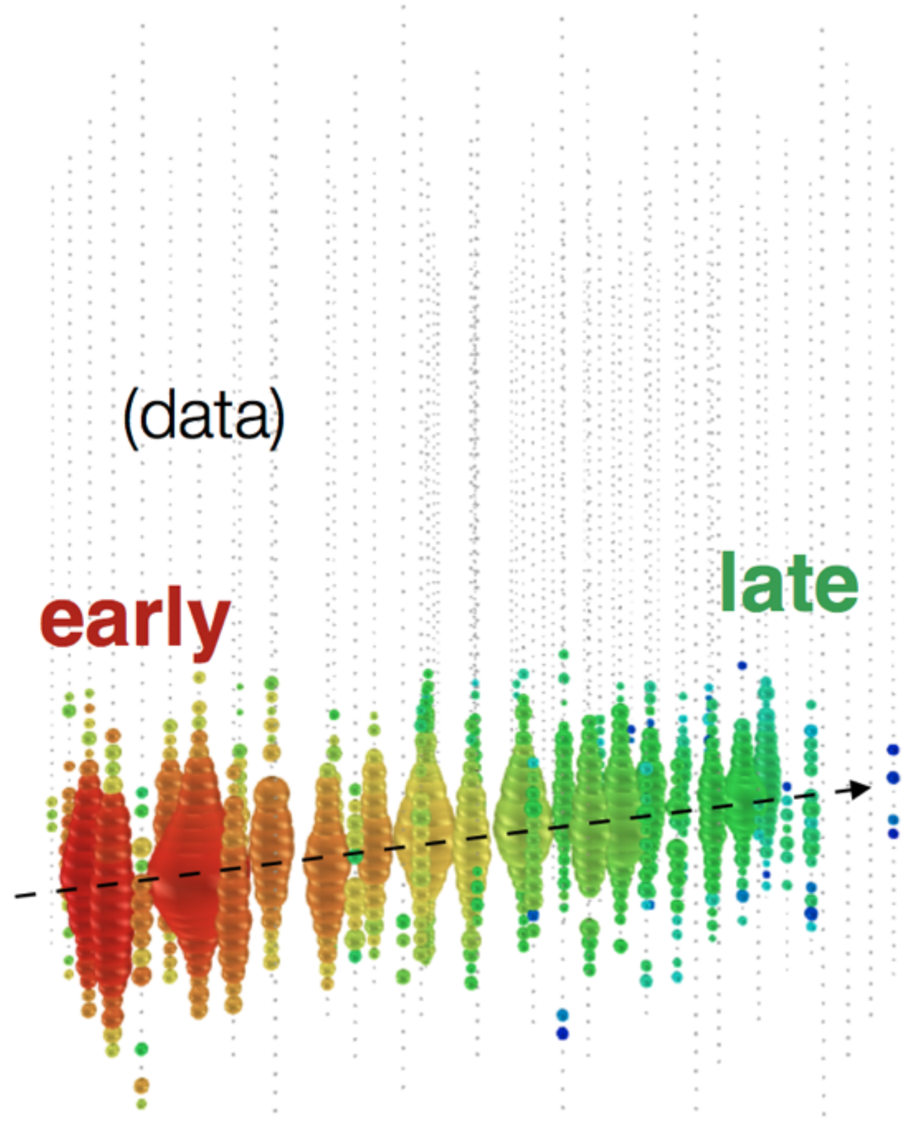
Flavour: event morphology

Neutral-current / ν_e



Isolated energy deposition (cascade) with no track

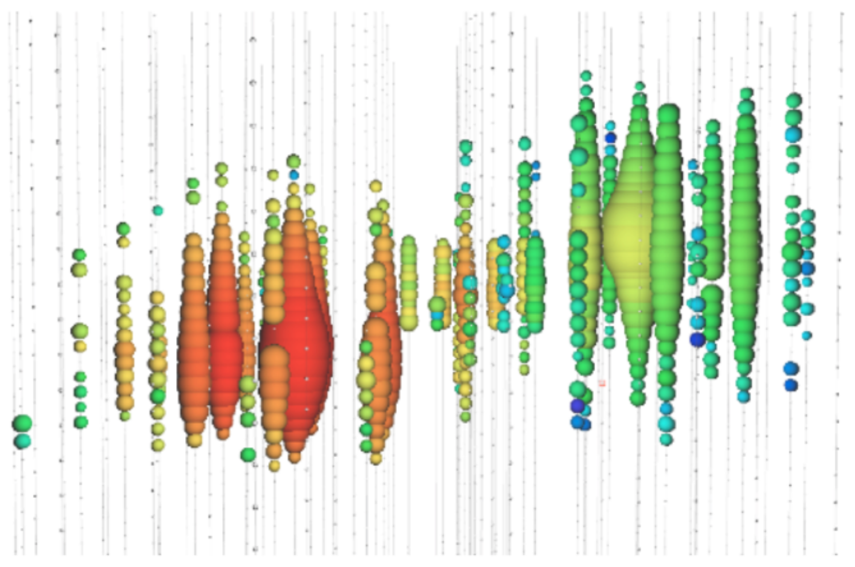
Charged-current ν_μ



Up-going track

Charged-current ν_τ

(simulation)



Double cascade

Images: icecube.wisc.edu

Flavour composition in astrophysical sources

(GRBs, AGNs, blazars, pulsars…)

$(\alpha_e : \alpha_\mu : \alpha_\tau)$

Pion sources (c.c. for π^-)

$$\pi^+ \to \mu^+ + \nu_\mu$$
$$\mu^+ \to e^+ + \nu_\mu + \bar{\nu}_e$$

(1 : 2 : 0)

“muon-damped” (c.c. for π^-)

$$\pi^+ \to \cancel{\mu^+} + \nu_\mu$$

(0 : 1 : 0)

“muon source” (c.c. for π^-)

$$\pi^+ \to \mu^+ + \cancel{\nu_\mu}$$
$$\mu^+ \to e^+ + \nu_\mu + \bar{\nu}_e$$

(1 : 1 : 0)

Neutron source

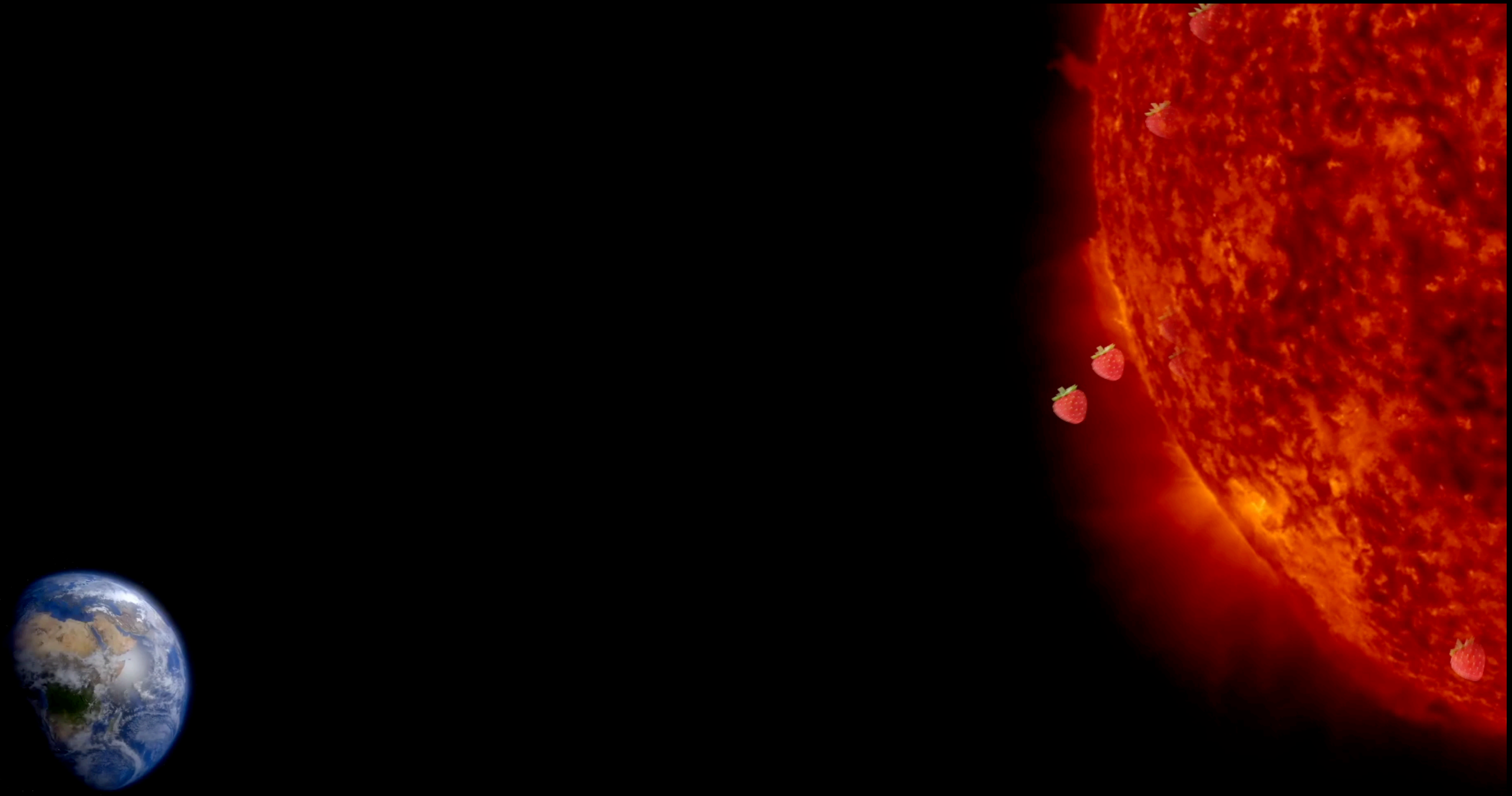
$$n \to p + e^- + \bar{\nu}_e$$

(1 : 0 : 0)

Different scenarios: different production environments

Flavour can be distinguished **statistically** in neutrino detectors: different charged-current interactions lead to different event **morphologies** (there is some degeneracy)

Can we learn the flavour composition at the source to understand the production of astrophysical neutrinos?



Stan Yen

Oscillation

Flavour eigenstates ($\alpha = e, \mu, \tau$) are not eigenstates of the Hamiltonian ($i = 1,2,3$)

$$
U = \begin{pmatrix} c_{12}c_{13} & s_{12}c_{13} & s_{13}e^{-i\delta_{\rm CP}} \\ -s_{12}c_{23} - c_{12}s_{13}s_{23}e^{i\delta_{\rm CP}} & c_{12}c_{23} - s_{12}s_{13}s_{23}e^{i\delta_{\rm CP}} & c_{13}s_{23} \\ s_{12}s_{23} - c_{12}s_{13}c_{23}e^{i\delta_{\rm CP}} & -c_{12}s_{23} - s_{12}s_{13}c_{23}e^{i\delta_{\rm CP}} & c_{13}c_{23} \end{pmatrix}
$$

$c_{ij} \equiv \cos\theta_{ij}$

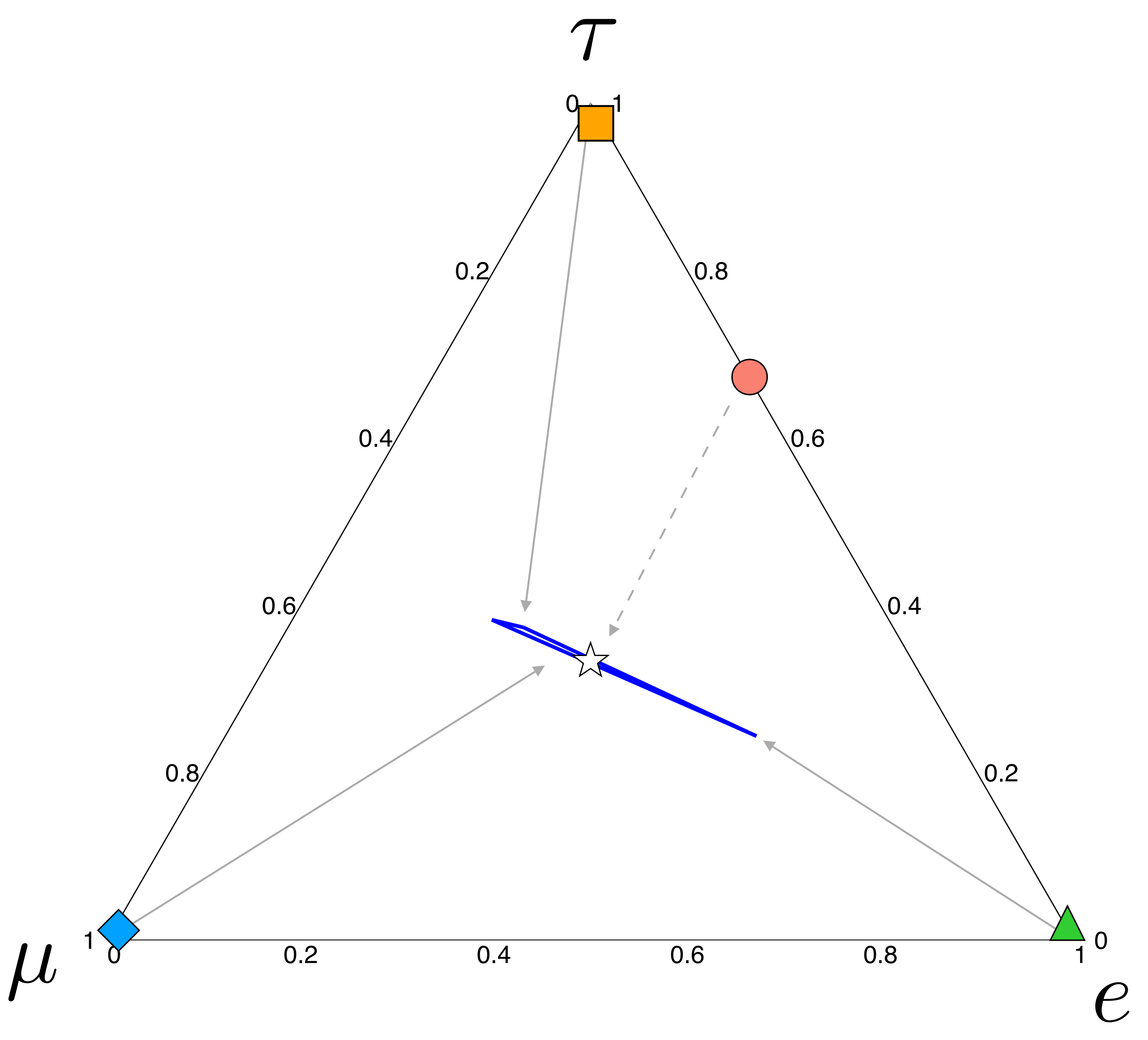
$s_{ij} \equiv \sin\theta_{ij}$

$$
|\nu_\alpha\rangle = \sum_{i=1}^{3} U^*_{\alpha i}|\nu_i\rangle,
$$

Flavour basis

PMNS mixing matrix

mass basis



Distances are **large and uncorrelated** -> mixing **averages out:**

$$
P_{\alpha\to\beta} = \sum_{i=1}^{3} |U_{\alpha i}|^2 |U_{\beta i}|^2
$$

$$
f_{\beta,\oplus} = \sum_{\alpha=e,\mu,\tau} P_{\alpha\beta} f_{\alpha,S}
$$

flavour composition at Earth

flavour composition at source

Flavour composition at Earth

$$U = \begin{pmatrix} c_{12}c_{13} & s_{12}c_{13} & s_{13}e^{-i\delta_{\rm CP}} \\ -s_{12}c_{23} - c_{12}s_{13}s_{23}e^{i\delta_{\rm CP}} & c_{12}c_{23} - s_{12}s_{13}s_{23}e^{i\delta_{\rm CP}} & c_{13}s_{23} \\ s_{12}s_{23} - c_{12}s_{13}c_{23}e^{i\delta_{\rm CP}} & -c_{12}s_{23} - s_{12}s_{13}c_{23}e^{i\delta_{\rm CP}} & c_{13}c_{23} \end{pmatrix}$$

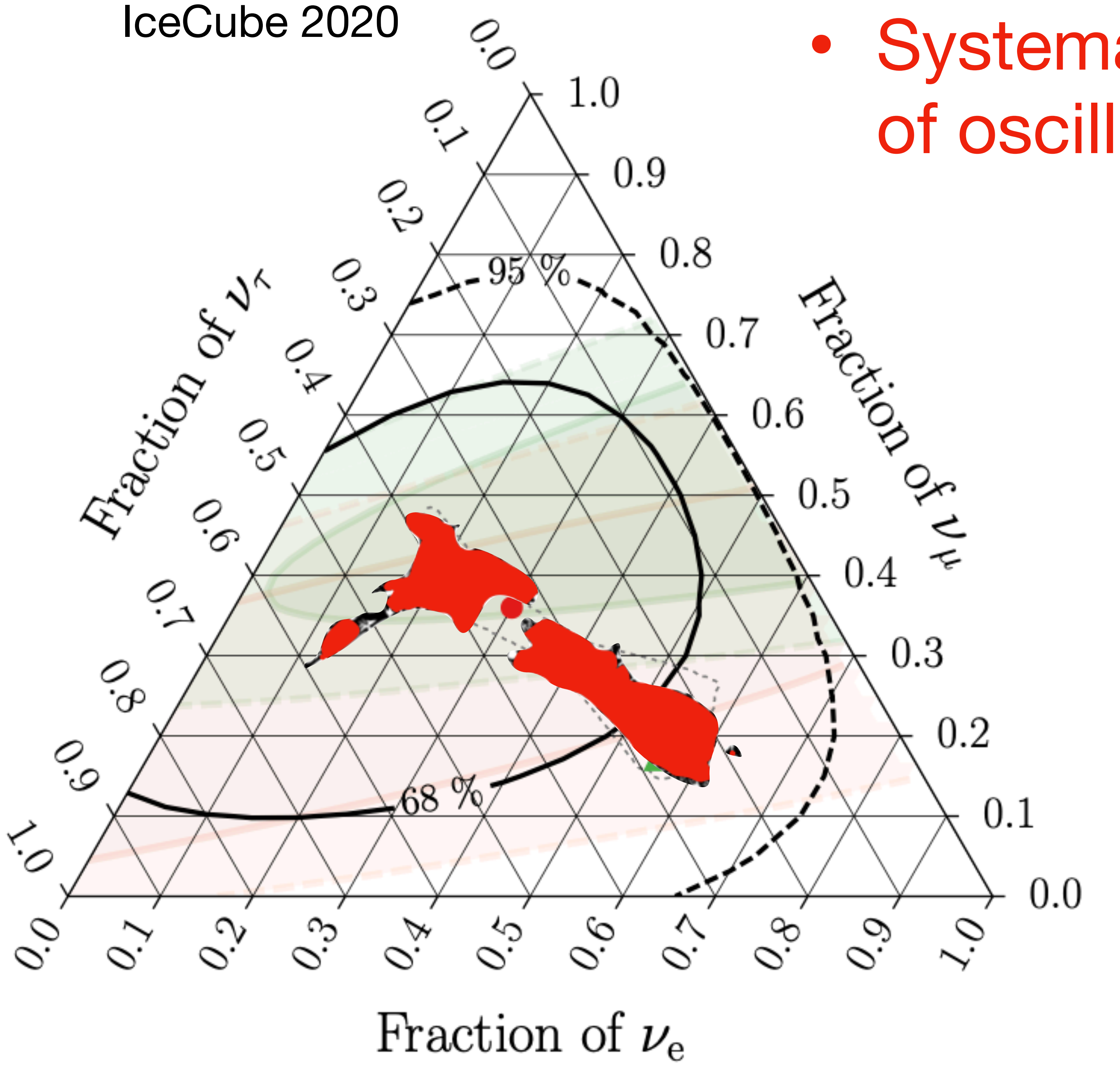
NuFit 5.0 global fit

Parameter	Normal ordering	Inverted ordering
$\sin^2\theta_{12}$	$0.304^{+0.012}_{-0.012}$	$0.304^{+0.013}_{-0.012}$
$\sin^2\theta_{23}$	$0.573^{+0.016}_{-0.020}$	$0.575^{+0.016}_{-0.019}$
$\sin^2\theta_{13}$	$0.02219^{+0.00062}_{-0.00063}$	$0.02238^{+0.00063}_{-0.00062}$
$\delta_{\rm CP}$ (°)	197^{+27}_{-24}	282^{+26}_{-30}

Two limits:

- Statistics (astrophysical neutrinos)
- Systematics: precise knowledge of oscillation parameters

IceCube 2020



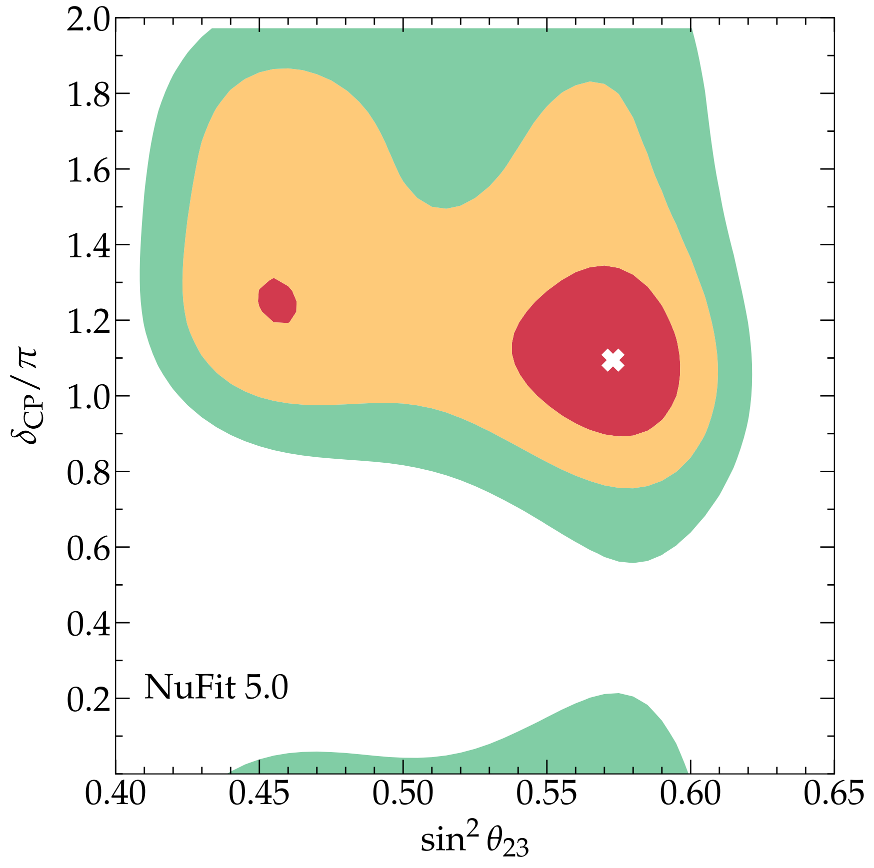
θ_{12} ("solar angle"): Solar, reactor experiments

θ_{23} ("atmospheric angle") Atmospheric, long-baseline

θ_{13} Reactor experiments

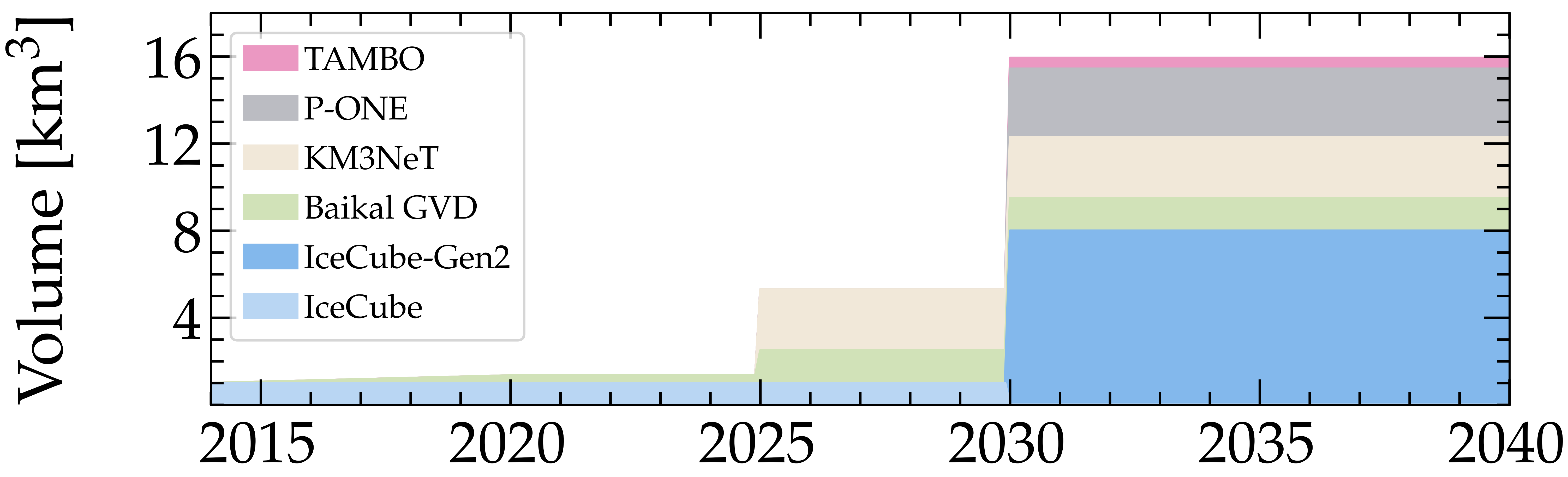
δ_{CP} Long-baseline experiments

Mostly uncorrelated except:



What does the future say about this?

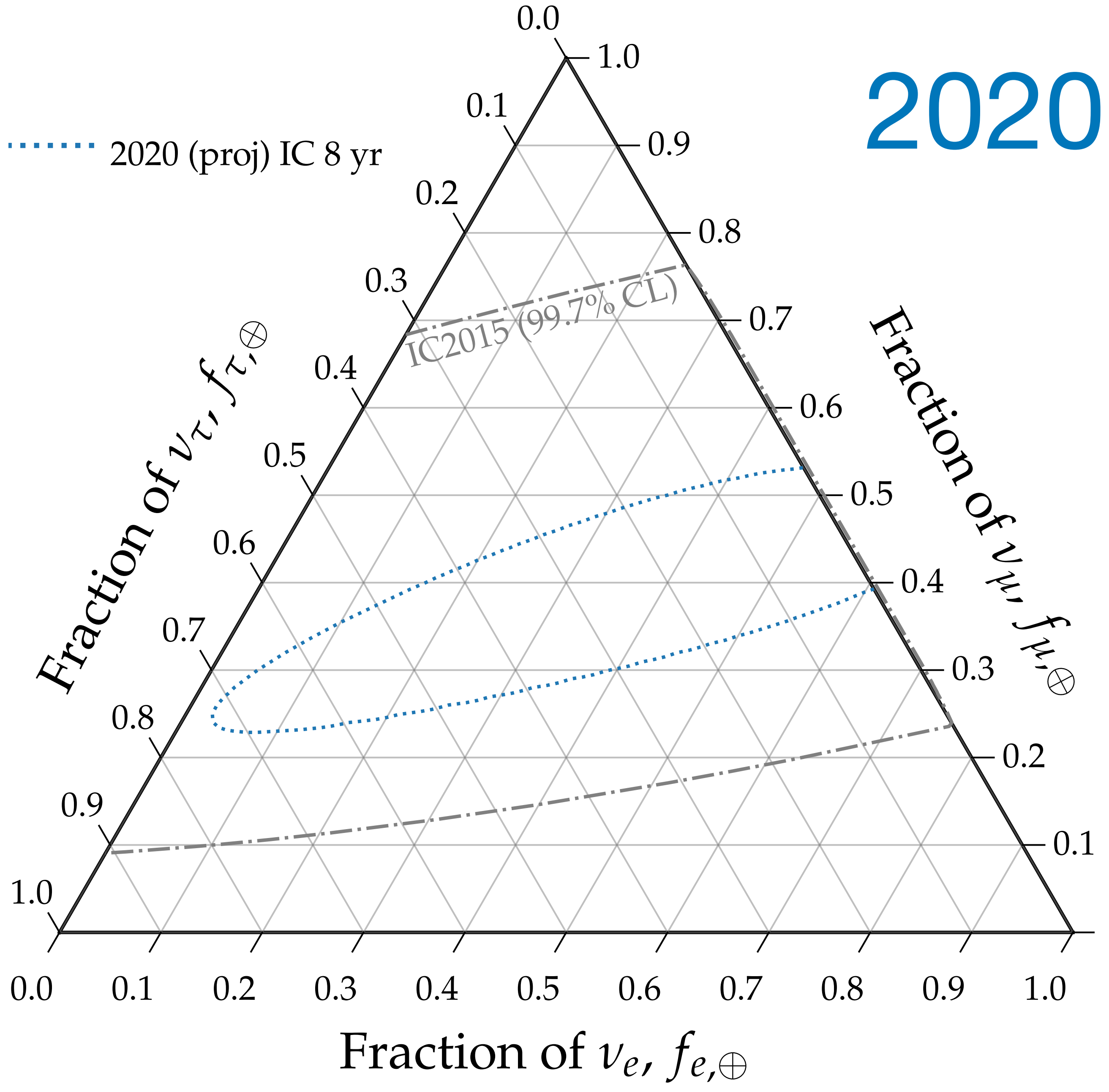
Statistics: need more Cherenkov telescopes!



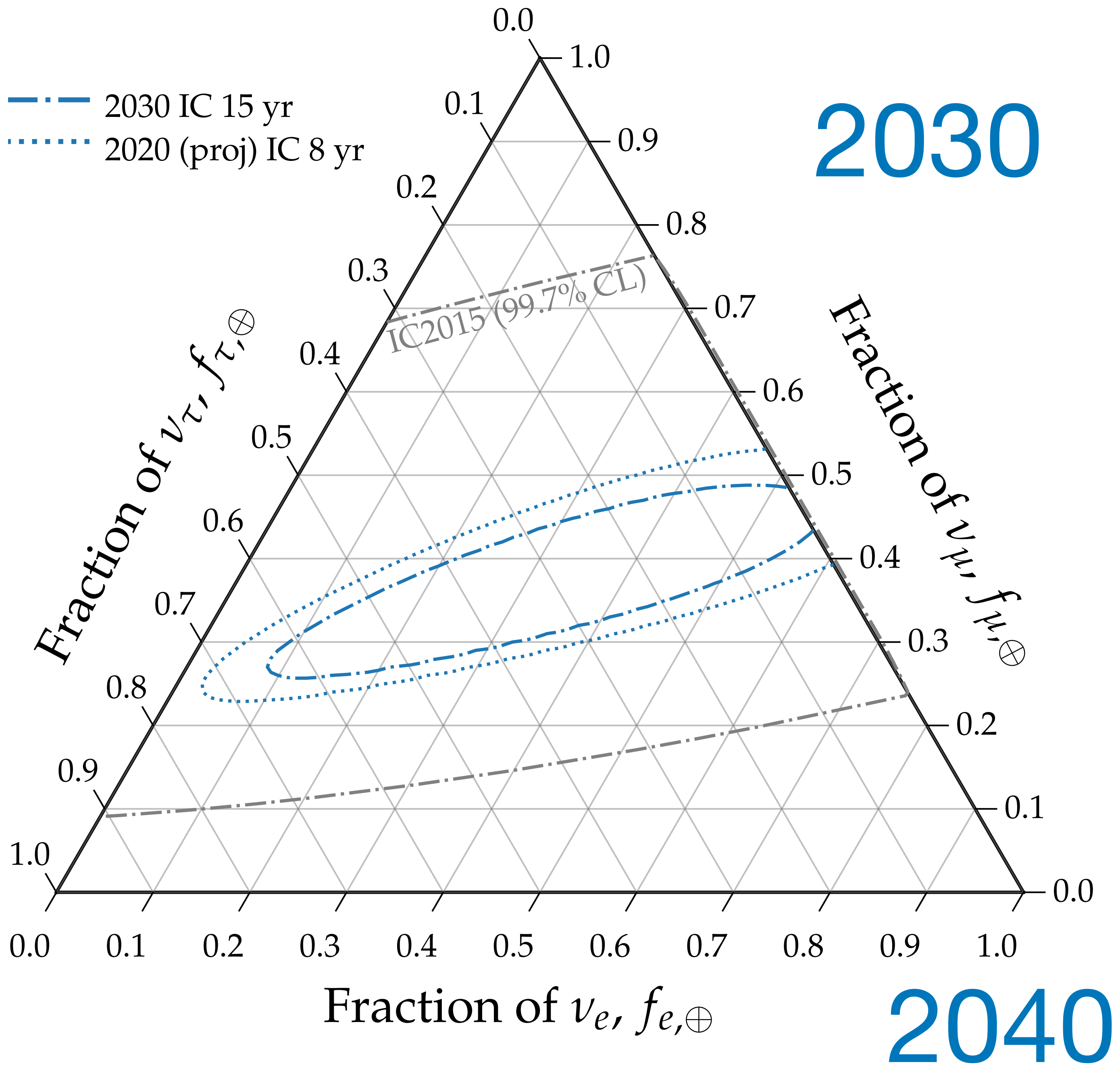
Telescope	Medium	Location	Exposure (km^3)
IceCube-Gen2	Ice	South pole (HE upgrade of IceCube)	~6-9
KM3NeT	Seawater	Mediterranean Sea (successor to ANTARES)	~2-3
GVD	Freshwater	Lake Baikal	1.5
P-ONE	Seawater	Cascadia Basin (Pacific Ocean)	π
TAMBO	Rock/air/water Cherenkov	Peru	~10 (very high E, tau only)

Statistics

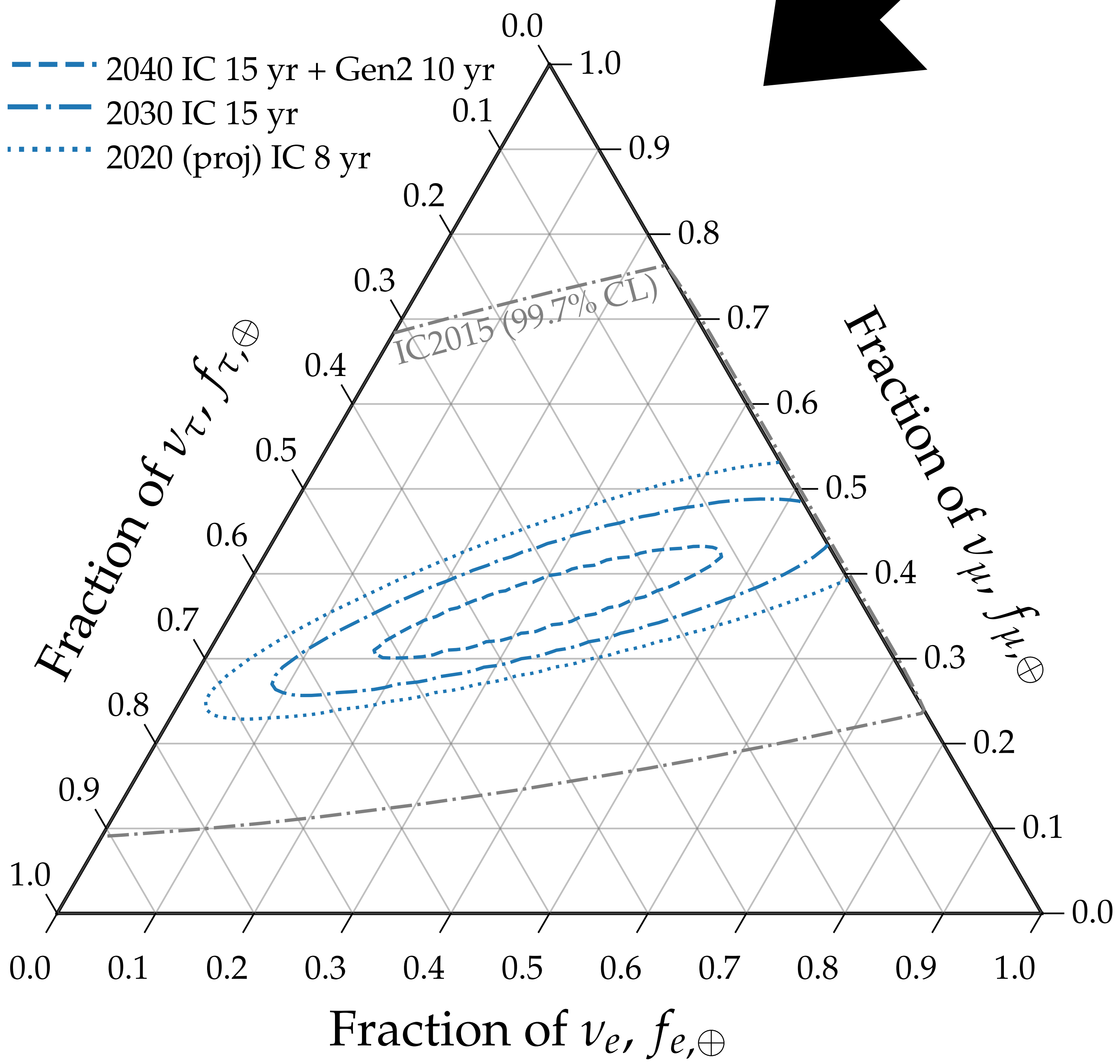
2020



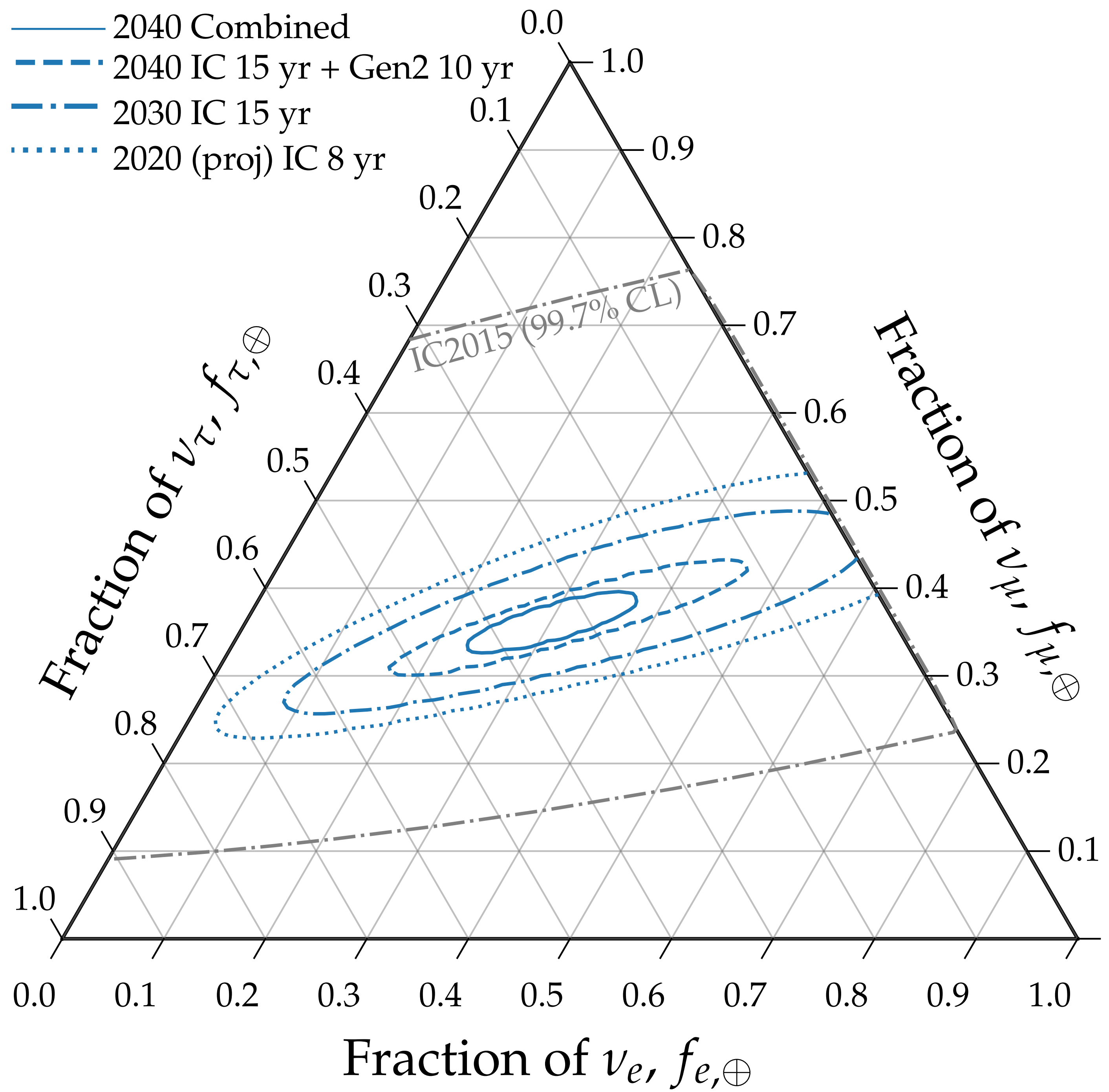
2030



2040 (IceCube-Gen2)

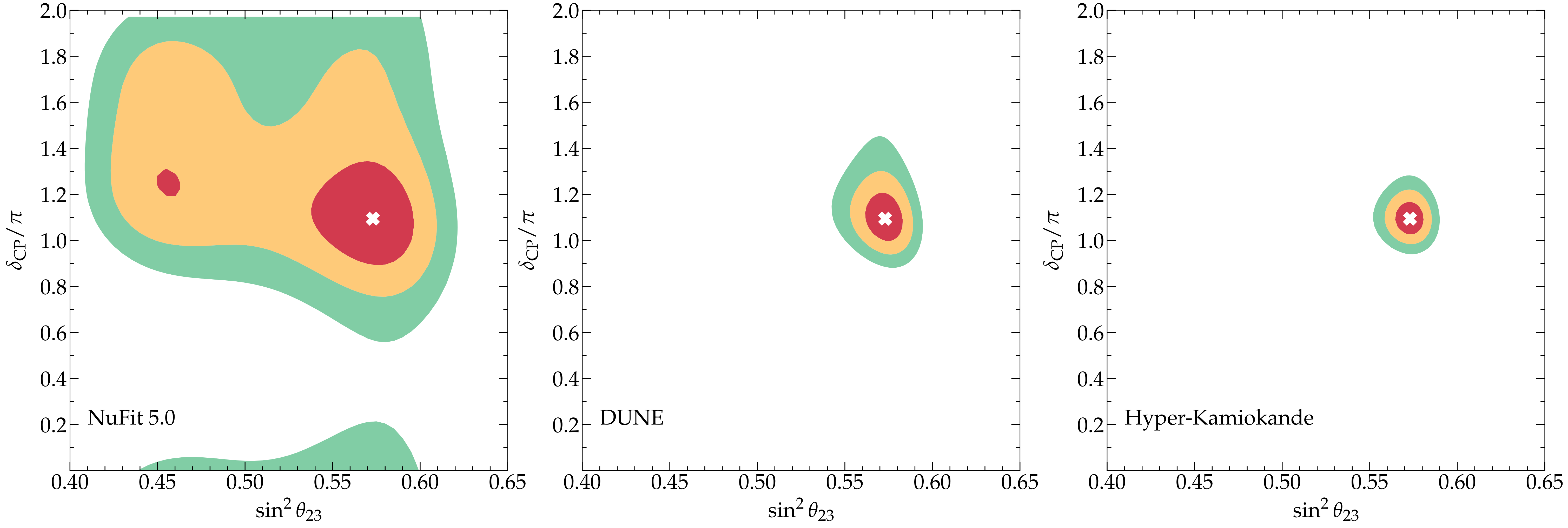


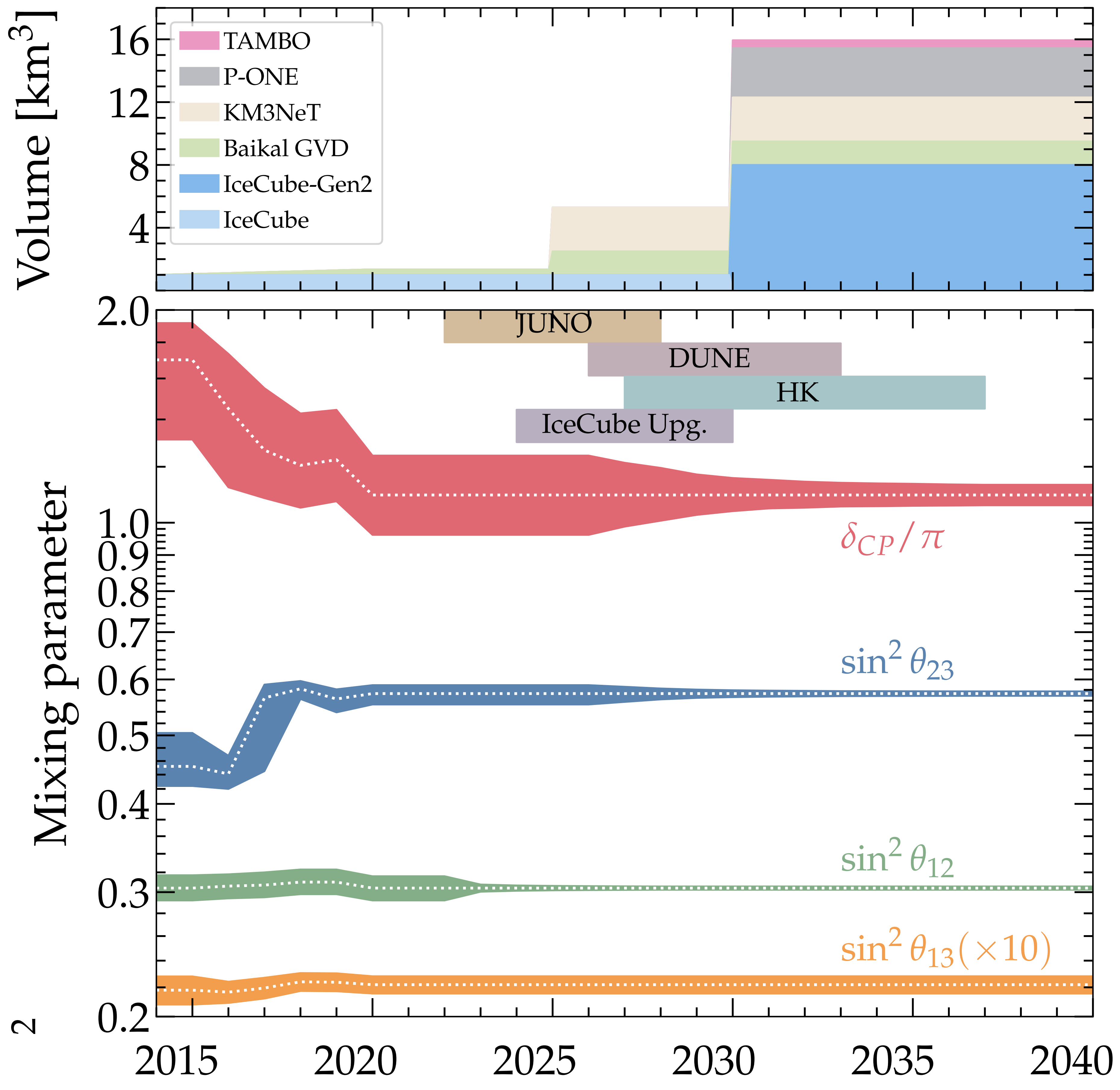
2040(all experiments)



Systematics: terrestrial experiments

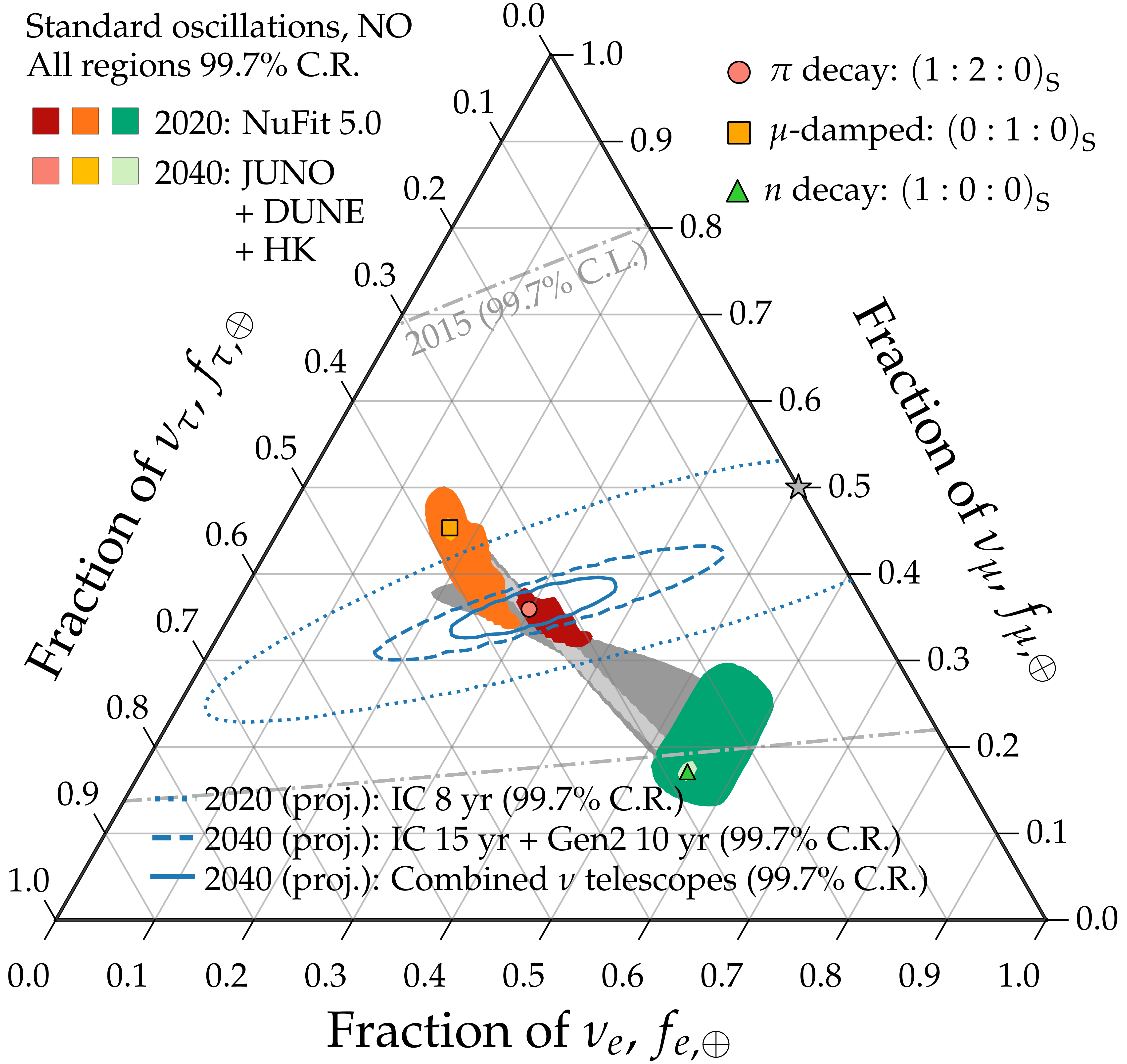
- JUNO: 2022-2028: 20kt liquid scintillator reactor measurement. 0.52% uncertainty on $\sin^2\theta_{12}$
- DUNE: ~2026-2033: 40kt liquid argon long baseline experiment. θ_{23} & δ_{CP}
- Hyper-Kamiokande: 187 kt water Cherenkov.θ_{23} & δ_{CP}
- IceCube Upgrade: dense instrumentation: constrain unitarity



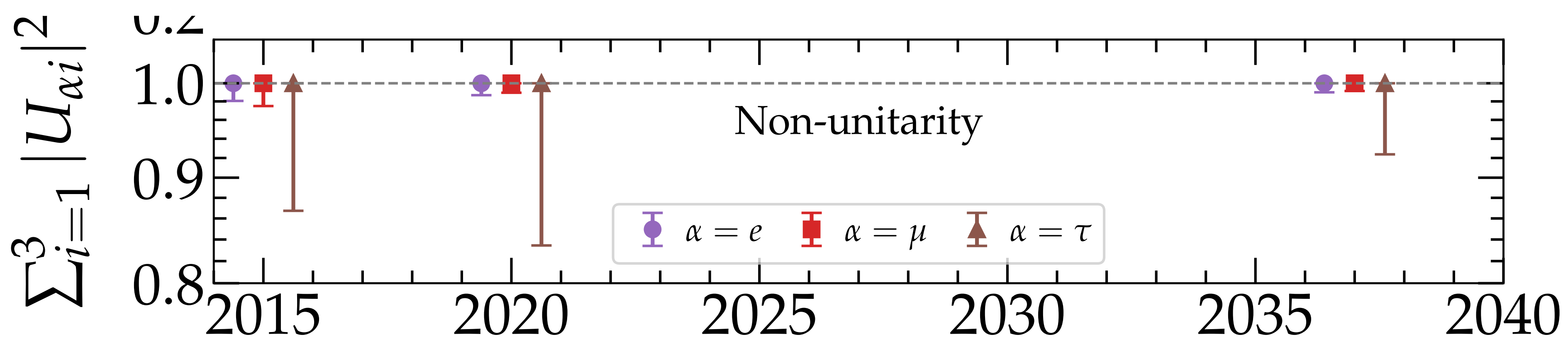


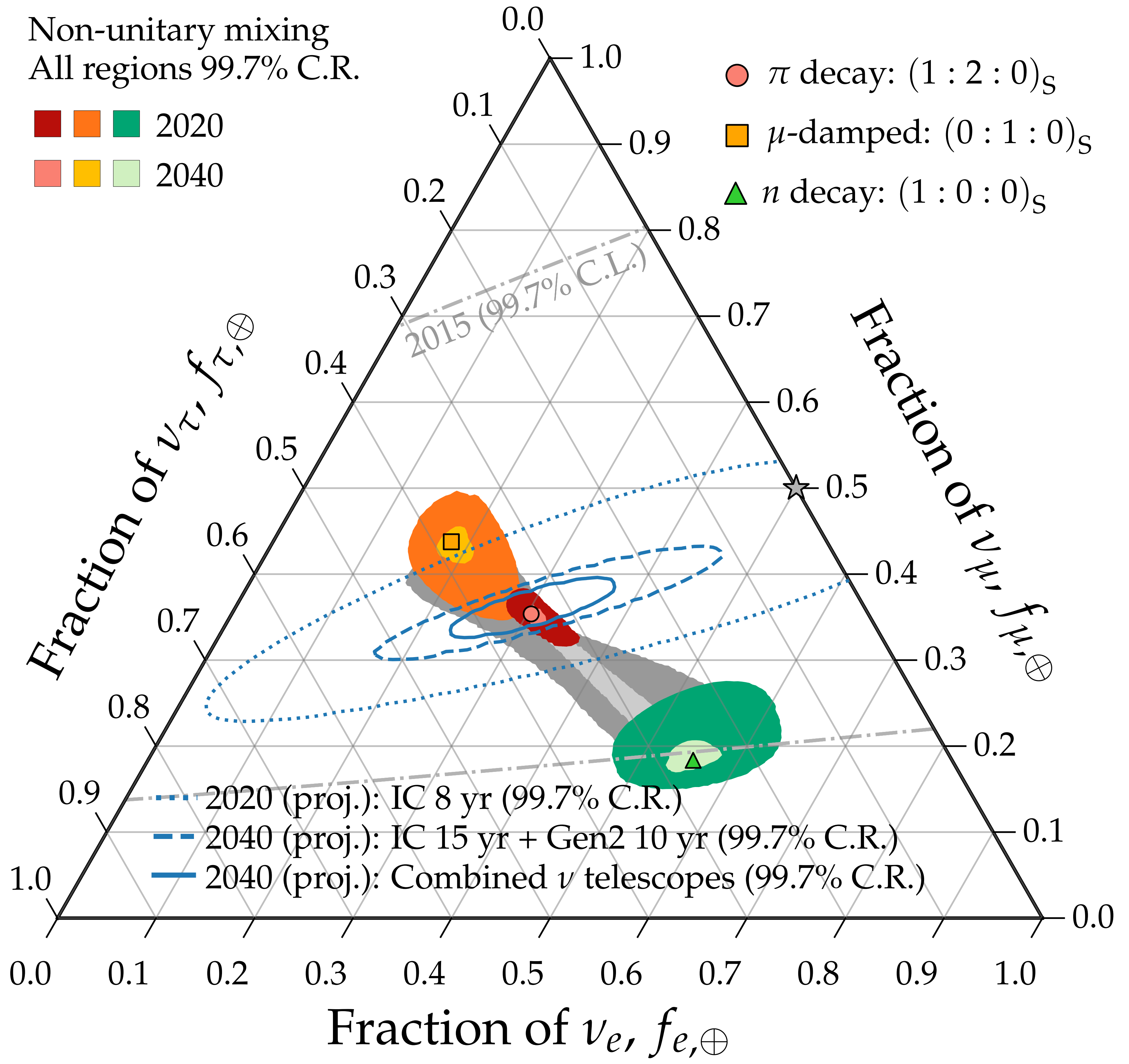
$$|\nu_\alpha\rangle = \frac{1}{\sqrt{N_\alpha}} \sum_{i=1}^{3} U^*_{\alpha i} |\nu_i\rangle$$

Without assuming unitary 3x3 PMNS matrix?

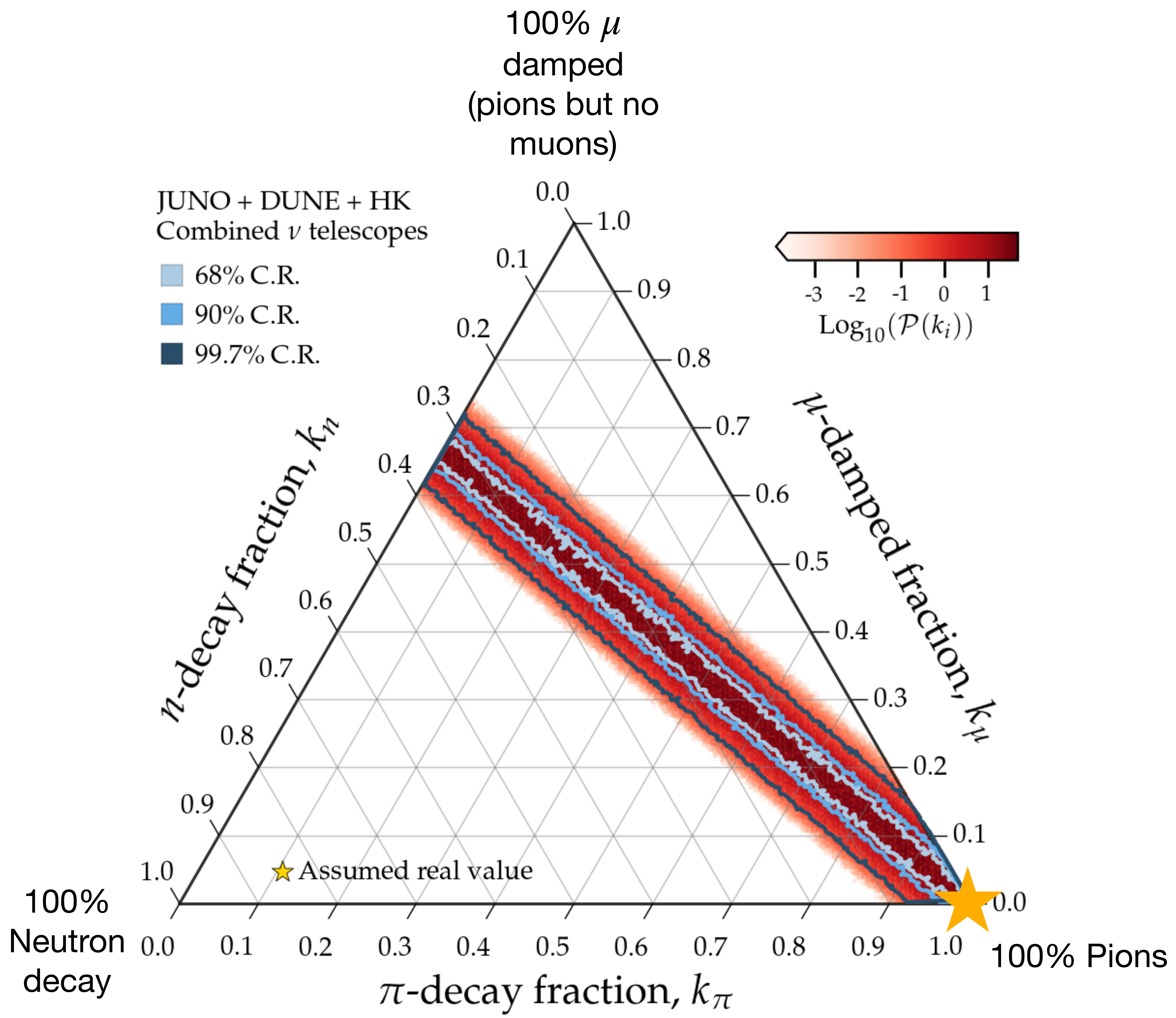


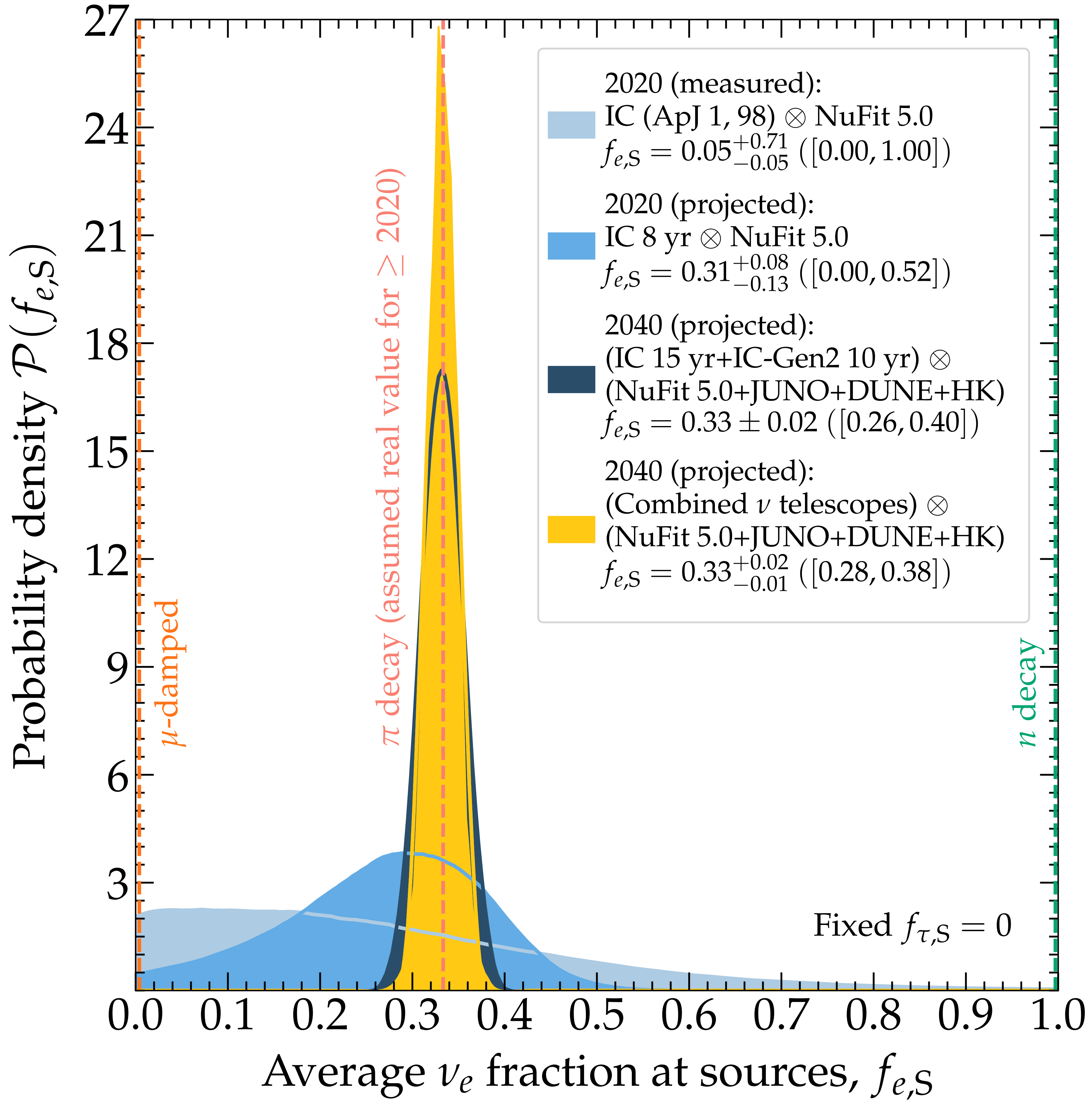
Grey: allowed region within the standard model
Colours: allowed region assuming a single source composition





Flavour composition at the source?

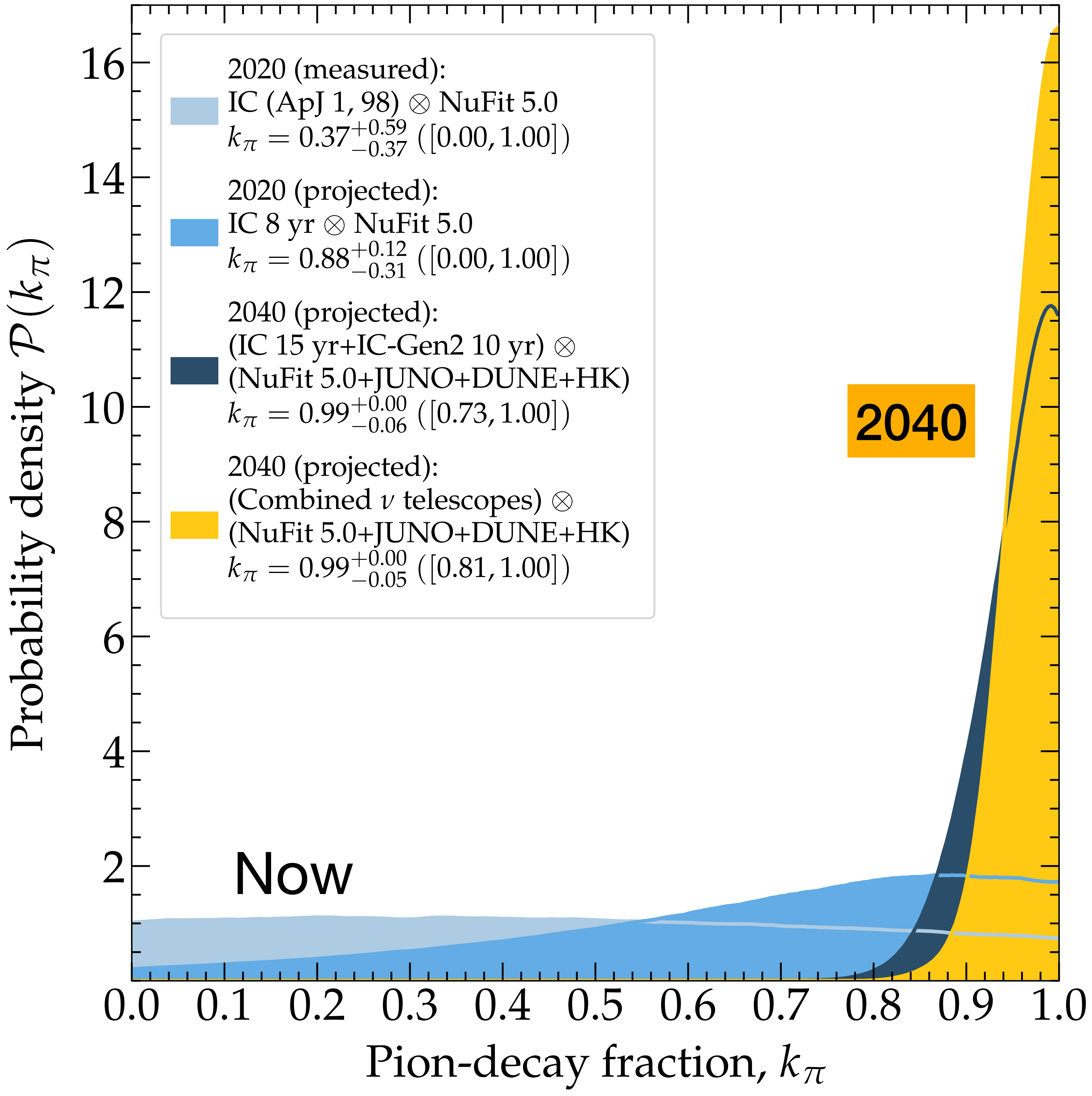




Flavour composition at the source

Dominant production mechanism can be pinned down to within 20% *using neutrino flavour alone*.

Assuming no neutron decay



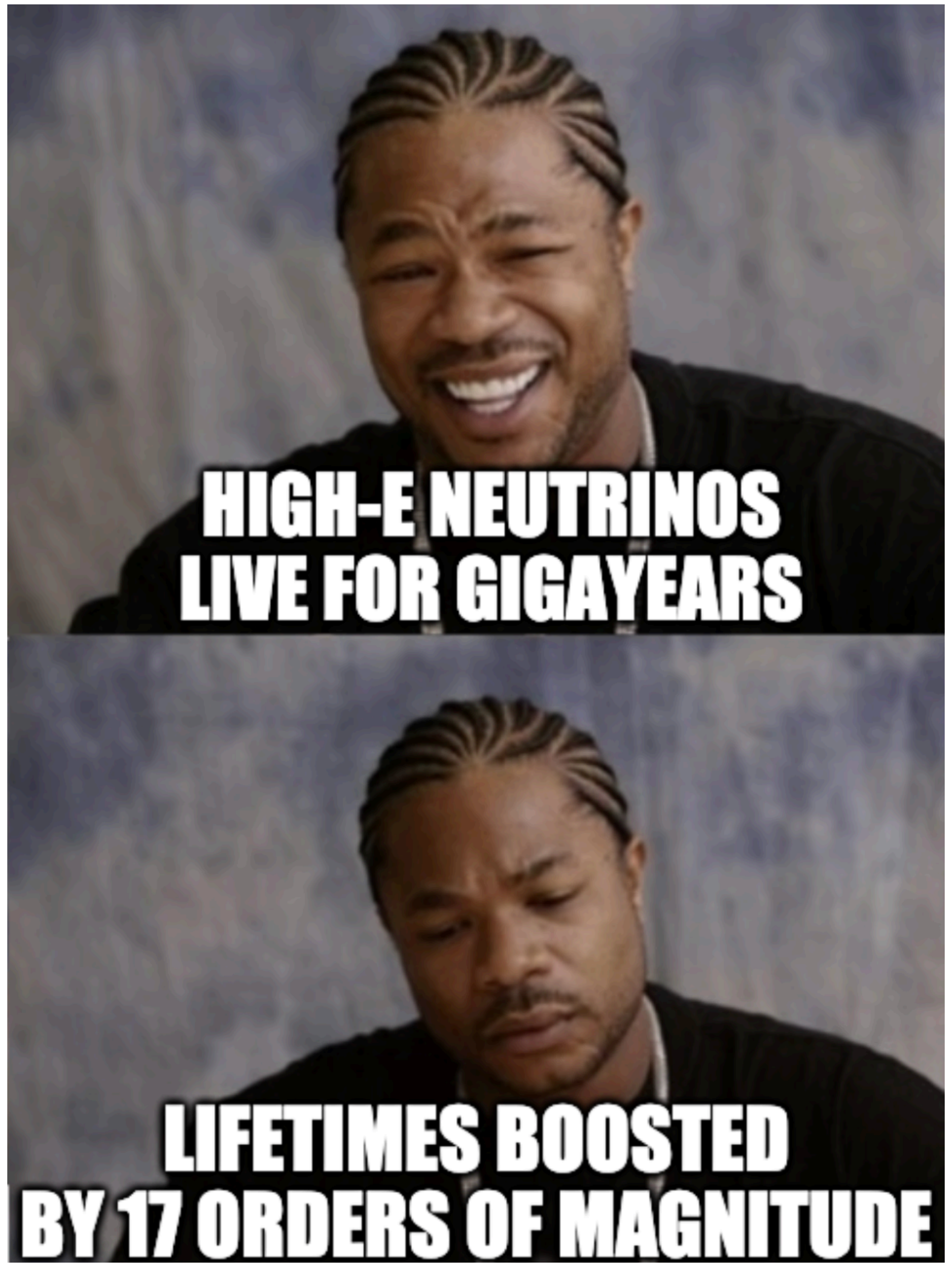
New physics: neutrino decay

Neutrino decay

Invisible decay: all but one mass eigenstate decays to invisible species.

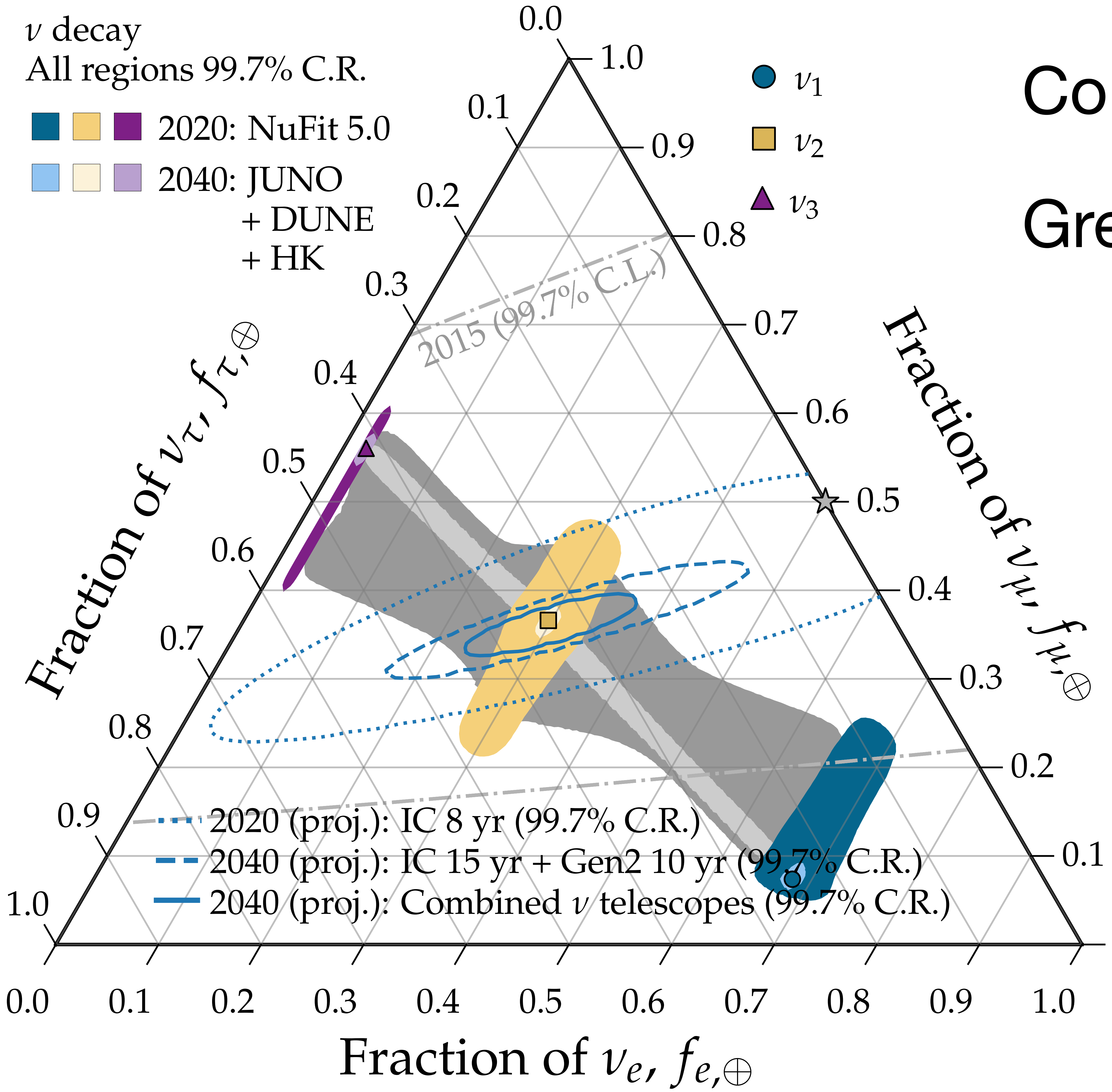
$$N_\nu = N(z_0)\exp\left\{ -\frac{m_\nu}{\tau E_\nu} \int_0^{z_0} \frac{dz}{(1+z)^2 H_0 \sqrt{\Omega(z)}} \right\}$$

neutrino lifetime at rest



Must be integrated over distribution of cosmic sources

See Abdullah & Denton 2005.07200 for a complete treatment of *visible* decay

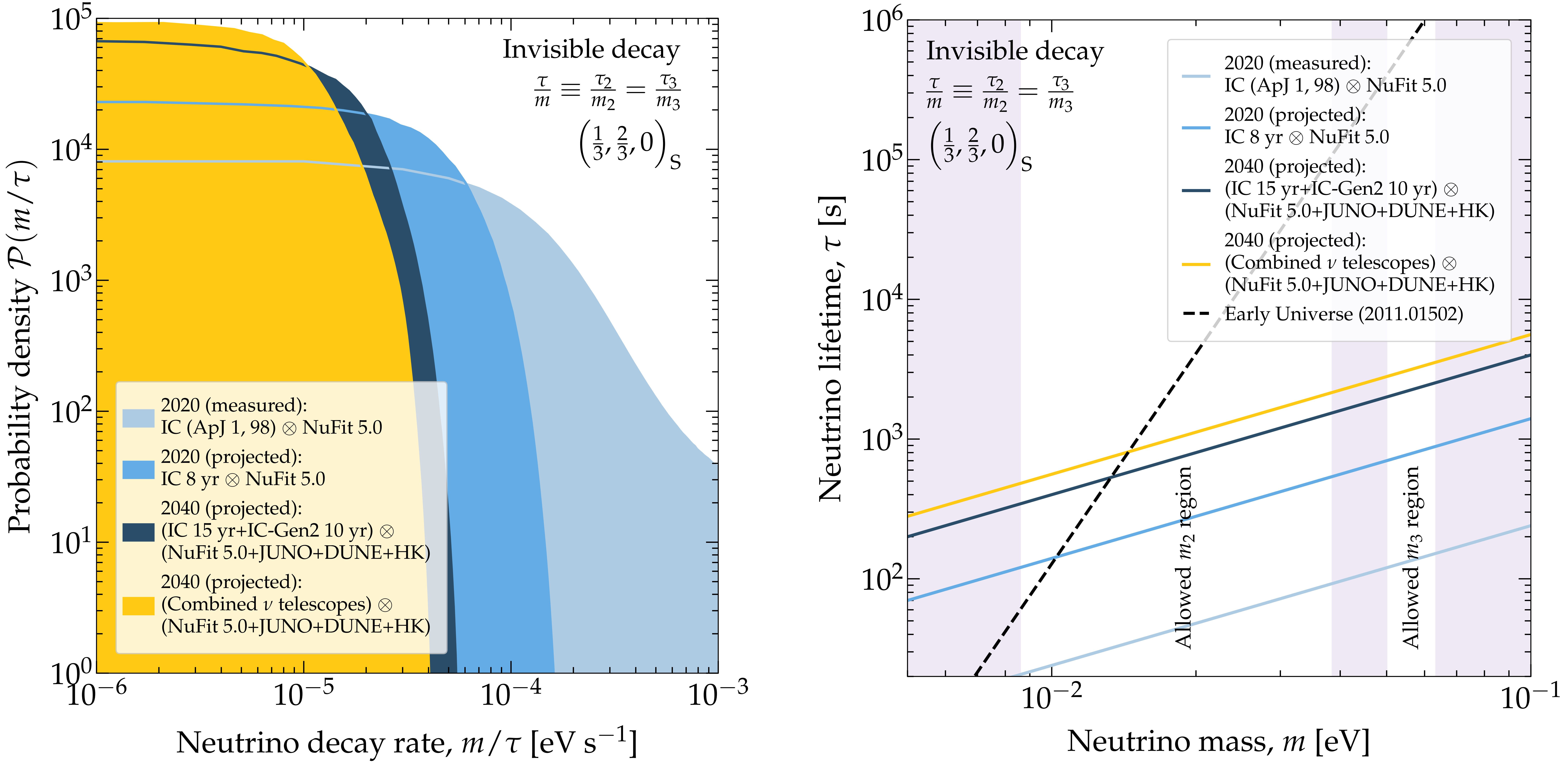


Colour: full decay

Grey: partial decay

Full decay of m_2 and m_3 almost excluded now

Sensitivity to single mass eigenstates



Invisible decay
$\frac{\tau}{m} \equiv \frac{\tau_2}{m_2} = \frac{\tau_3}{m_3}$
$\left(\frac{1}{3}, \frac{2}{3}, 0\right)_S$
Probability density $\mathcal{P}(m/\tau)$
Neutrino decay rate, m/τ [eV s^{-1}]
2020 (measured):
IC (ApJ 1, 98) ⊗ NuFit 5.0
2020 (projected):
IC 8 yr ⊗ NuFit 5.0
2040 (projected):
(IC 15 yr+IC-Gen2 10 yr) ⊗
(NuFit 5.0+JUNO+DUNE+HK)
2040 (projected):
(Combined ν telescopes) ⊗
(NuFit 5.0+JUNO+DUNE+HK)
Early Universe (2011.01502)
Neutrino lifetime, τ [s]
Neutrino mass, m [eV]
Allowed m_2 region
Allowed m_3 region

Dark matter

Dark Matter Annihilation to Neutrinos

Carlos A. Argüelles, Alejandro Diaz, Ali Kheirandish, Andrés Olivares-Del-Campo, Ibrahim Safa, Aaron C. Vincent

[Accepted/Reviews in Modern Physics] https://arxiv.org/abs/1912.09486

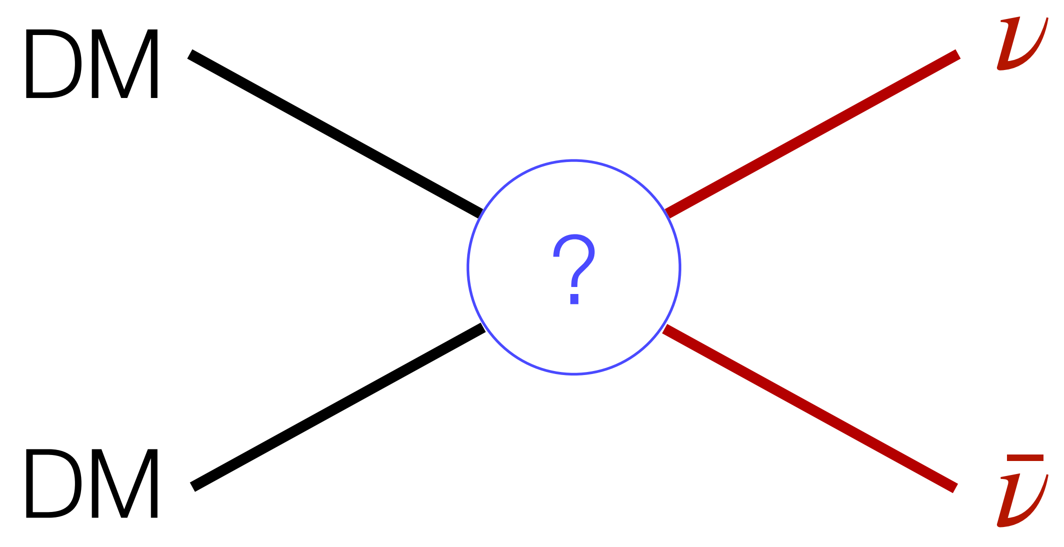
What is the sensitivity of neutrino detectors to new physics?

Illustrate with DM annihilation to neutrinos

Indirect searches $\chi\chi \to SM, SM$: gammas dominate, *except* if neutrinos are the only product

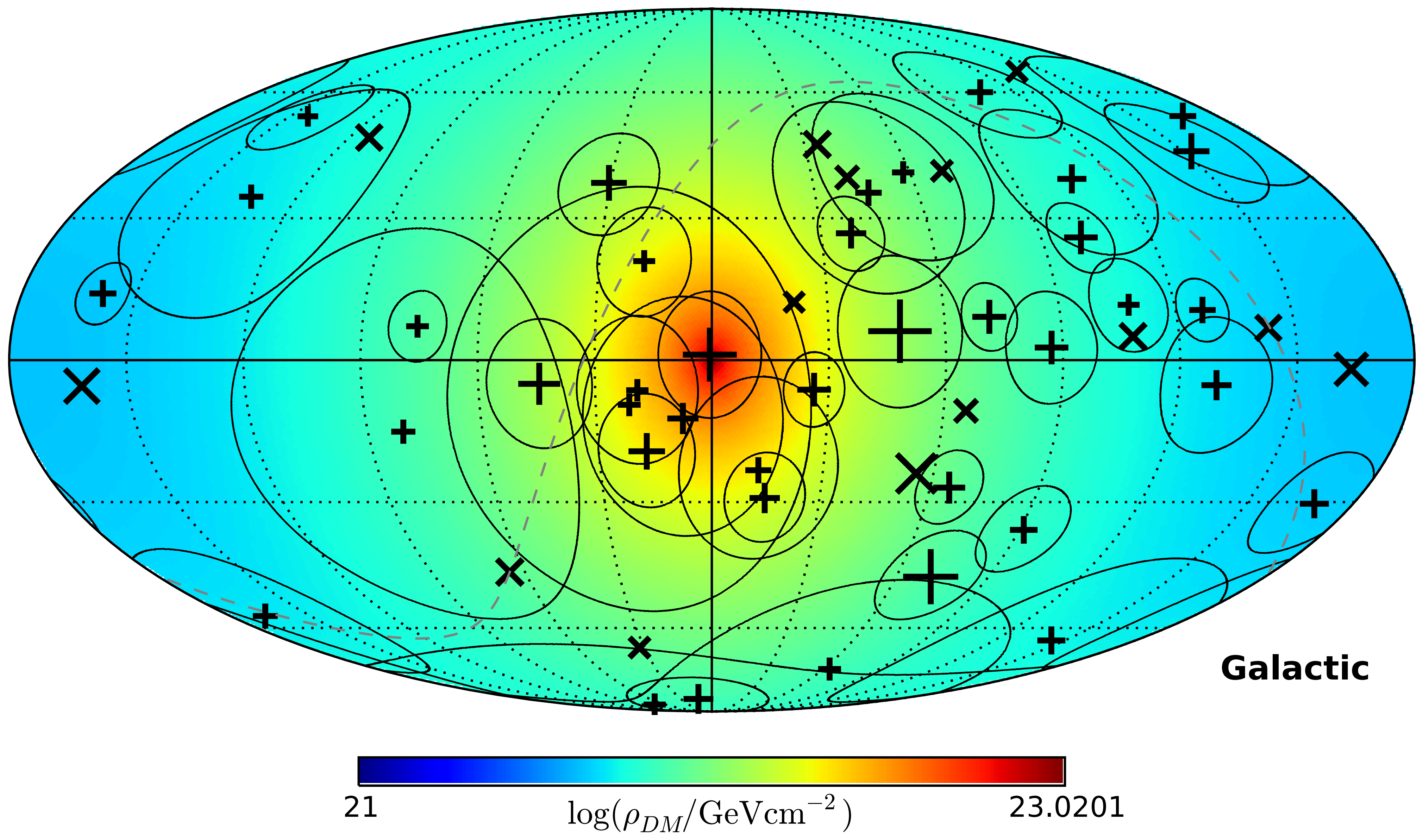
$$\frac{d\Phi_{\nu+\bar{\nu}}}{dE_\nu} = \frac{1}{4\pi}\frac{\langle\sigma v\rangle}{\kappa m_\chi^2}\frac{1}{3}\frac{dN_\nu}{dE_\nu}J(\Omega)$$

$$J \equiv \int_{\Delta\Omega} d\Omega \int_{\mathrm{l.o.s.}} \rho_\chi^2(x)dx$$



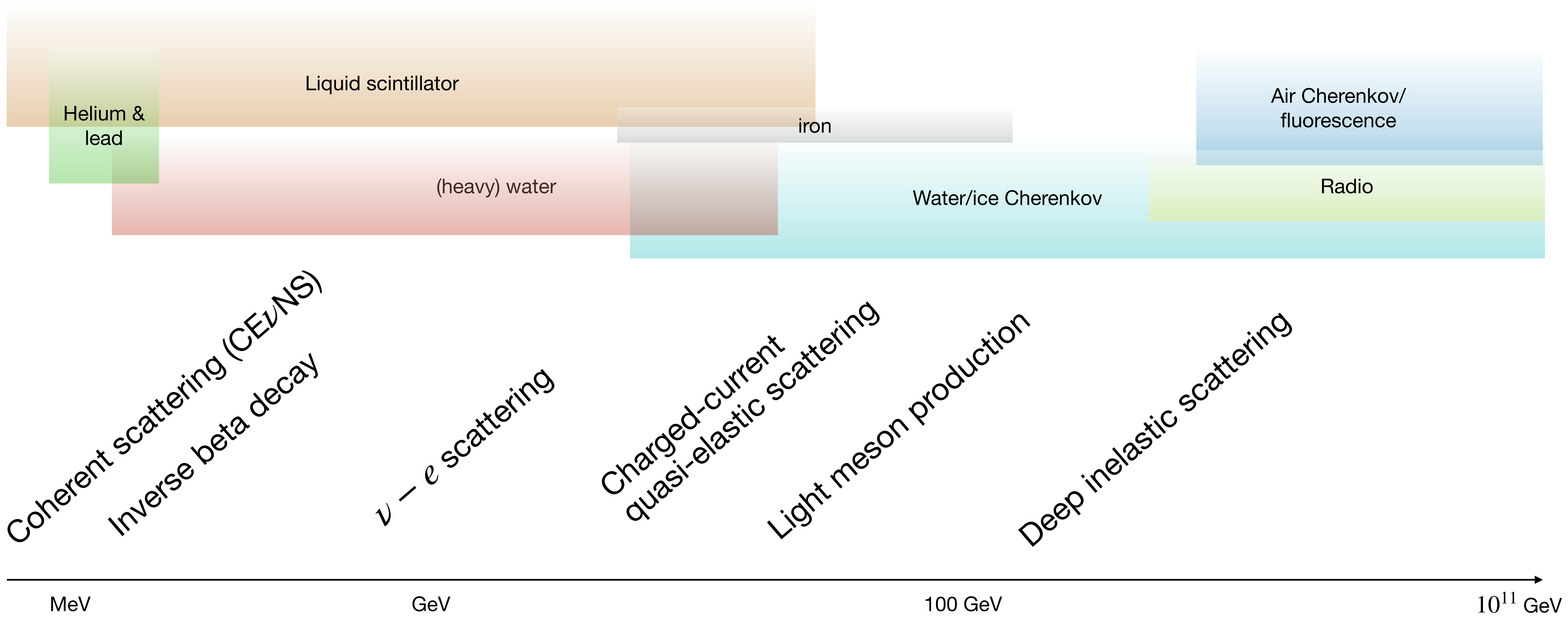
$$\frac{dN_\nu}{dE_\nu} = 2\delta(1 - E/m_\chi)m_\chi/E^2 .$$

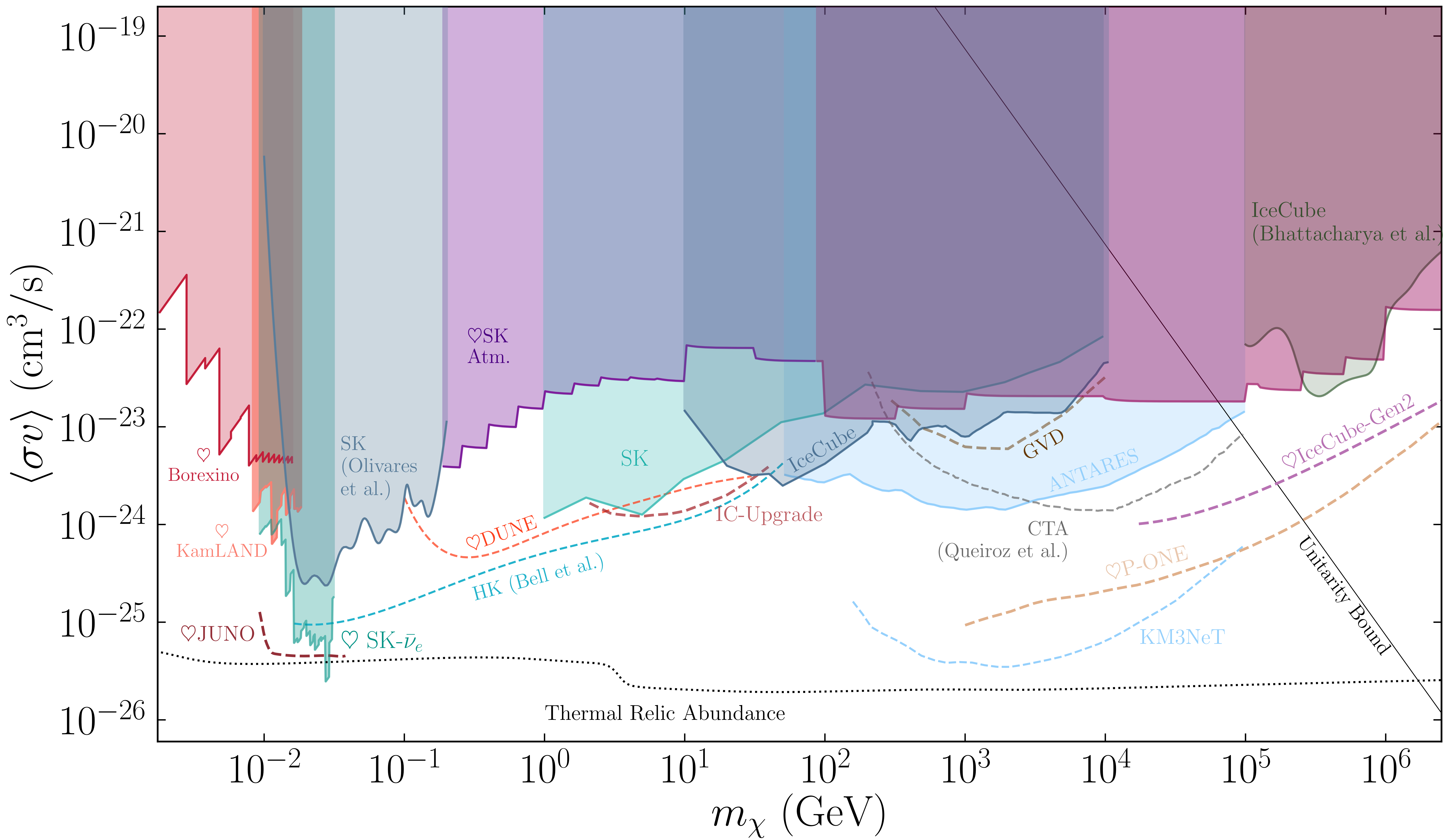
Dark matter column density



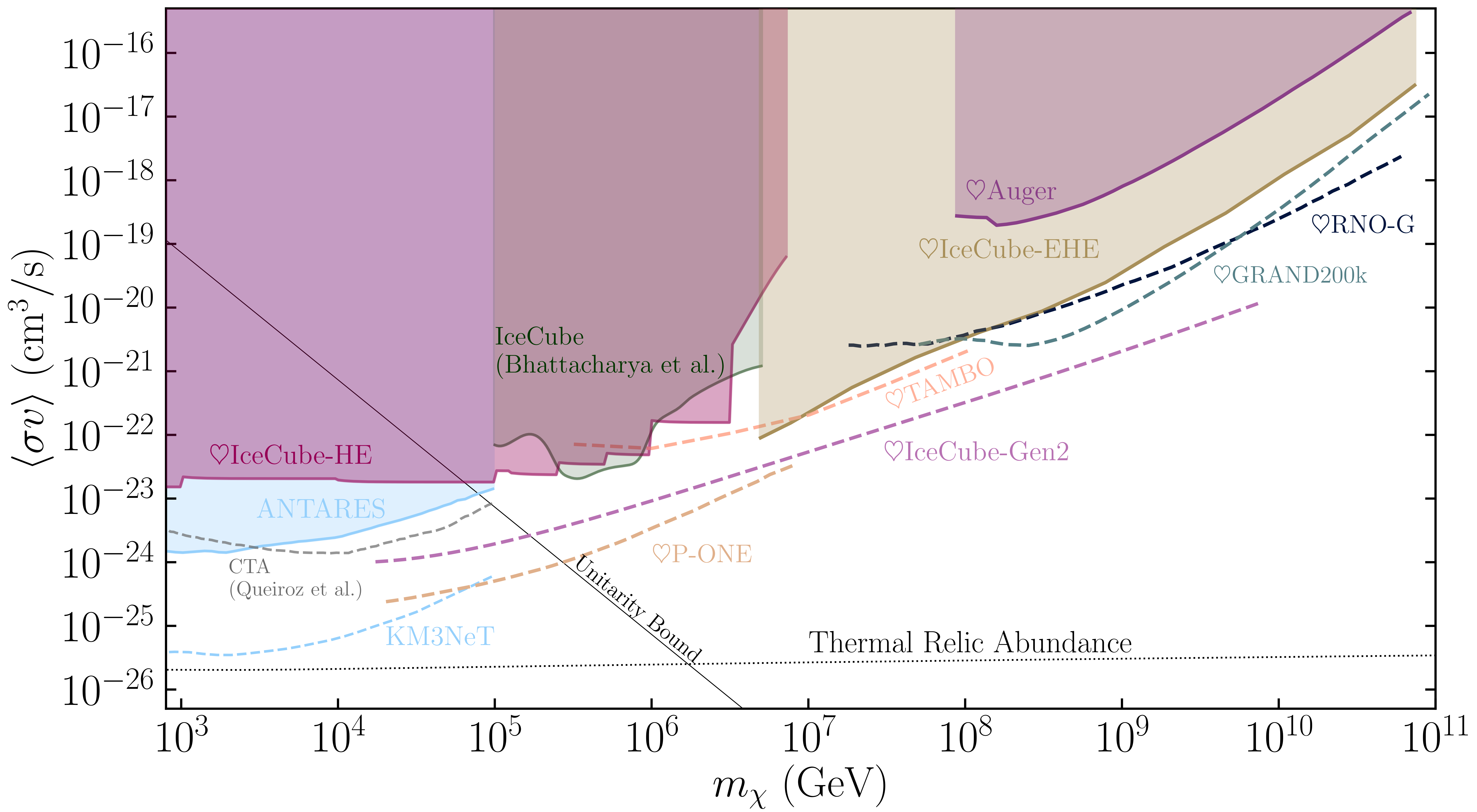
*We also spent a long time calculating extragalactic constraints. They are subdominant though

What if we looked at *every neutrino telescope in the world?*





10^{-19}
10^{-20}
10^{-21}
10^{-22}
10^{-23}
10^{-24}
10^{-25}
10^{-26}
$\langle \sigma v \rangle$ (cm^3/s)
10^{-2}
10^{-1}
10^{0}
10^{1}
10^{2}
10^{3}
10^{4}
10^{5}
10^{6}
m_χ (GeV)
♡ Borexino
♡ KamLAND
♡JUNO
SK (Olivares et al.)
♡SK Atm.
SK
♡ SK-$\bar{\nu}_e$
♡DUNE
HK (Bell et al.)
IC-Upgrade
IceCube
GVD
ANTARES
CTA (Queiroz et al.)
♡P-ONE
KM3NeT
♡IceCube-Gen2
IceCube (Bhattacharya et al.)
Unitarity Bound
Thermal Relic Abundance



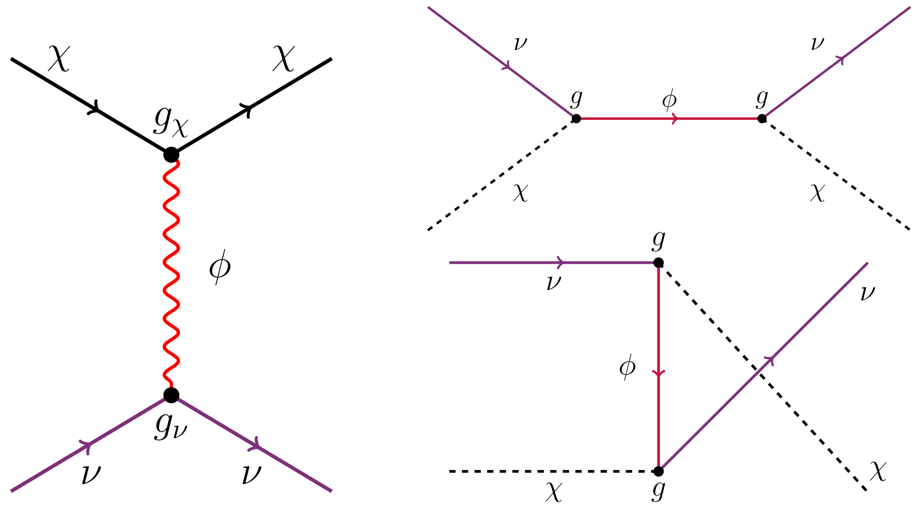
10^{-16}
10^{-17}
10^{-18}
10^{-19}
10^{-20}
10^{-21}
10^{-22}
10^{-23}
10^{-24}
10^{-25}
10^{-26}
$\langle \sigma v \rangle$ (cm^3/s)
10^3
10^4
10^5
10^6
10^7
10^8
10^9
10^{10}
10^{11}
m_χ (GeV)
♡Auger
♡IceCube-EHE
♡RNO-G
♡GRAND200k
IceCube
(Bhattacharya et al.)
♡TAMBO
♡IceCube-HE
♡IceCube-Gen2
ANTARES
CTA
(Queiroz et al.)
♡P-ONE
KM3NeT
Unitarity Bound
Thermal Relic Abundance

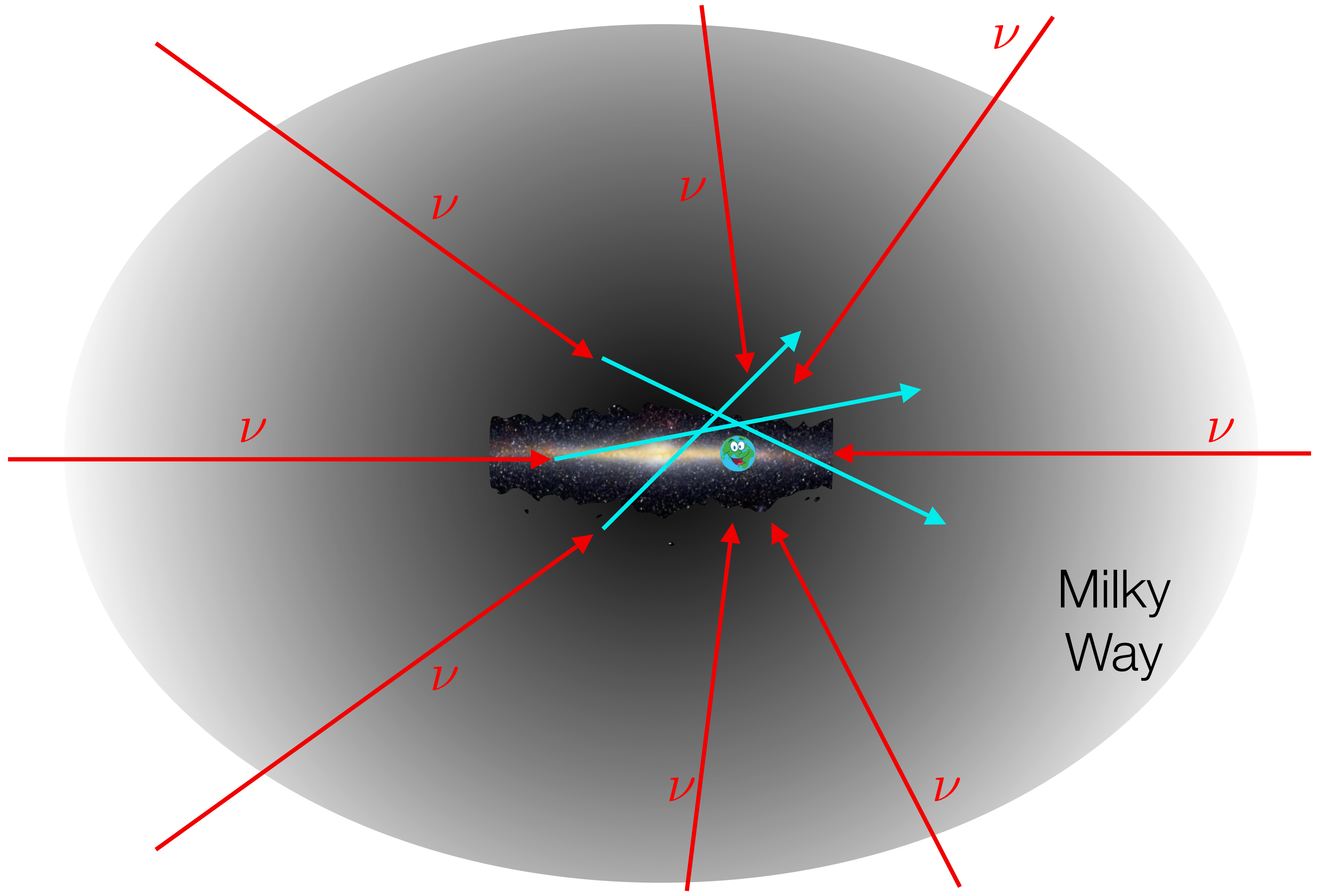
Neutrino-dark matter *elastic scattering*

Imaging Galactic Dark Matter with High-Energy Cosmic Neutrinos

Carlos A. Argüelles, Ali Kheirandish, Aaron C. Vincent

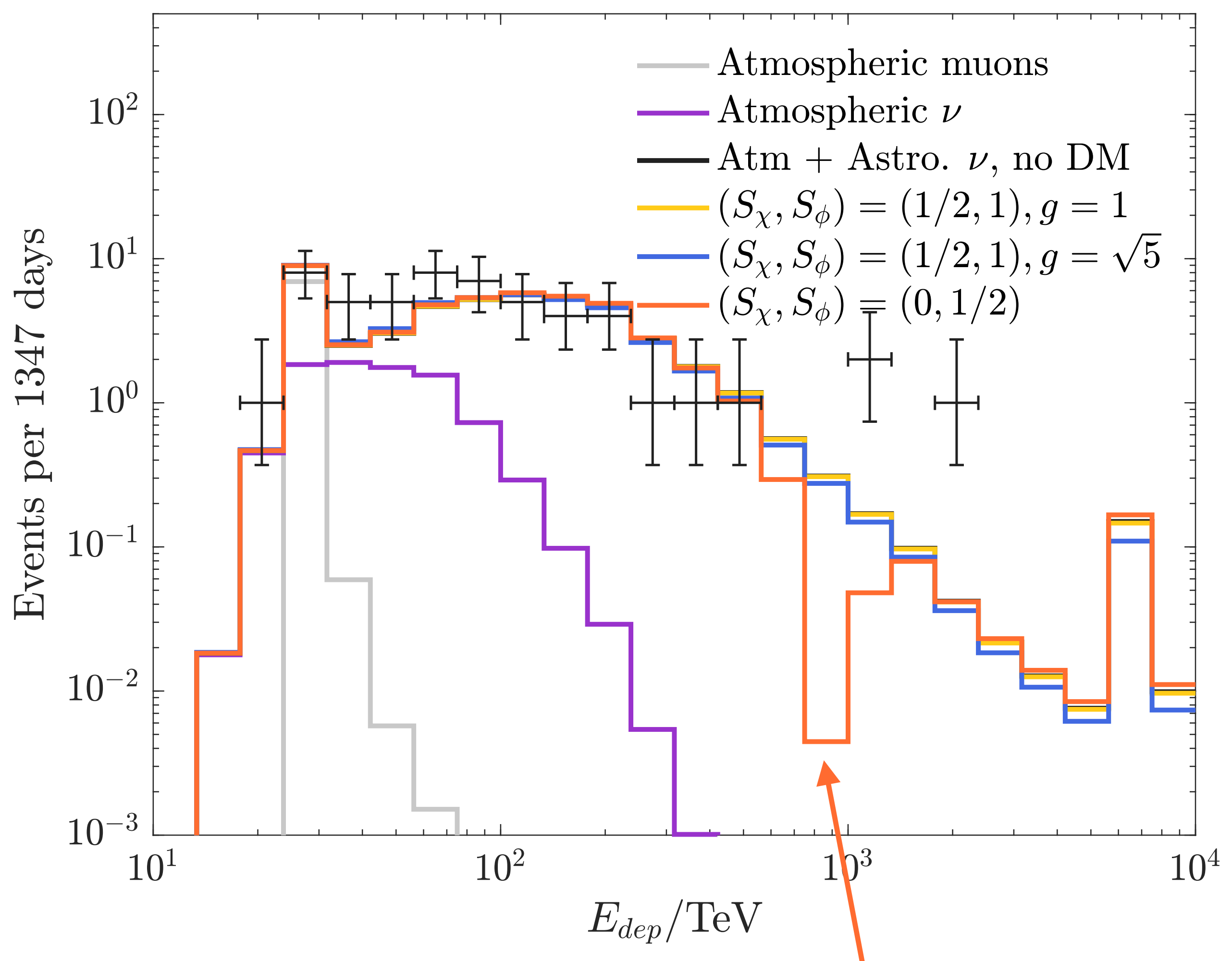
PRL https://arxiv.org/abs/1703.00451





Anisotropic deflection/ energy loss

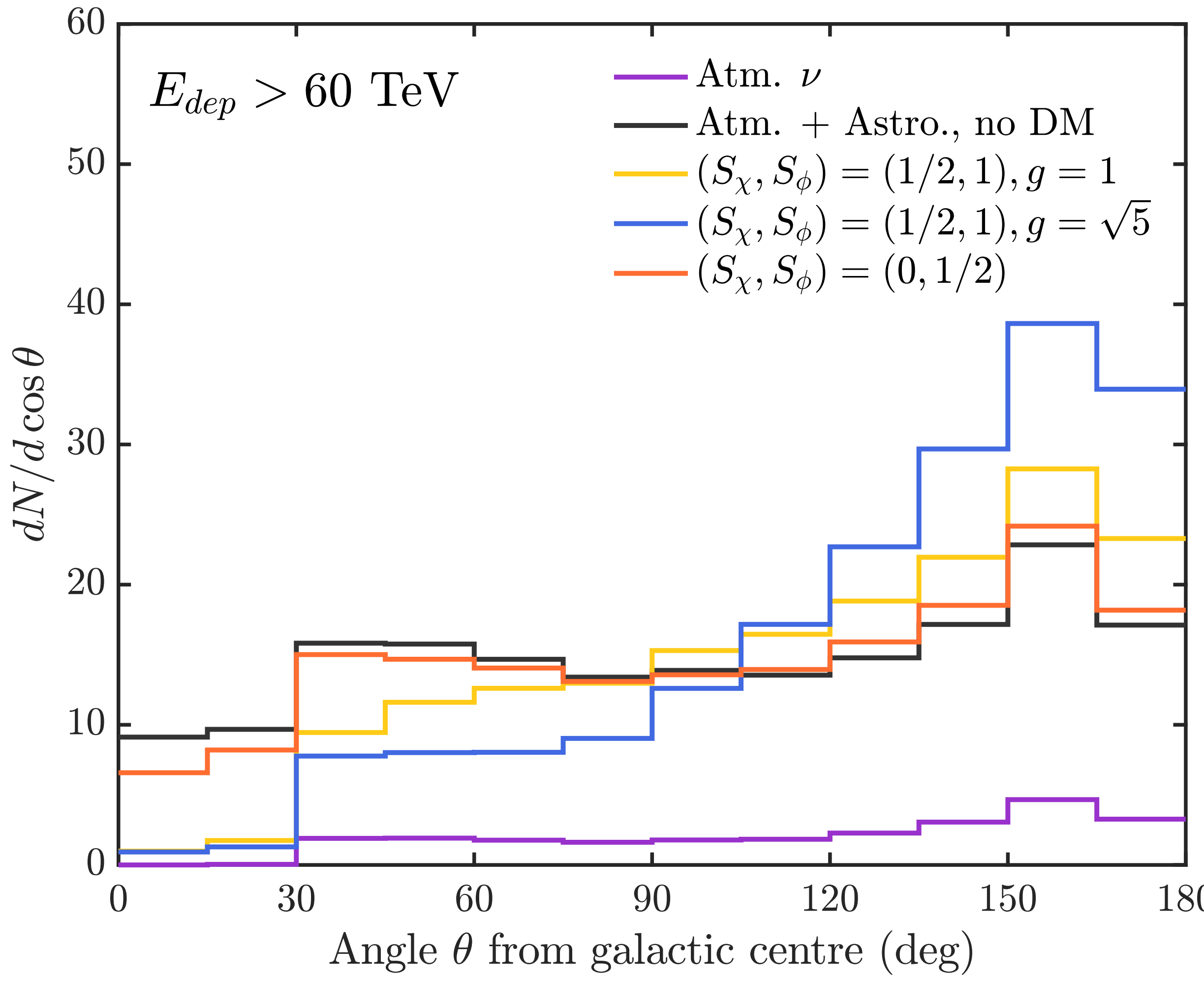
Energy



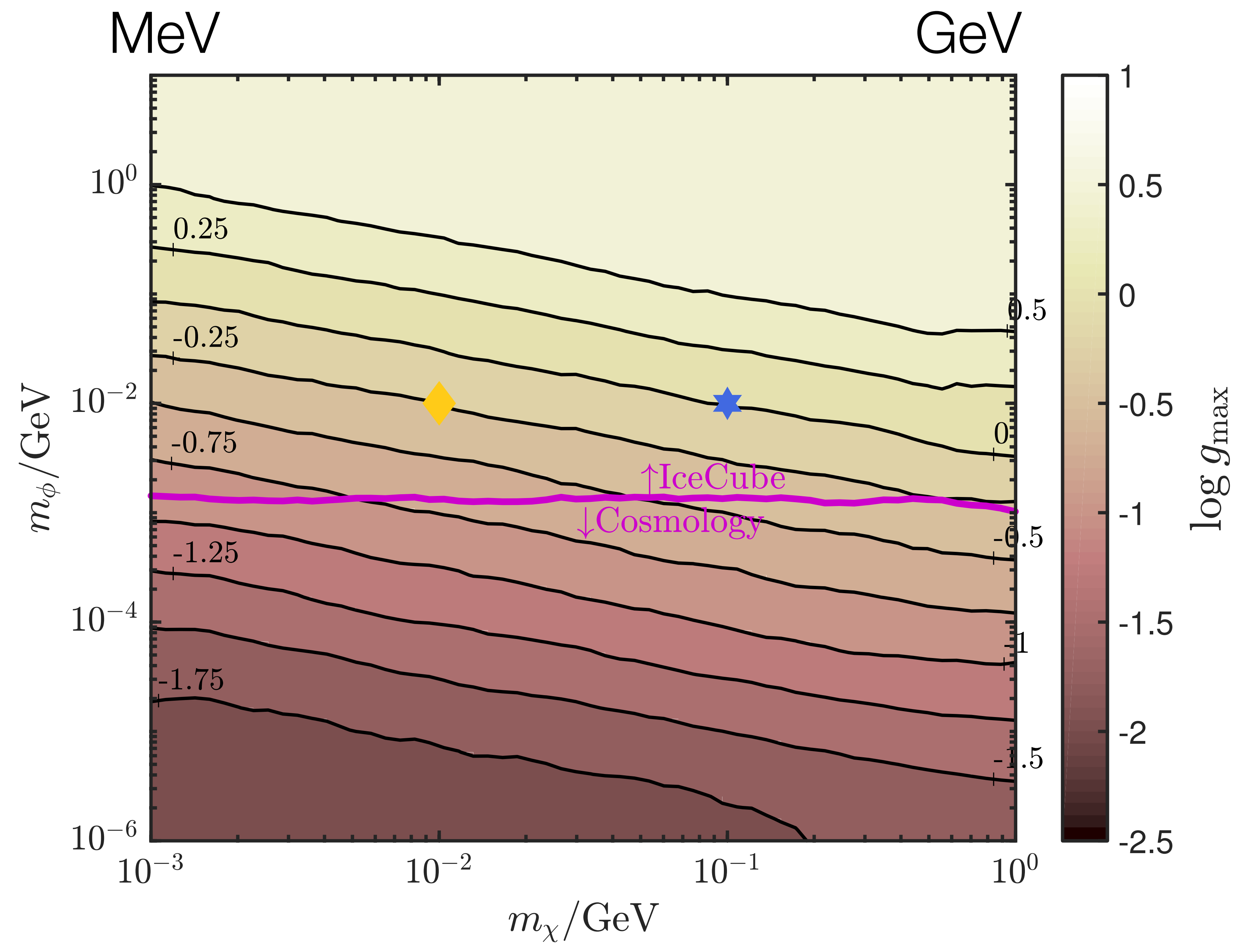
Resonance @ 810 TeV

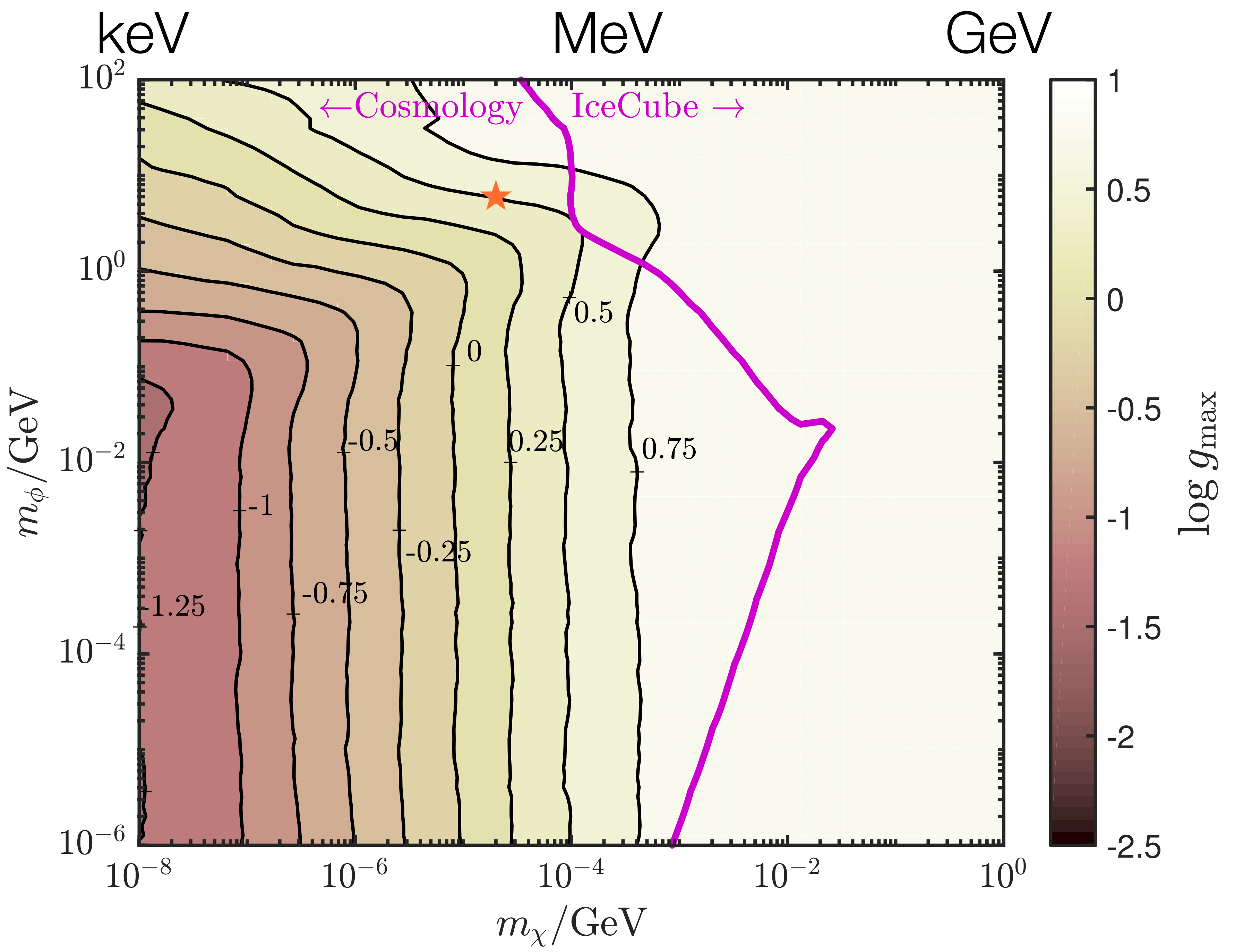
IceCube HESE events

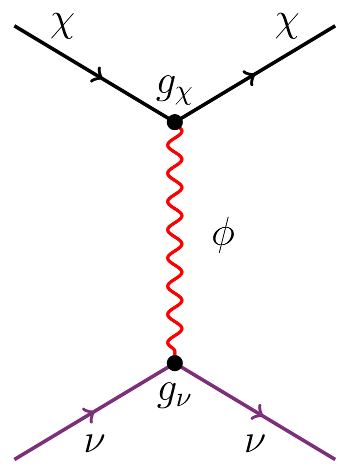
Angle from galactic centre

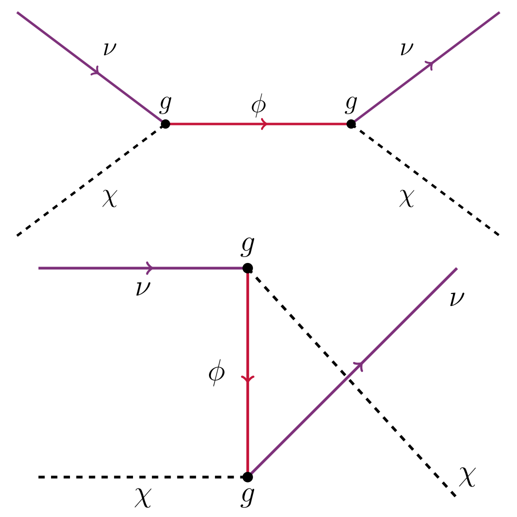


DM-neutrino elastic scattering: 4 years of data







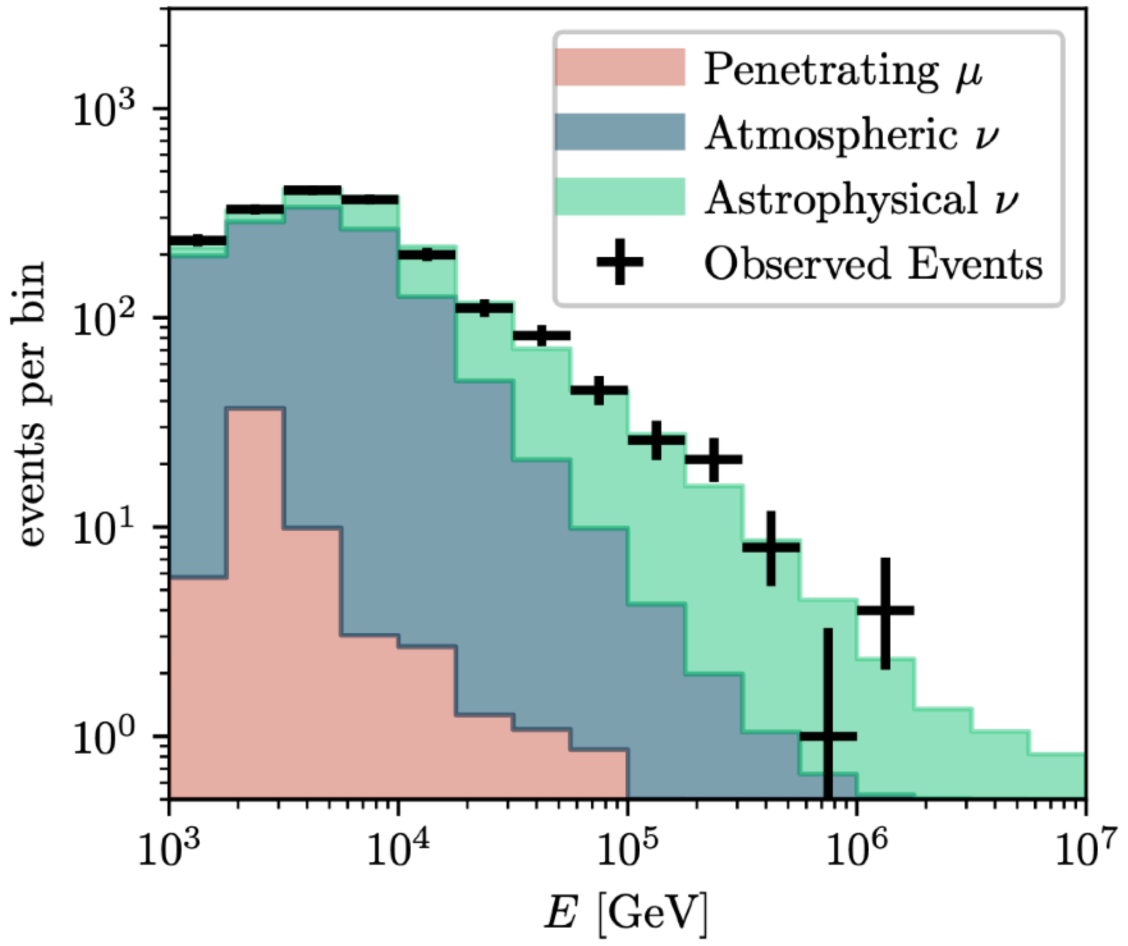


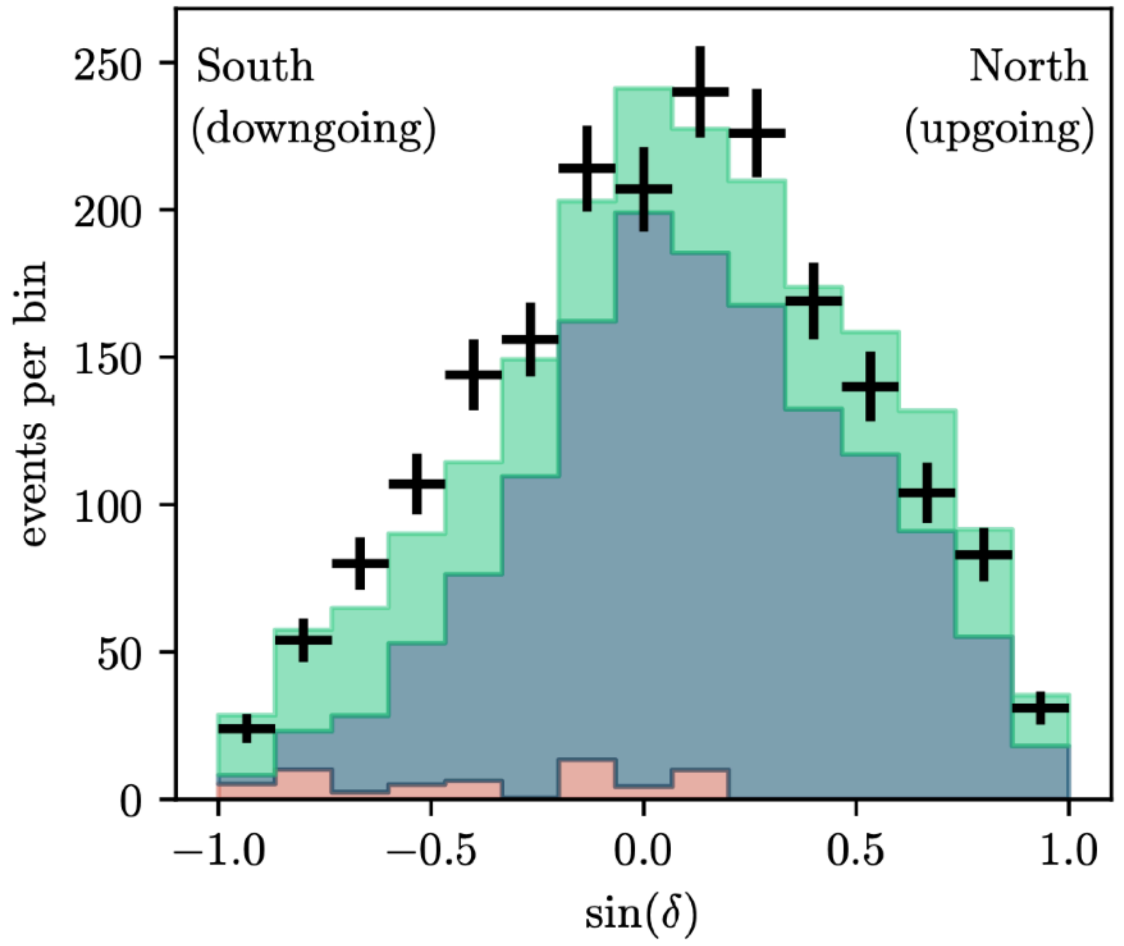
Only 53 events:
already eating into
cosmology parameter
space

DM-neutrino elastic scattering: all energies?

Currently working on official IceCube analysis with O(1000) events.
DarkFATE code, adapted from nuFATE (neutrino fast attenuation through Earth https://github.com/aaronvincent/nuFATE)

IceCube https://arxiv.org/abs/1907.06714





(official IceCube Analysis by Adam McMullen - preliminary results not public)

Large extra dimensions

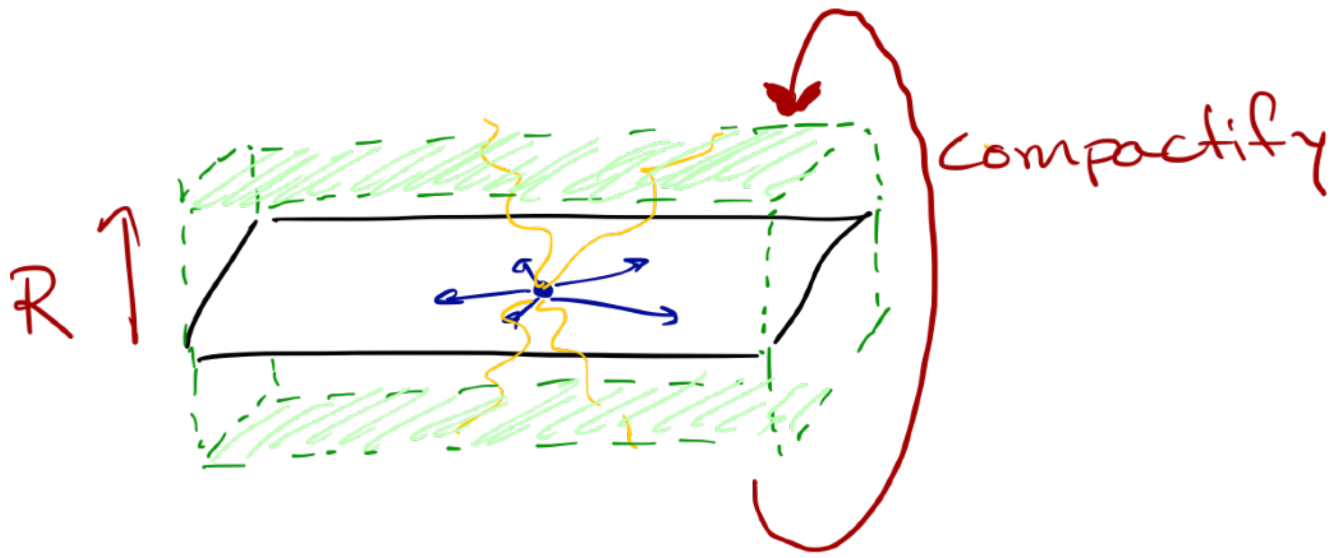
Signatures of microscopic black holes and extra dimensions at future neutrino telescopes

Katherine J. Mack, Ningqiang Song, Aaron C. Vincent

JHEP https://arxiv.org/abs/1912.06656

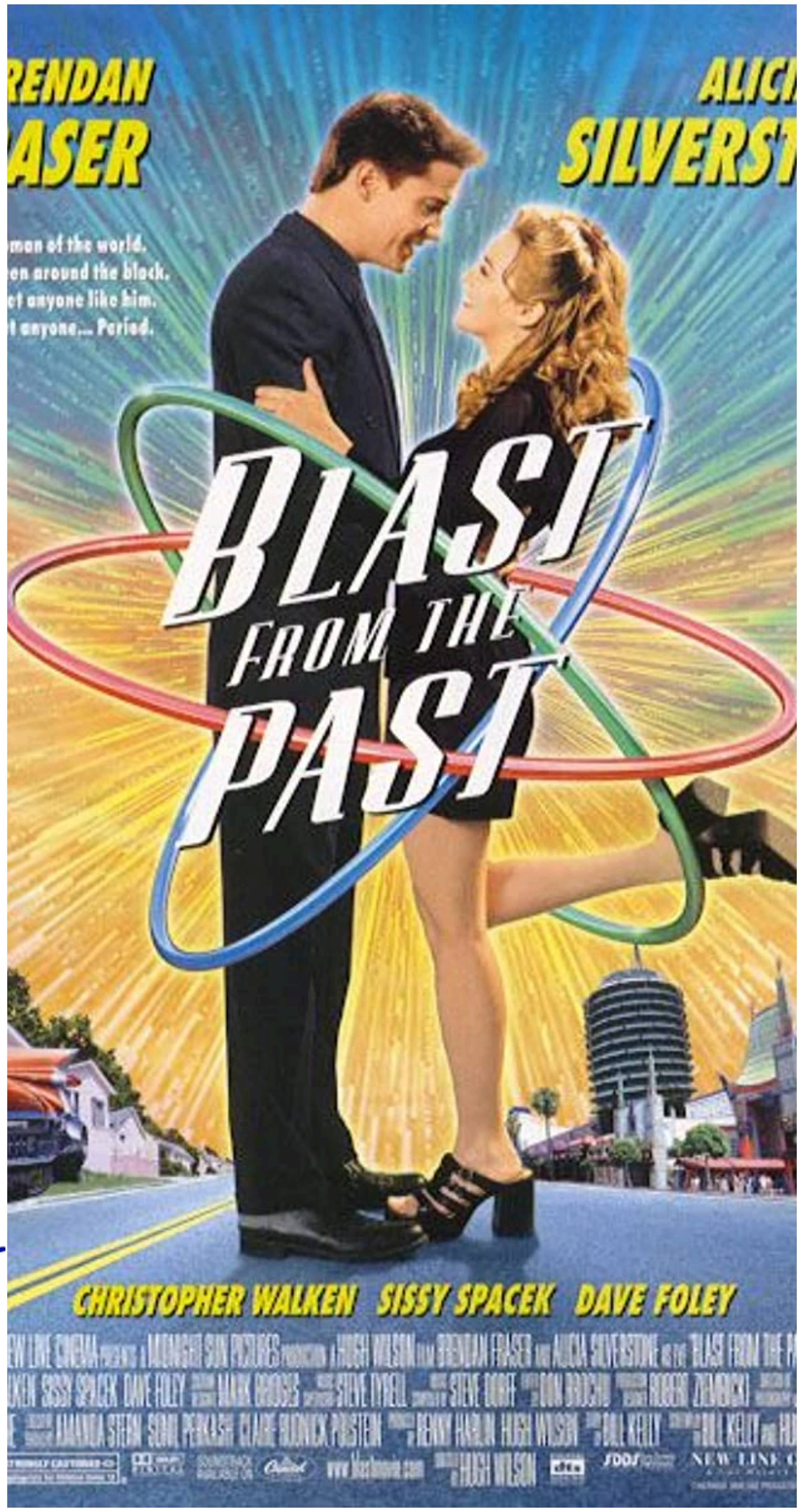
High energies: Black holes from large extra dimensions (ADD)

If the scale of gravity is set by extra dimensions,



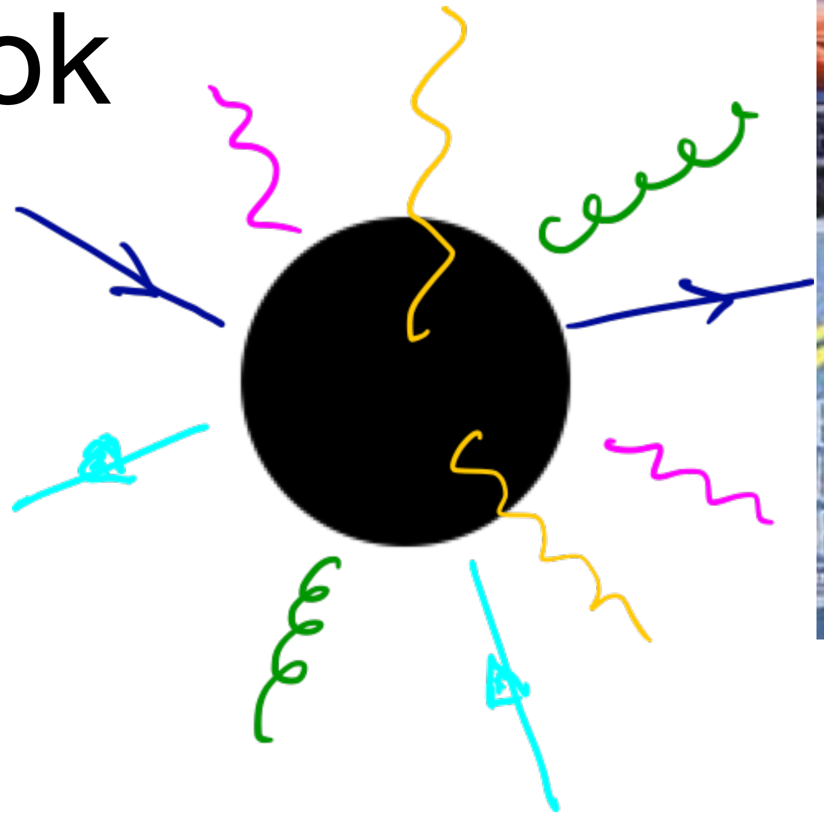
$$V(r) \sim \frac{m_1 m_2}{M_{Pl(4+n)}^{n+2}} \frac{1}{r^{n+1}}, \ (r \ll R).$$

$$V(r) \sim \frac{m_1 m_2}{M_{Pl(4+n)}^{n+2} R^n} \frac{1}{r}, \ (r \gg R)$$



If true Planck scale $M_\star$~10 TeV: can produce microscopic black holes in high-energy collisions.

These evaporate immediately to high-energy products. Since these **anything coupled to gravity**, and most of the standard model is **hadronic**, these showers will look hadronic (as opposed to electroweak).

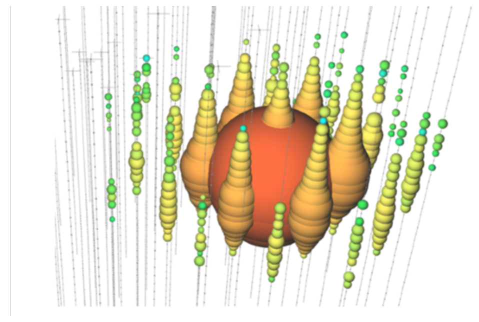


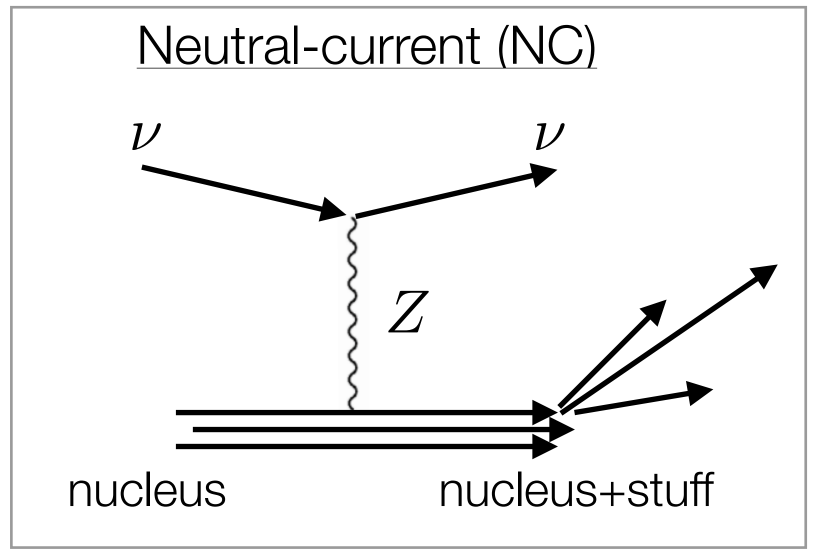
Neutrinos with E > PeV can produce CM collisions at higher energies than LHC

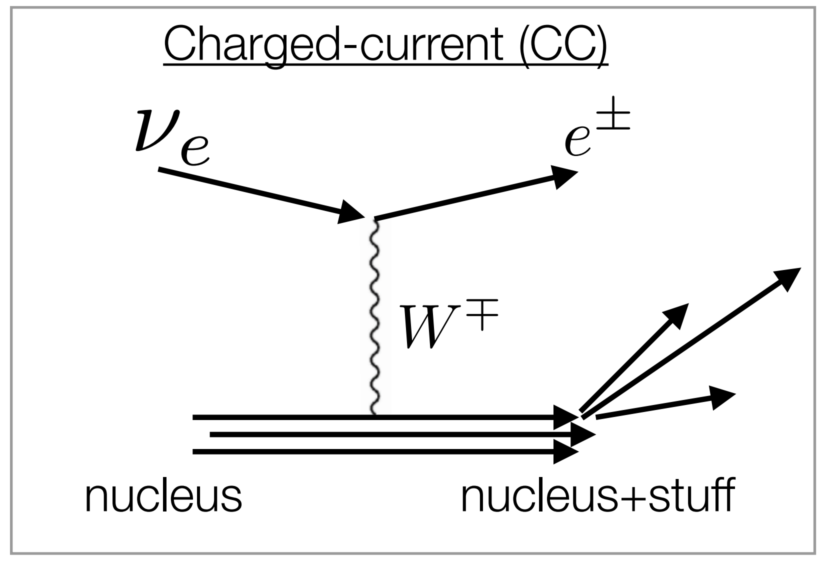
SM

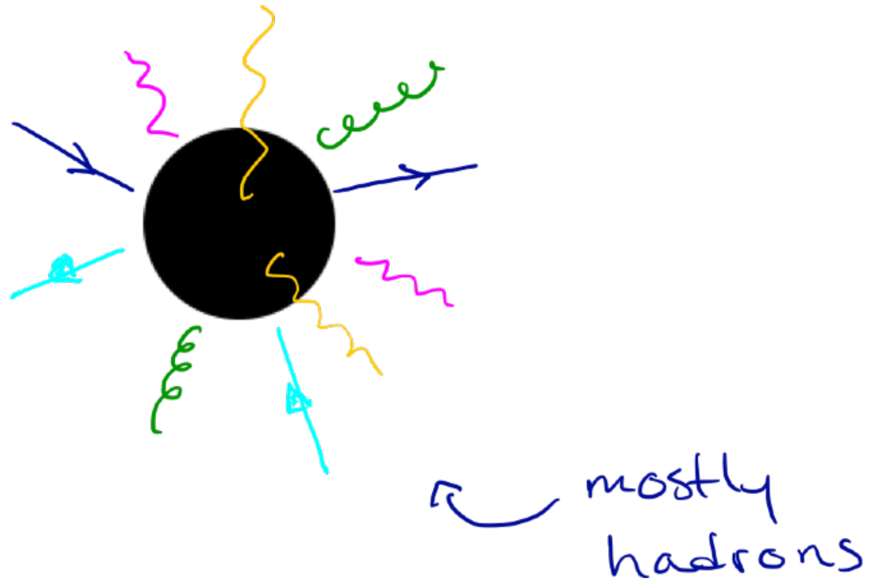
Black hole from $\nu -$nucleus collision

Showers



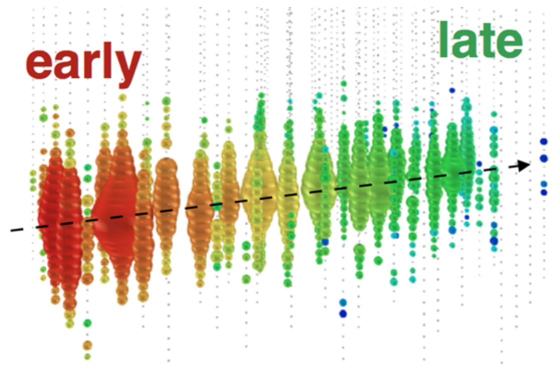


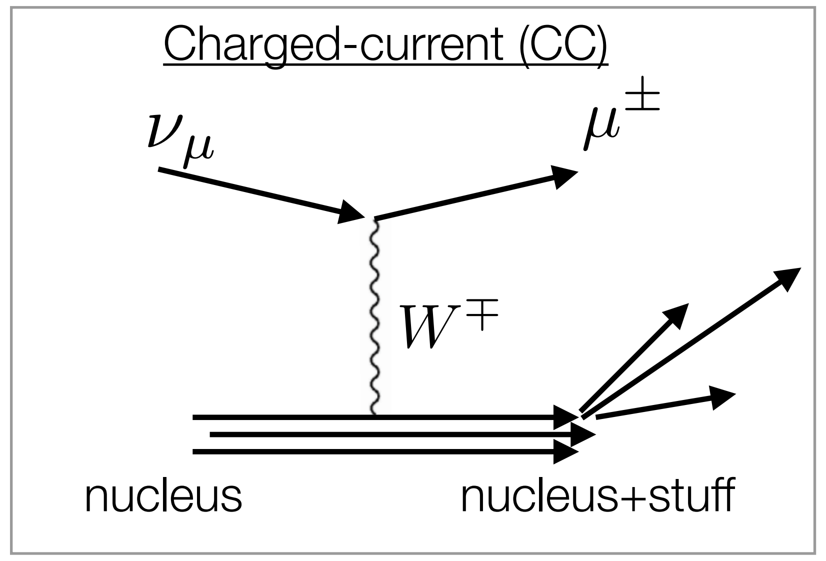


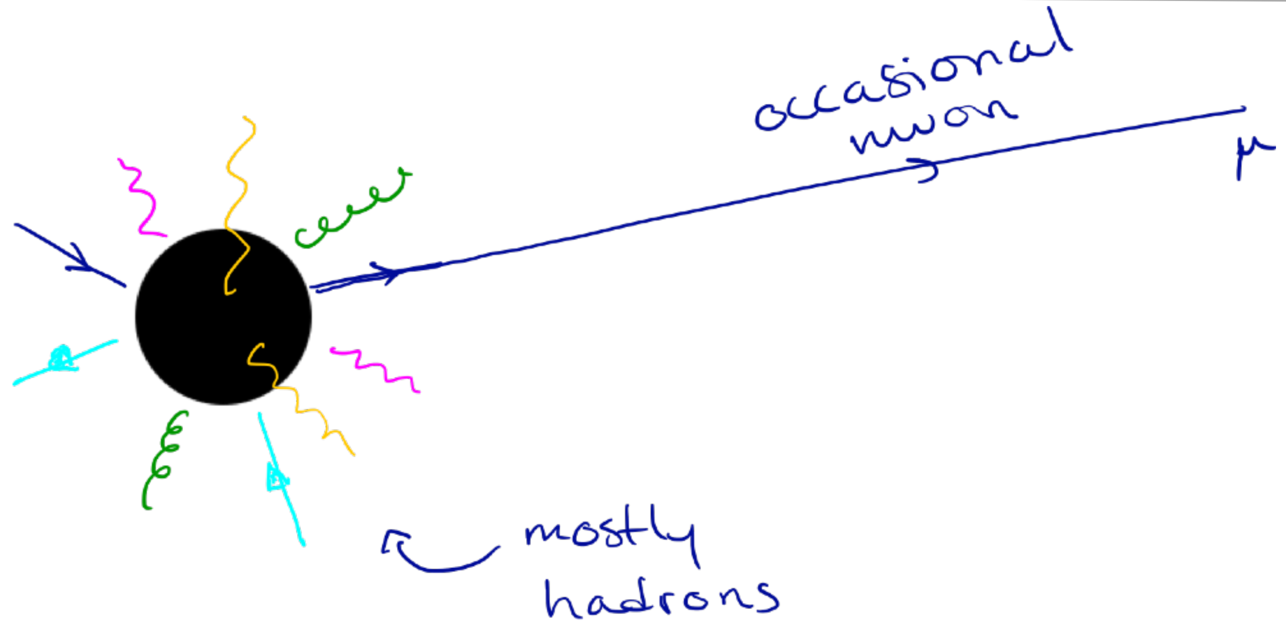


some electrons

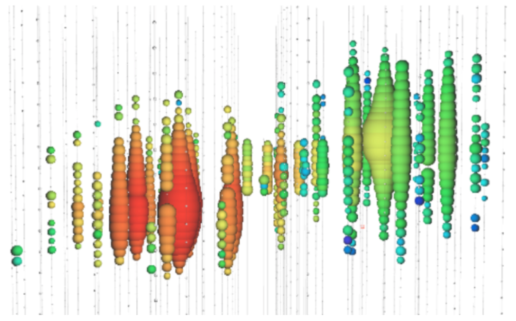
Tracks

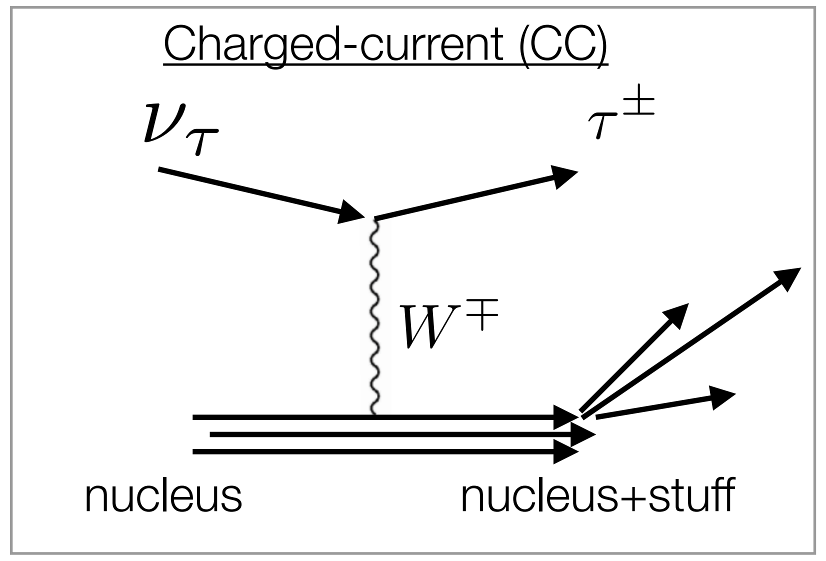


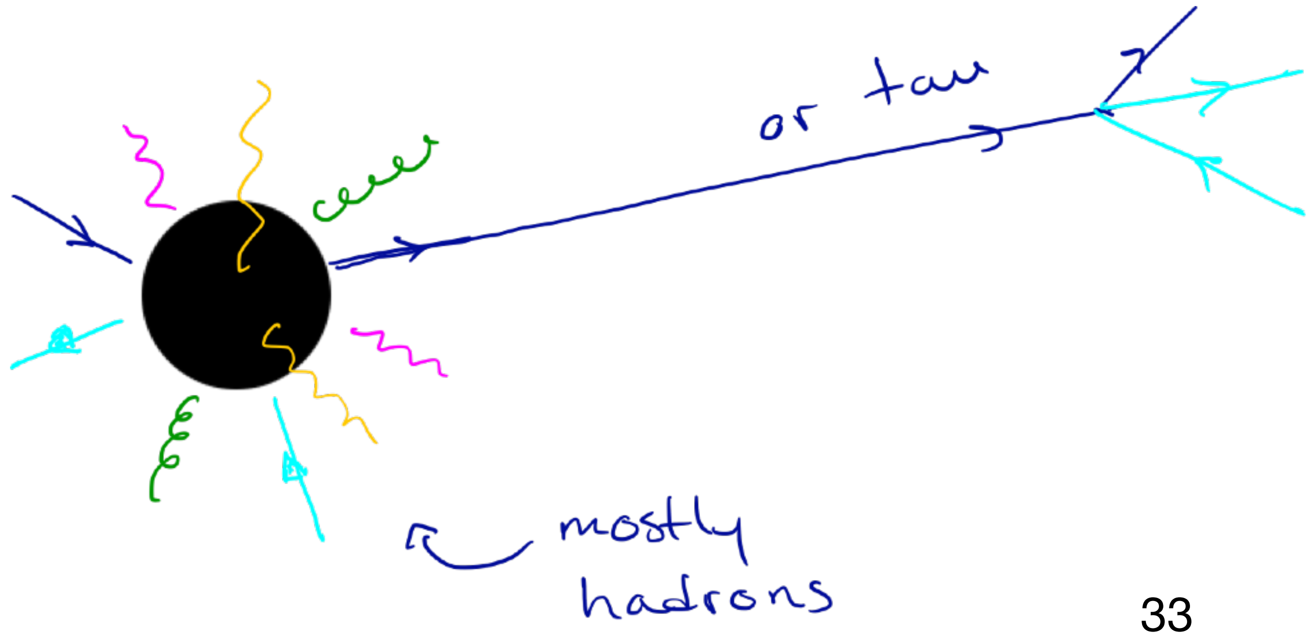




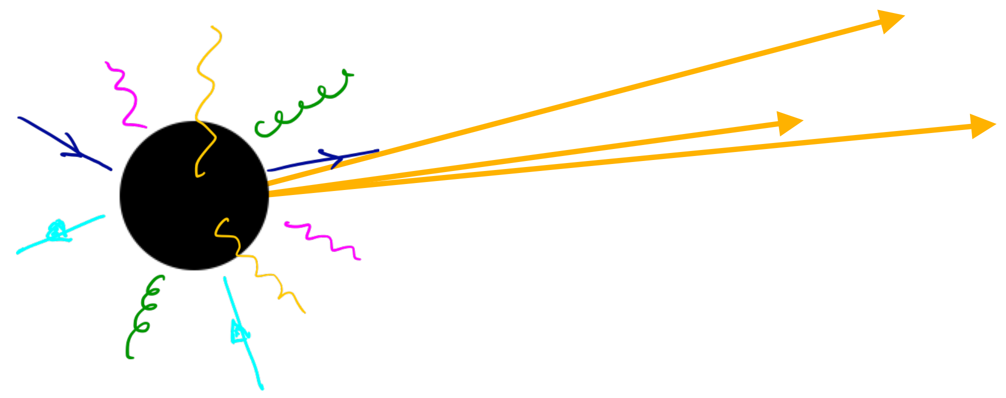
double-bangs



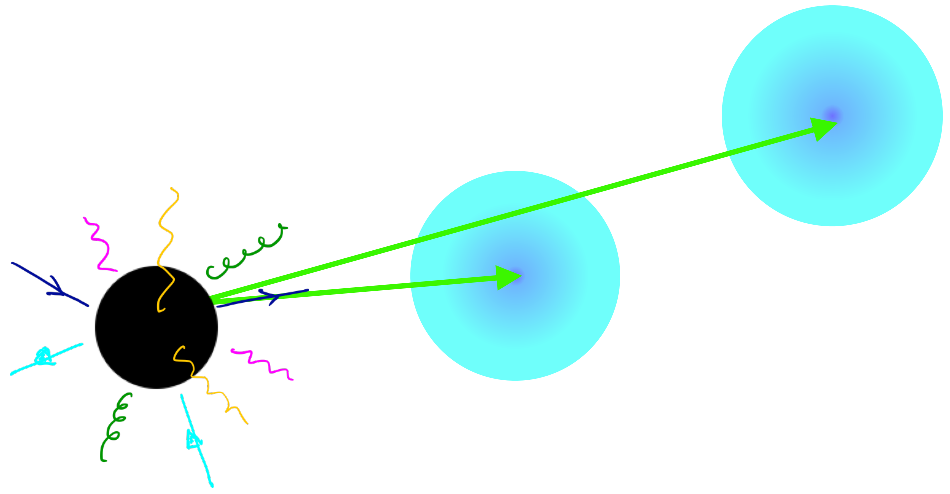




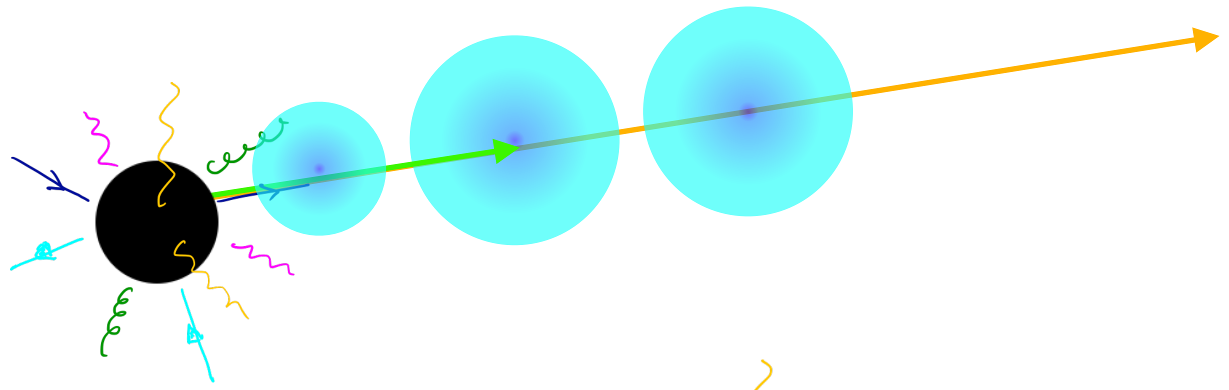
Other crazy morphologies that don't occur in the SM



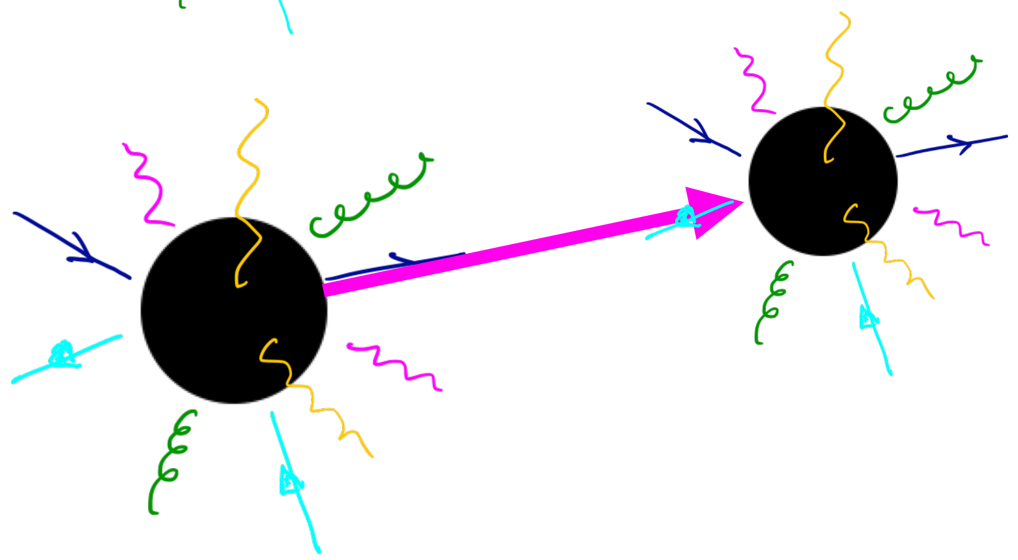
Multitrack (hard to see)



n-bang (*only* 0.2% of black hole events)



***Kebab*:** (About 3% of cases)

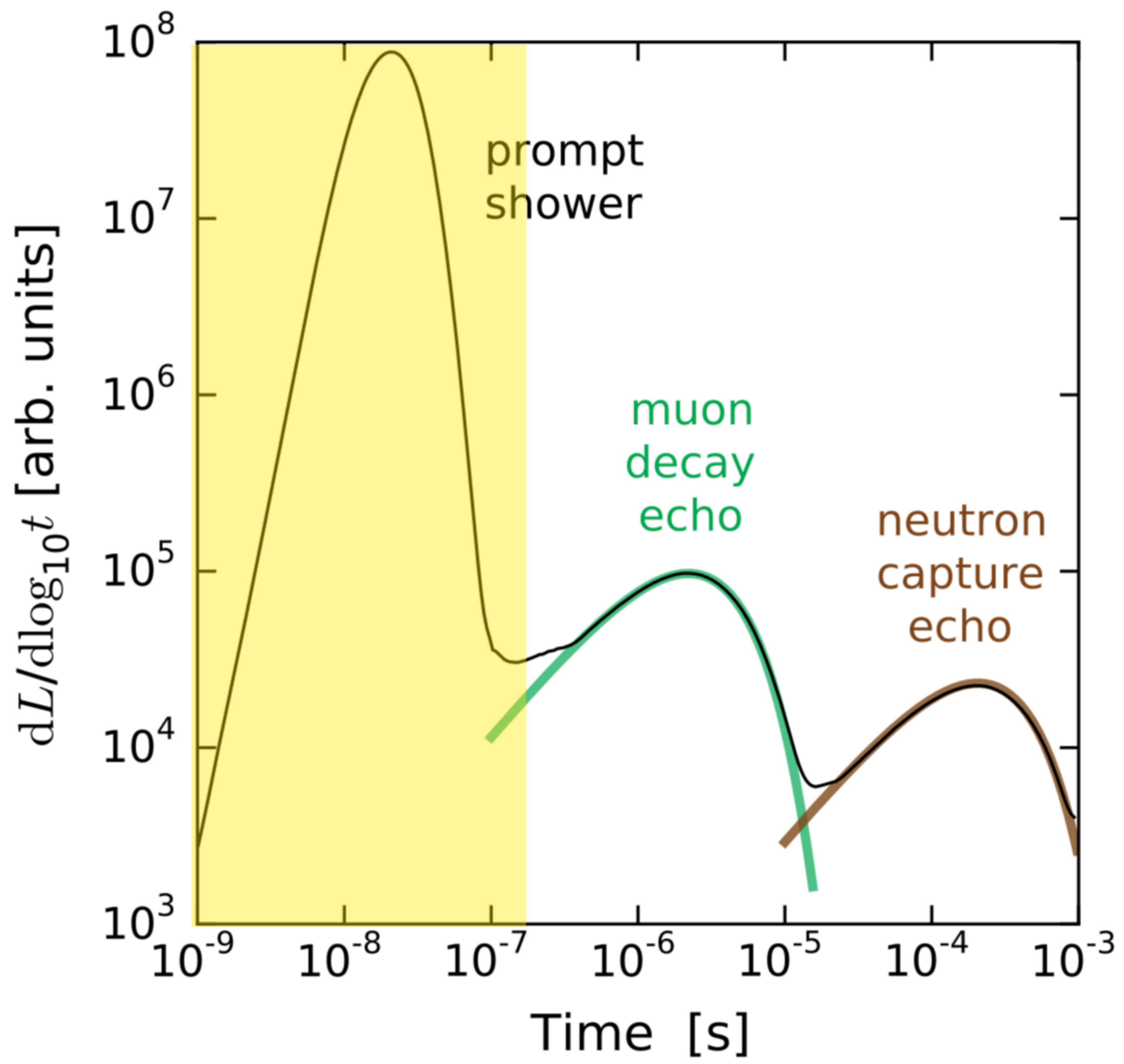


***Double black hole bang*:** (very rare!!)

These are rare, but if we see even one
we can suspect LEDs are involved!!

Hadronic vs electromagnetic energy deposition: **Cherenkov light echoes**

First interaction of neutrinos in ice produces a large **prompt** Cherenkov burst that lasts $\sim 10^{-7}$ s, proportional to the total event energy.



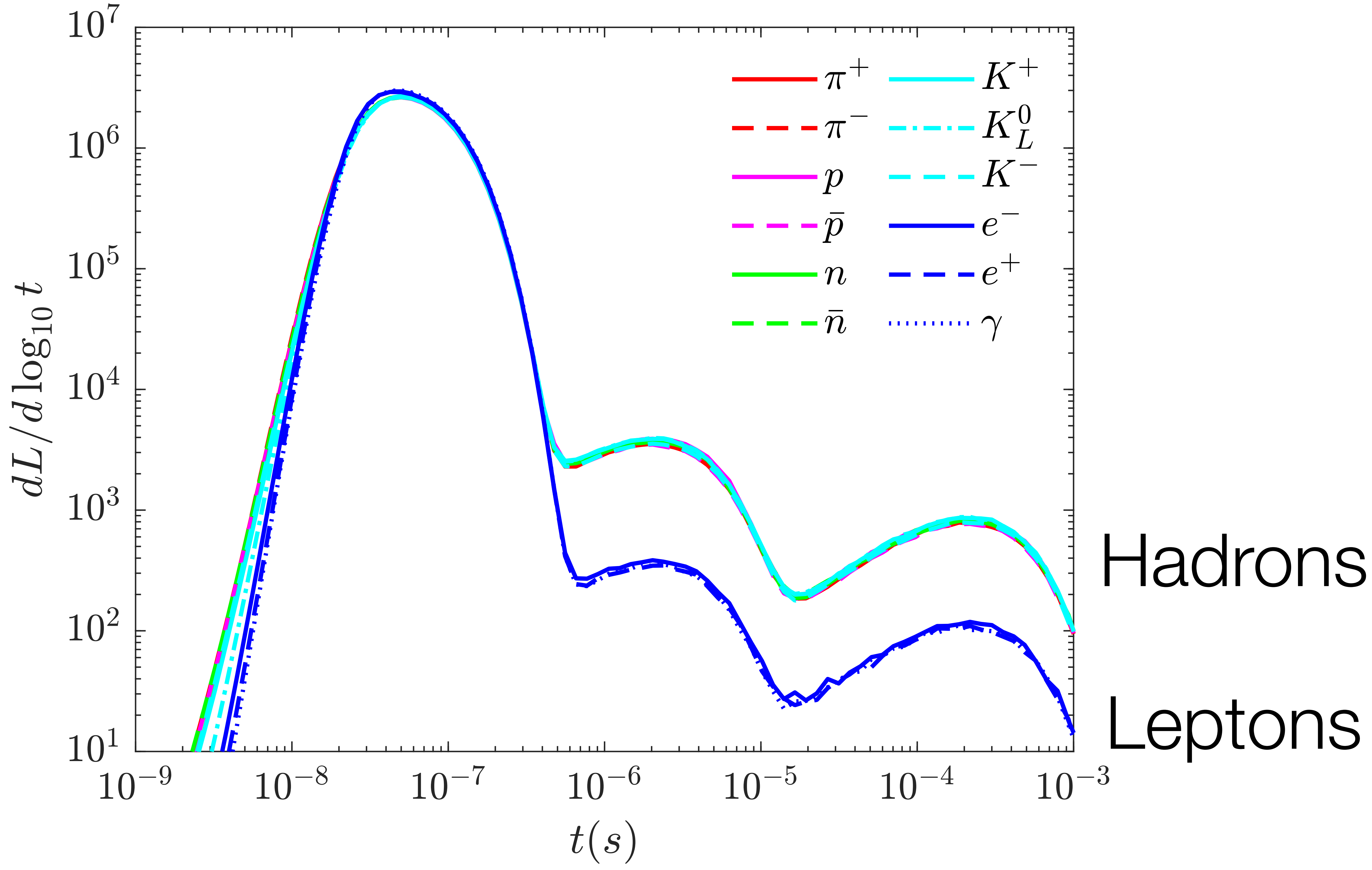
Li, Bustamante, Beacom 1606.06290

Muons can be copiously produced at low energies, and live $\sim 10^{-6}$s, leading to a second **muon echo** as they decay

Neutrons can live for up to .1 ms before being captured, leading to a third **neutron capture echo**

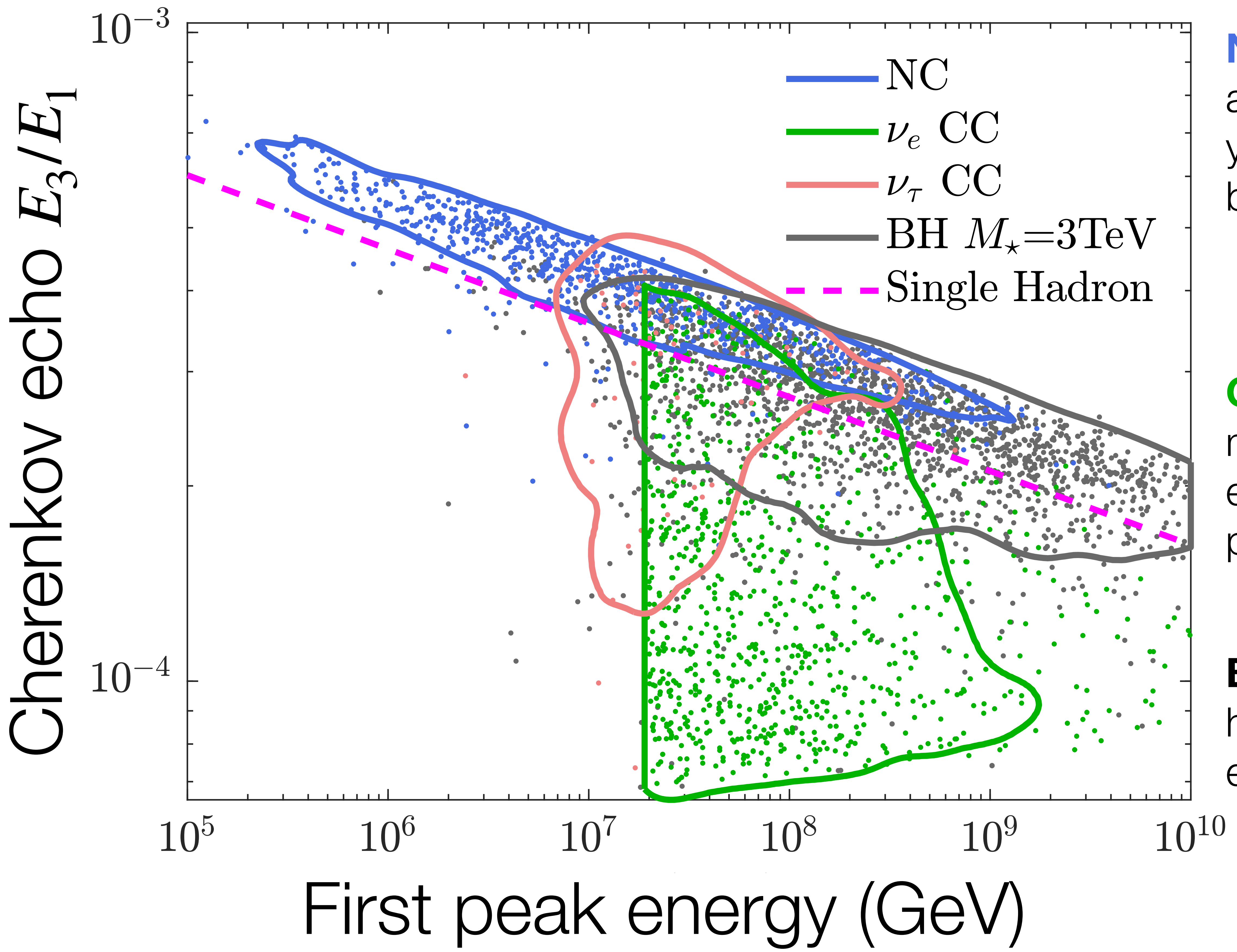
Cherenkov light echoes

Cherenkov light generation for specific particles injected in the ice



Does not seem to be possible at IceCube due to PMT afterpulses, but future telescopes could observe this

Messy in reality (statistics to the rescue!)



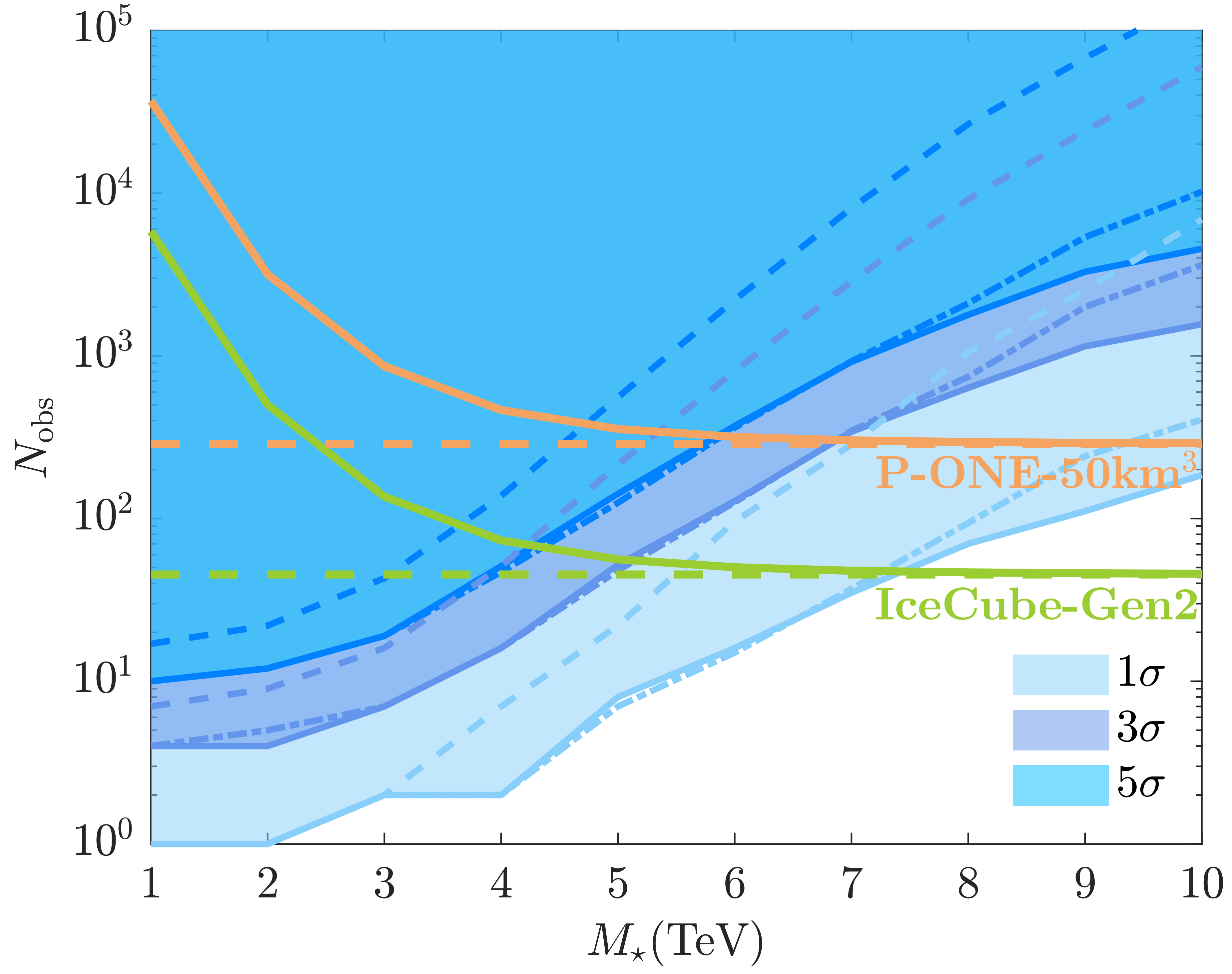
Neutral current events: above the hadron line since hadronization yields mostly hadrons + a few γ. Low energy because neutrino takes away most of the E.

Charged current events: much lower muon/neutron light echo, because most energy injection is from an electron or positron

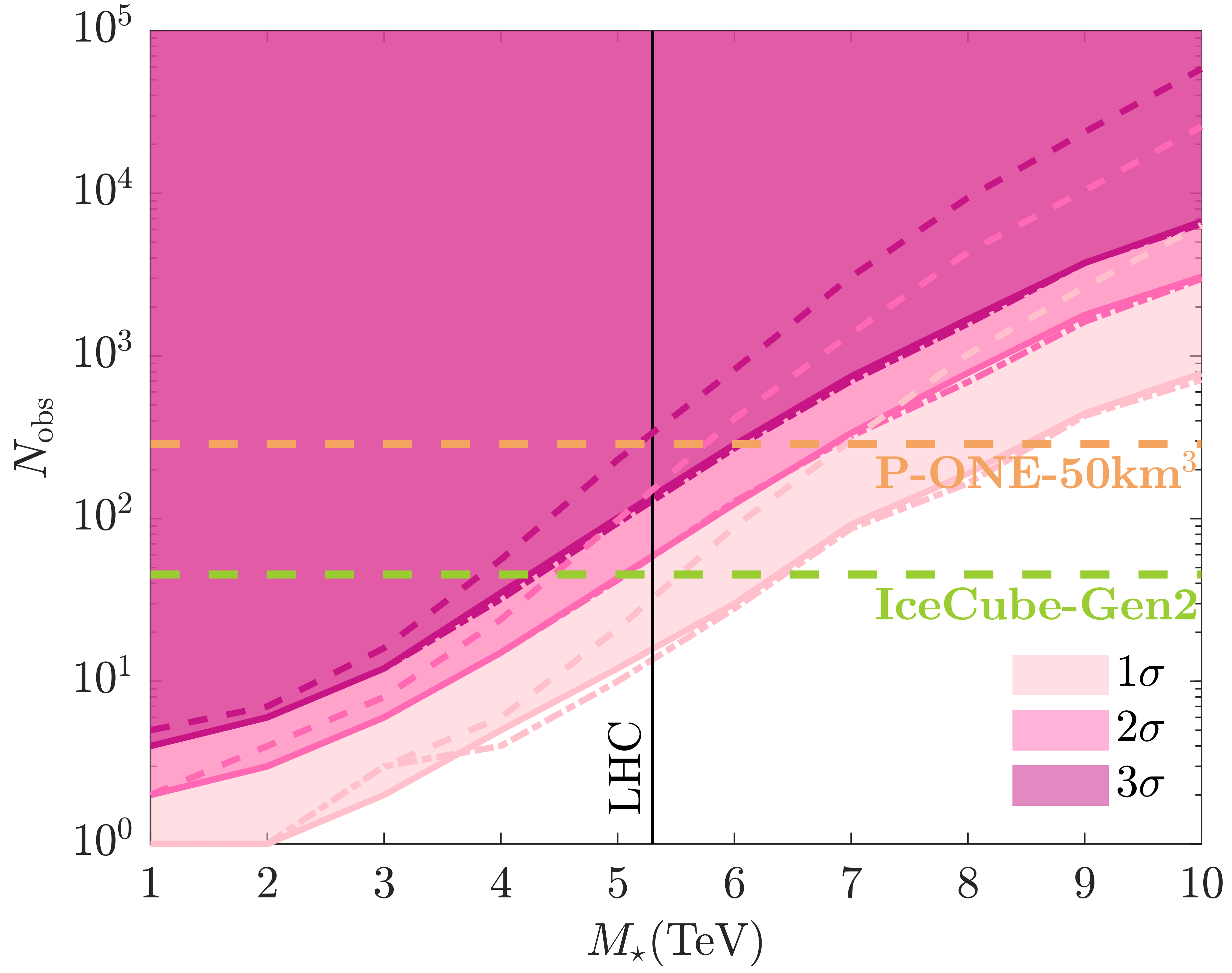
Black Holes: Most of the energy is hadronic: high energy and large Cherenkov echo.

Detecting large extra dimensions with neutrino telescopes

Detection prospects



Exclusion prospects



Summary

- Our understanding of the high-energy neutrino sky will become **1-2 orders of magnitude more precise** over the coming two decades
- Neutrino telescopes cover at least **14 orders of magnitude in energy** & can say all sorts of things about the dark sector & new physics
    - neutrino decay
    - Dark matter
    - large extra dimensions

Thank you