An Analytic Approach to Light Dark Matter Propagation

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March 24, 2023

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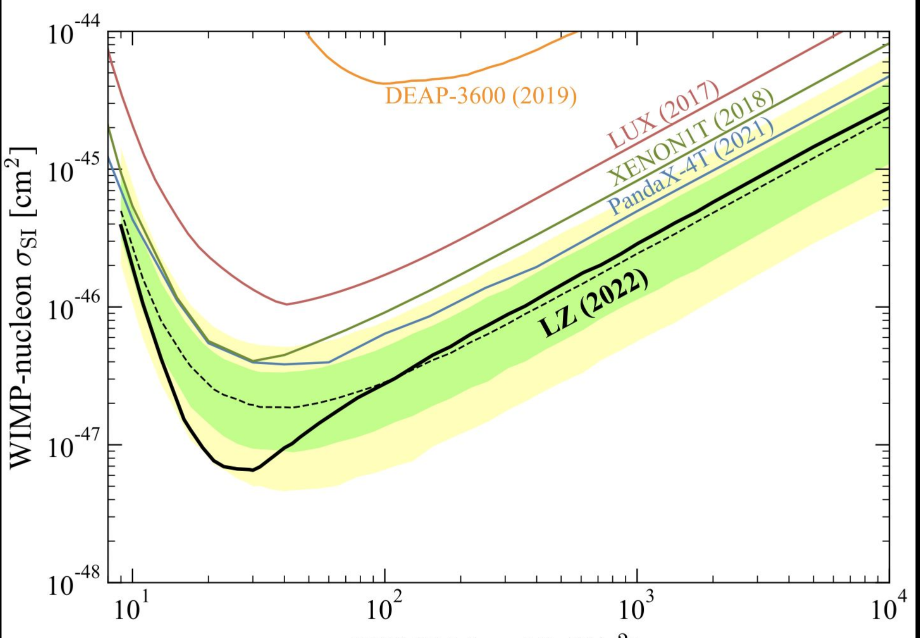
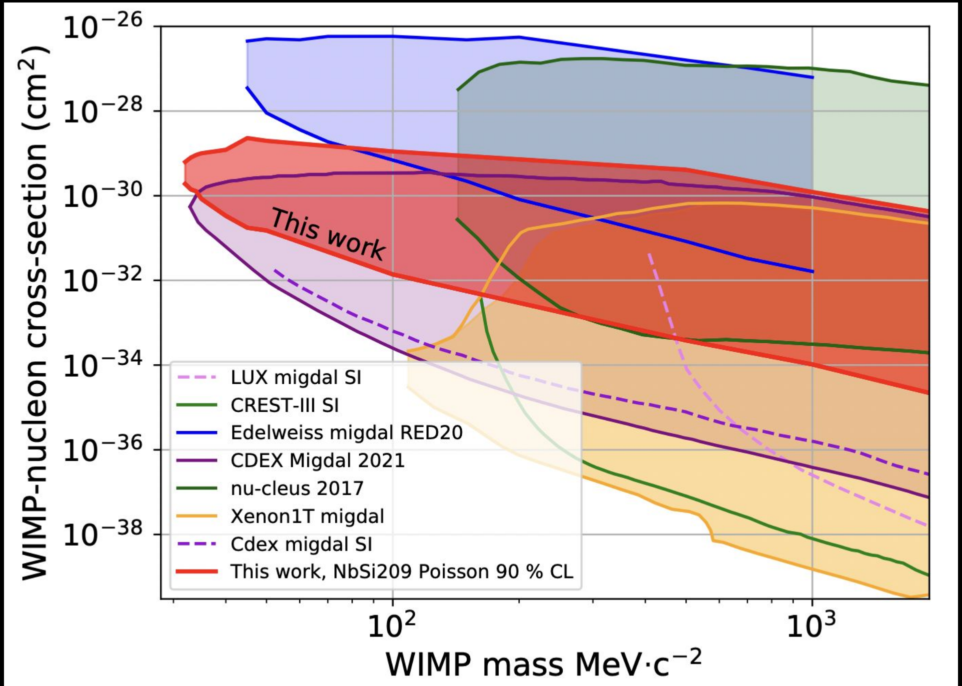
An Intro to Dark Matter Attenuation

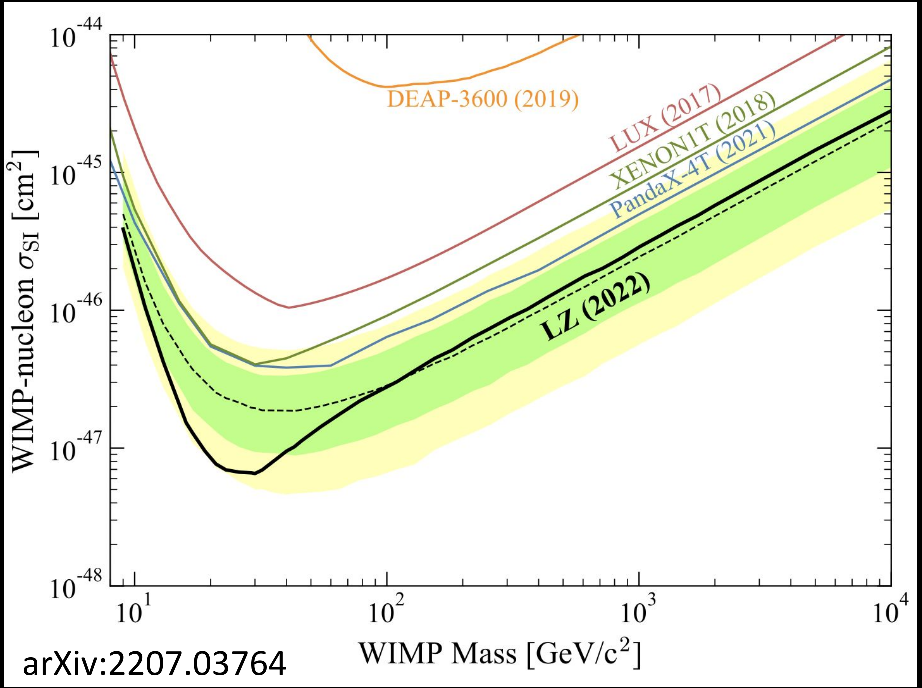
Direct Detection Limits



arXiv:2207.03764

Direct Detection Limits





arXiv:2203.03993 (Edelweiss)

Attenuation: Straight Line Approximation

PHYSICAL REVIEW D VOLUME 41, NUMBER 12 15 JUNE 1990

Opening the window on strongly interacting dark matter

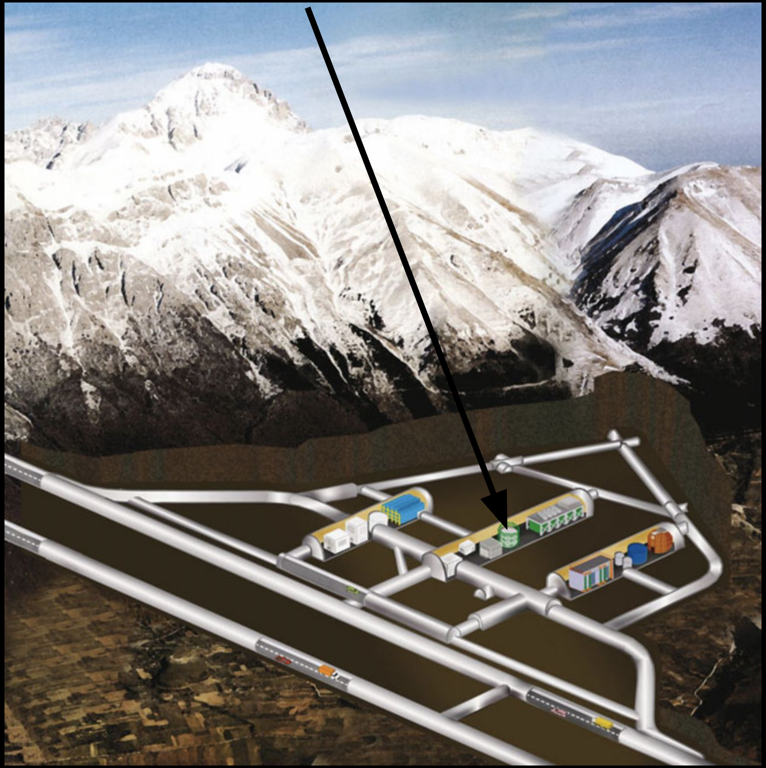
Glenn D. Starkman and Andrew Gould
Institute for Advanced Study, Princeton, New Jersey 08540

Rahim Esmailzadeh
Center for Particle Astrophysics, University of California, Berkeley, California 94720

Savas Dimopoulos*
CERN TH-Division, 1211 Geneva 23, Switzerland
(Received 2 February 1990)

We discuss the possibility that the dark matter consists of strongly interacting massive particles (SIMP's) which have cross sections with ordinary matter which are larger than characteristic weak-interaction cross sections. We show that, while results from $\beta\beta$ decay, cosmic-ray detectors, galactic-halo stability, the cooling of molecular clouds, proton-decay detectors, and the existence of old neutron stars and the Earth constrain the interactions of the missing matter with ordinary matter over a broad range of parameter space, there still exist several windows for SIMP's. It is noteworthy that there are two regions of less than geometric cross sections: one with masses of 10^5-10^7 GeV and another with masses above 10^{10} GeV.

Attenuation: Straight Line Approximation



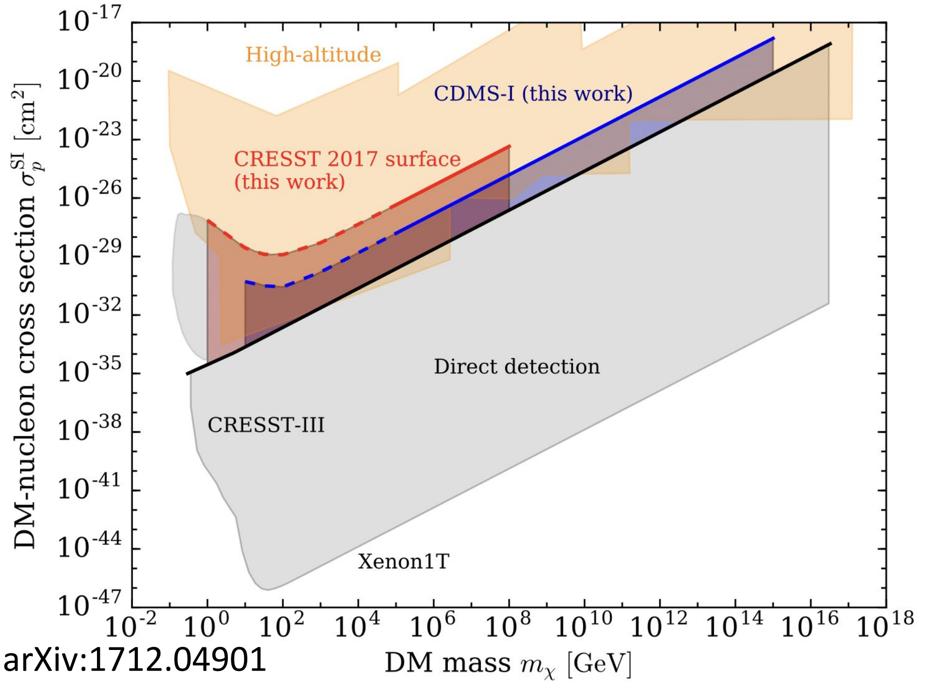
Assume that dark matter follows a straight (ballistic) trajectory, every particle losing the average possible amount of energy:

$$\frac{\mathrm{d}\langle E_\chi \rangle}{\mathrm{d}t} = -\sum_i n_i(\mathbf{r}) \, \langle E_R \rangle_i \, \sigma_i(v) \, v \,,$$

- n_i : number density of nuclei in overburden
- $\langle E_R \rangle_i$: average recoil energy
- σ_i : DM-nucleus cross section

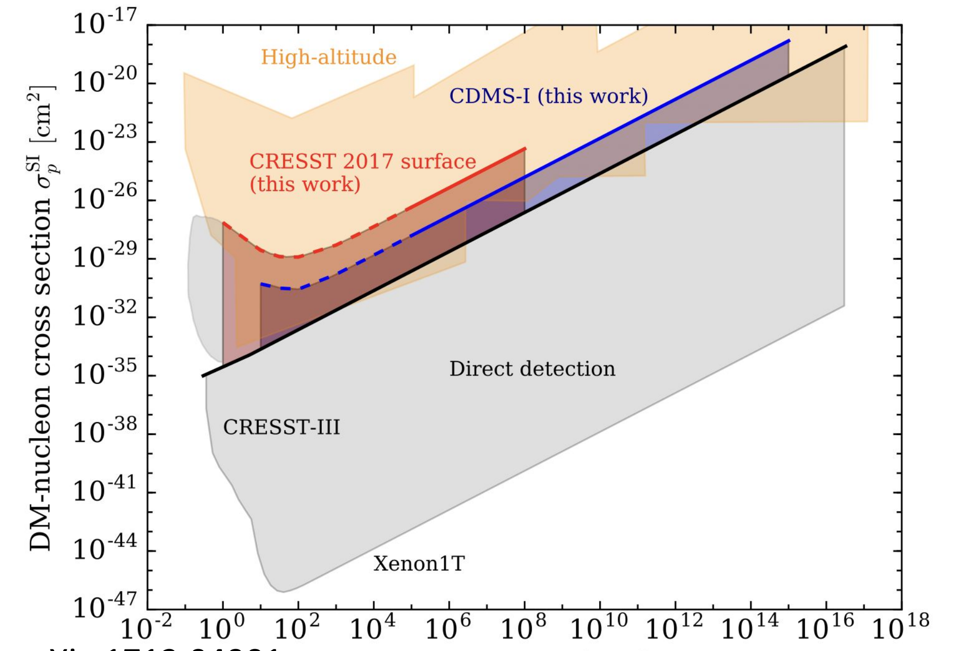
Left: Schematic of LNGS
-Credit: https://www.appec.org

Attenuation: Straight Line Approximation



arXiv:1712.04901

Attenuation: Straight Line Approximation



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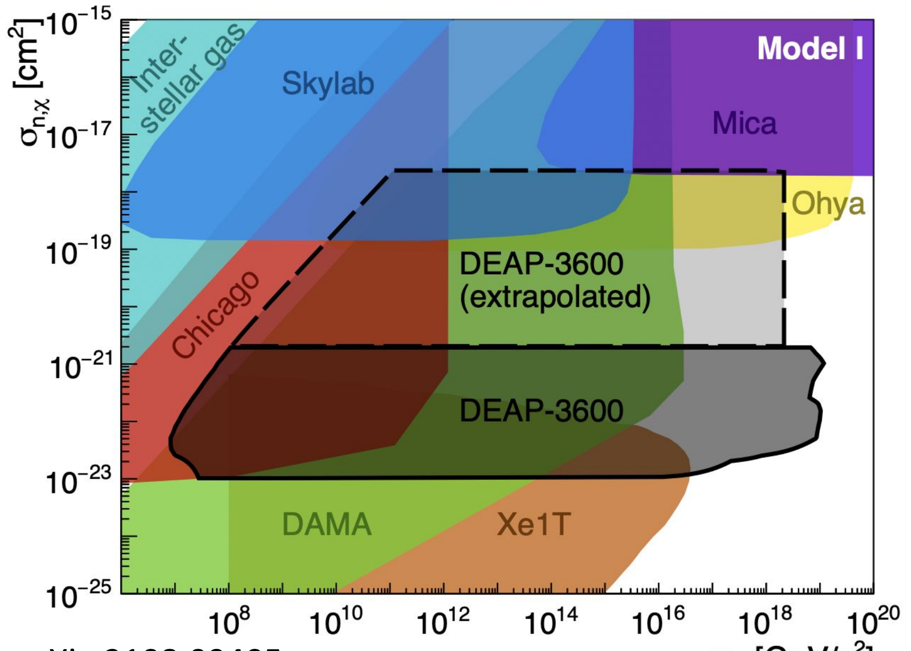
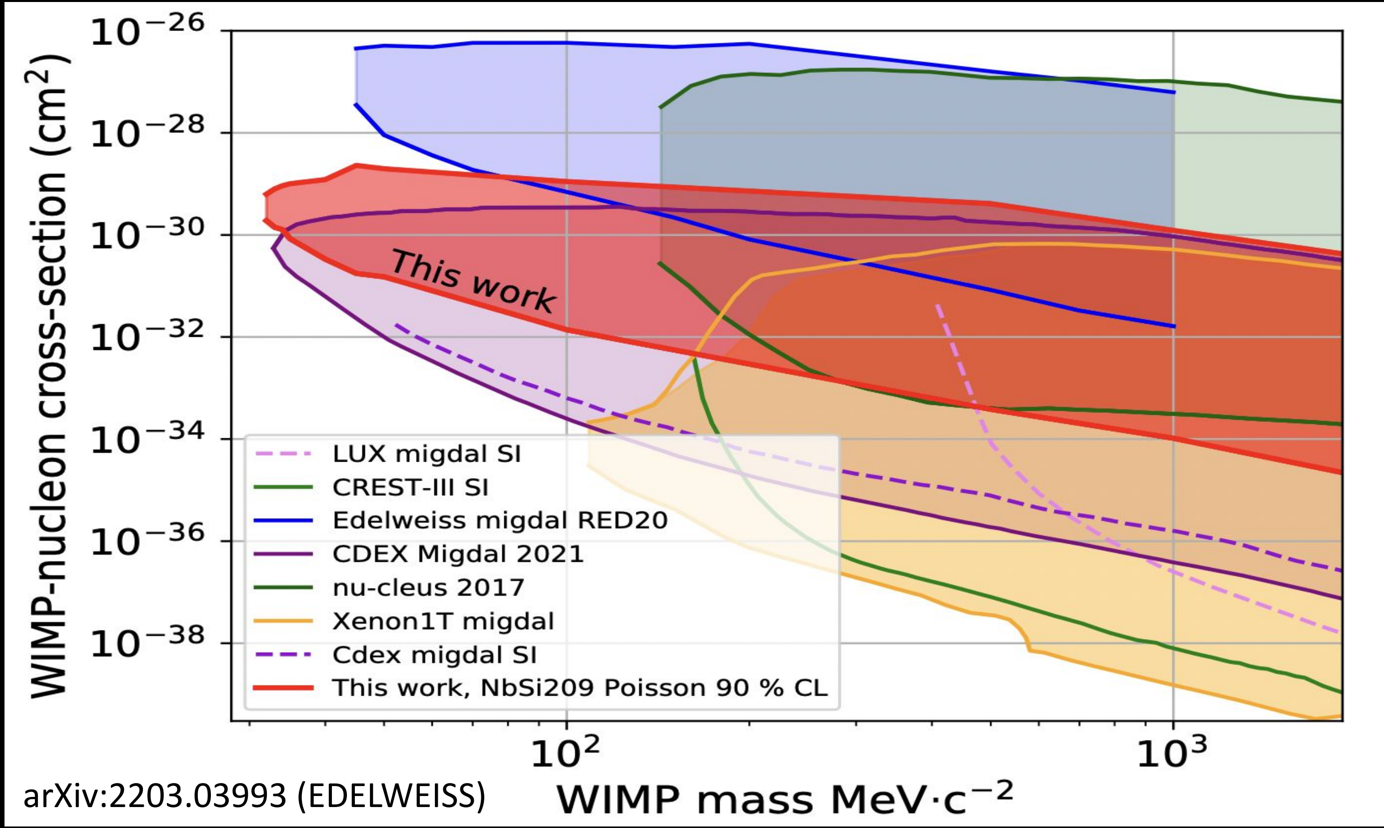


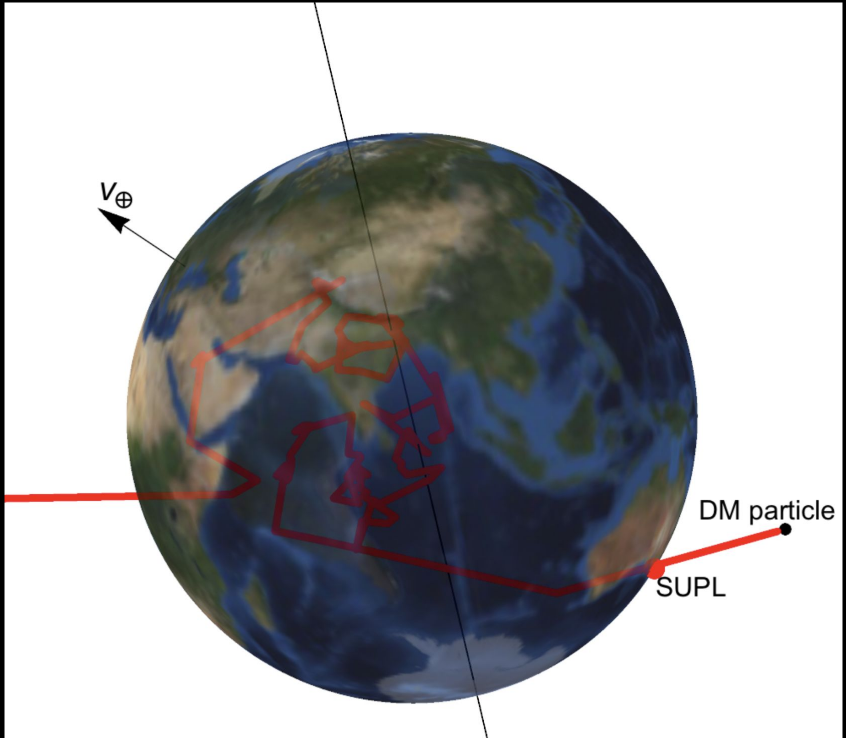
arXiv:2108.09405

Attenuation: Straight Line Approximation



arXiv:2203.03993 (EDELWEISS)

Attenuation: Monte Carlo Simulations

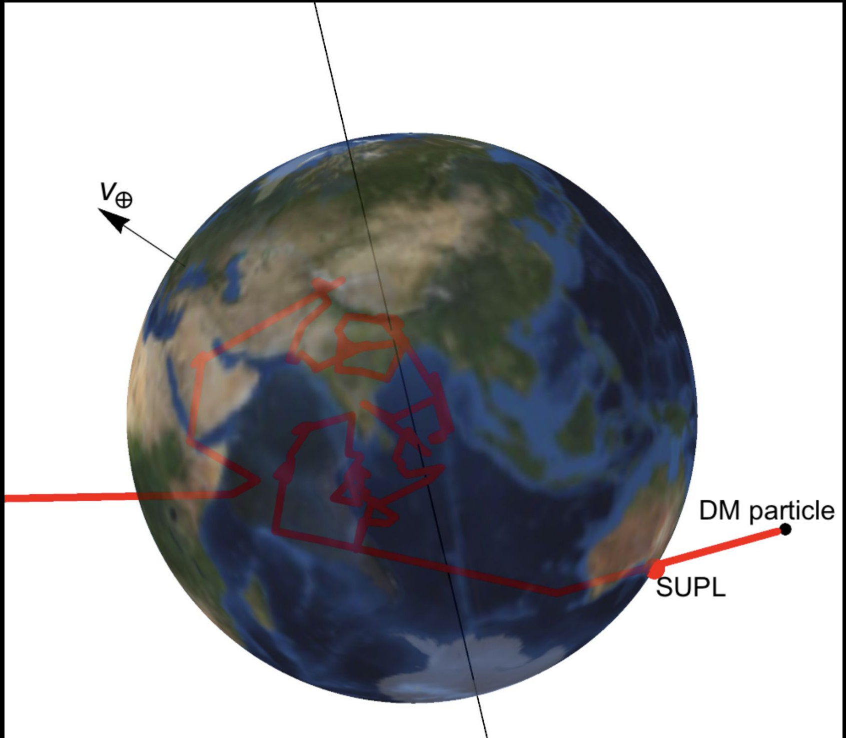


timonemken.com/code, arXiv:1706.02249

Alternative approach: numerically simulate millions of dark matter particles scattering in Earth

- Initialize a particle above the Earth with initial velocity **v**
- Choose distance particle travels without scattering from exponential path length distribution
- Choose target + scattering angle, update velocity
- Stop if particle is detected, loses too much energy, or leaves Earth

Attenuation: Monte Carlo Simulations



timonemken.com/code, arXiv:1706.02249

Examples:

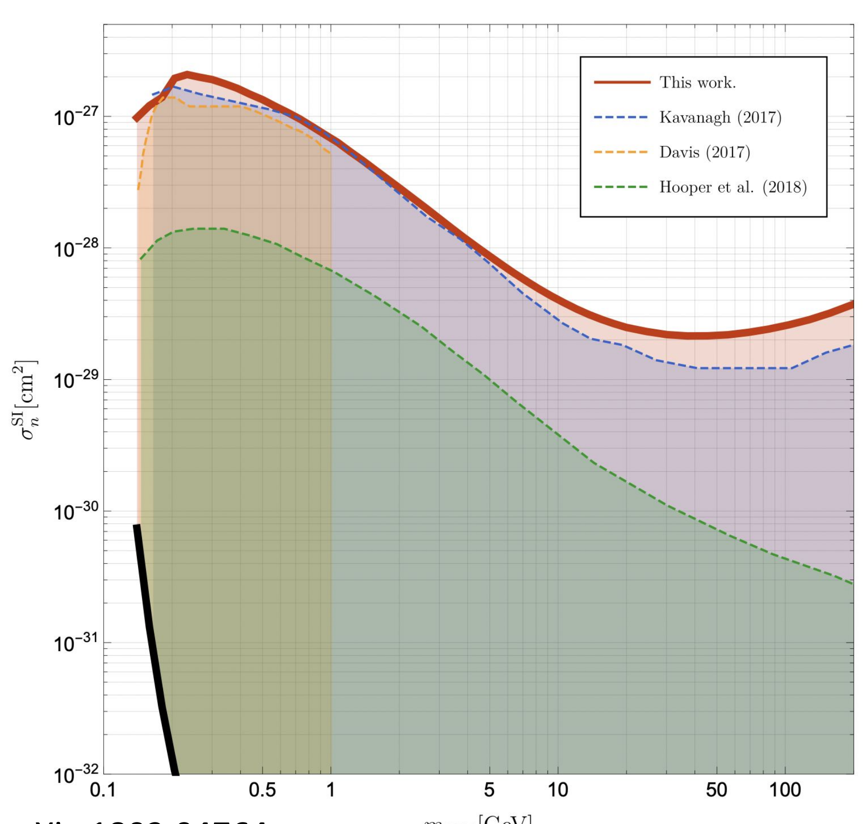
- DaMaSCUS: Dark Matter Simulation Code for Underground Scatterings, Timon Emken
- DMATIS: Dark Matter Attenuation Importance Sampling, M. Shafi Mahdawi
- DarkProp (for cosmic ray boosted dark matter), Chen Xia

See also:

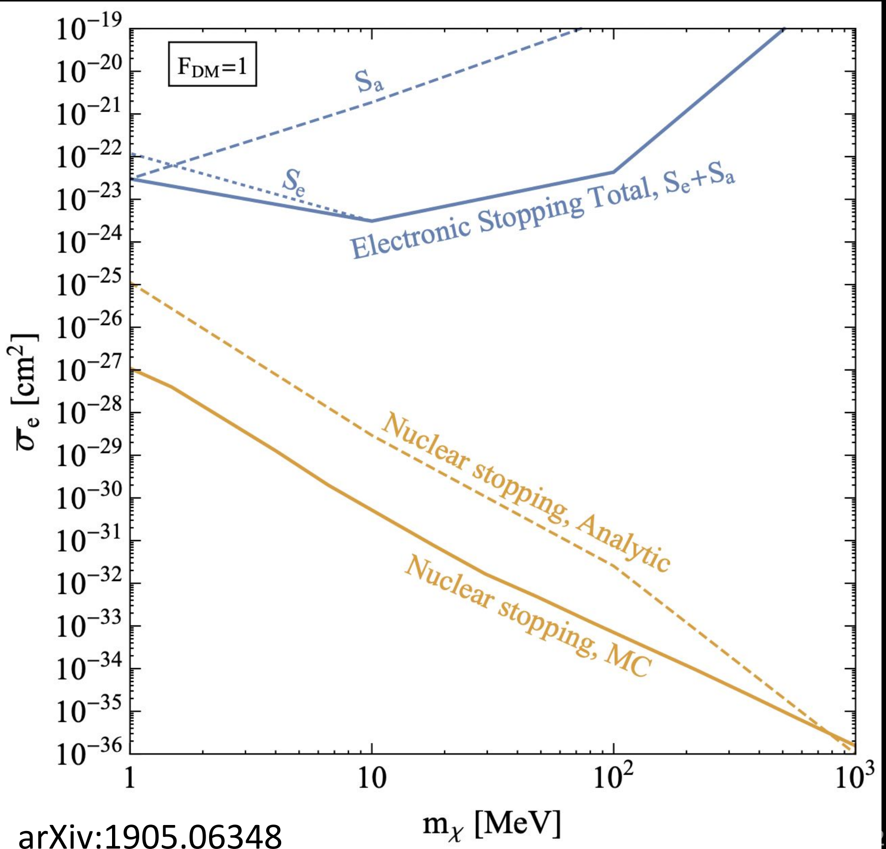
CVC and Beacom: PRD 100, 103011 (2019)

PROSPECT Collaboration, **CVC**: PRD 104, 012009 (2021)

Straight Line vs. Monte Carlo



arXiv:1802.04764

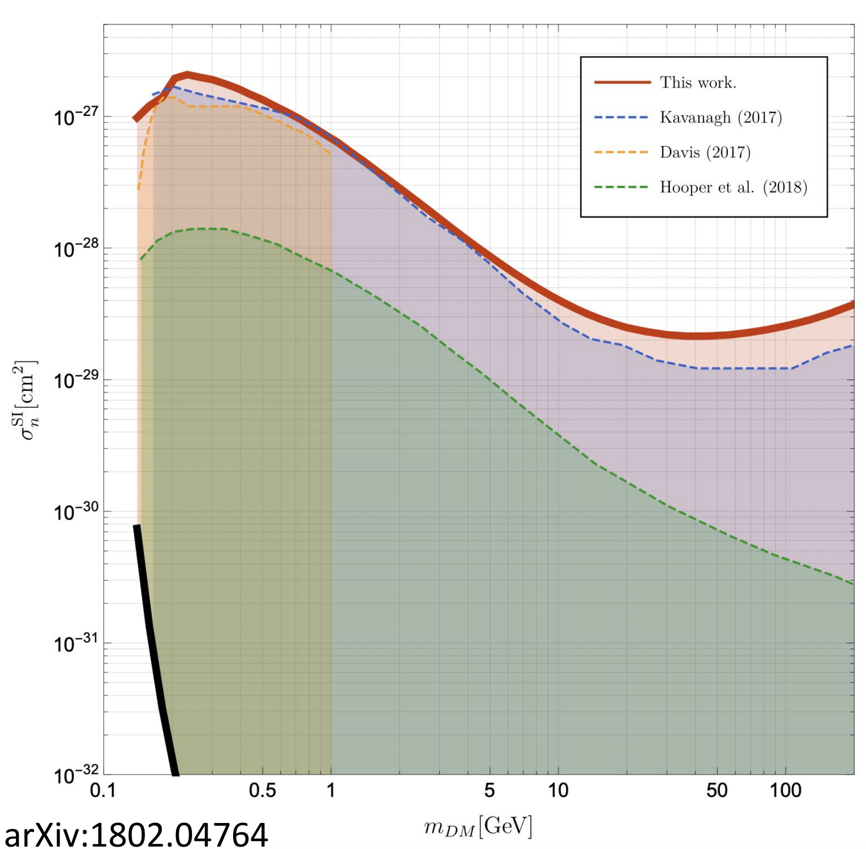


arXiv:1905.06348

A New Analytic Approximation

Based on arXiv:2301.07728

Probability of Not Scattering



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We want an alternative to both Monte Carlo and straight-line approaches

Hooper & Mcdermott (Green):

- Compute fraction of dark matter reaching CRESST surface detector without scattering

$$P_{\text{initial}}(x) = \frac{1}{l} e^{-x/l}$$

$$P(z, \theta) = \frac{1}{l \cos(\theta)} e^{-z/(l \cos(\theta))}$$

$$P(z) = \int_0^1 \frac{\mathrm{d}\cos(\theta)}{l \cos(\theta)} e^{-z/(l \cos(\theta))} = \frac{1}{l} \Gamma(0, z/l)$$

What about scattering?

Signatures of Earth-scattering in the direct detection of Dark Matter

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[c]CP^3-Origins, University of Southern Denmark,
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Received November 21, 2016
Accepted December 19, 2016
Published January 9, 2017

Abstract. Direct detection experiments search for the interactions of Dark Matter (DM) particles with nuclei in terrestrial detectors. But if these interactions are sufficiently strong, DM particles may scatter in the Earth, affecting their distribution in the lab. We present a new analytic calculation of this 'Earth-scattering' effect in the regime where DM particles scatter at most once before reaching the detector. We perform the calculation self-consistently, taking into account not only those particles which are scattered away from the detector, but also those particles which are deflected towards the detector. Taking into account a realistic model of the Earth and allowing for a range of DM-nucleon interactions, we present the EarthShadow code, which we make publicly available, for calculating the DM velocity distribution after Earth-scattering. Focusing on low-mass DM, we find that Earth-scattering reduces the direct detection rate at certain detector locations while increasing the rate in others. The Earth's rotation induces a daily modulation in the rate, which we find to be highly sensitive to the detector latitude and to the form of the DM-nucleon interaction. These distinctive signatures would allow us to unambiguously detect DM and perhaps even identify its interactions in regions of the parameter space within the reach of current and future experiments.

JCAP01(2017)012

What about scattering?

$$P_0(z) = \frac{1}{l}\Gamma(0, z/l)$$

$$P_1(z) = \int P_0(z')P(z - z')dz'$$

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Assume isotropic scattering

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$$P_n(z) = \int_0^\infty P_{n-1}(z')\frac{1}{2l}\Gamma(0, |z-z'|/l)dz'$$

This process is iterative!
Can compute probabilities through n scatterings

Energy Loss

$$P(\Delta E) = \frac{\sum_A \frac{n_A \sigma_{\chi A}}{E_{max,A}} \theta(E_{max,A} - \Delta E)}{\sum_A n_A \sigma_{\chi A}}$$

$$\frac{dN}{dE}_n(E) = \int \left(\frac{E}{E - \Delta E} \right) \frac{dN}{dE}_{n-1} \left(\frac{E^2}{E - \Delta E} \right) P(\Delta E) \, d\Delta E$$

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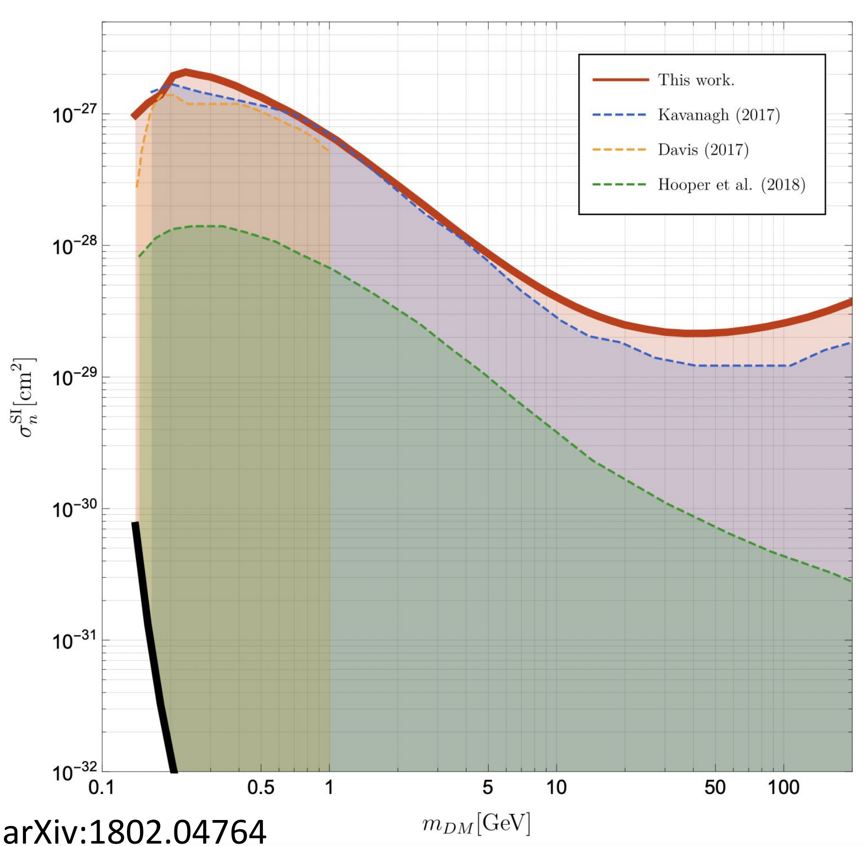
$$\frac{dN}{dE}_{total}(z, E) = \sum_{n=0}^{\infty} \int_z^{\infty} P_n(z') dz' \frac{dN}{dE}_n(E)$$

Velocity Distribution



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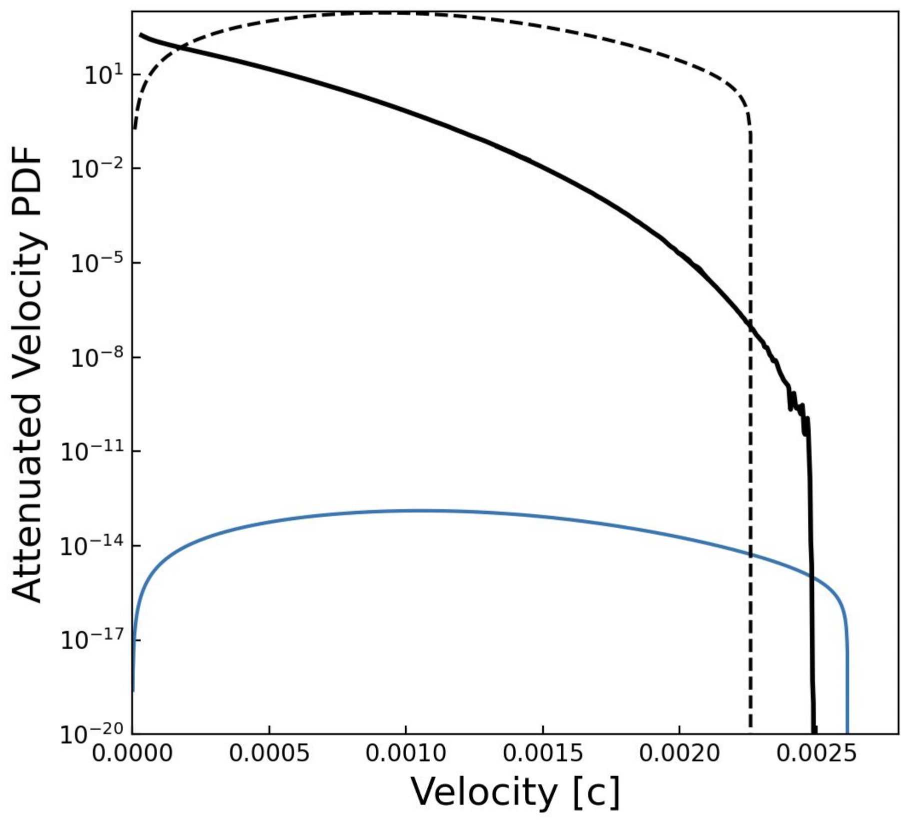
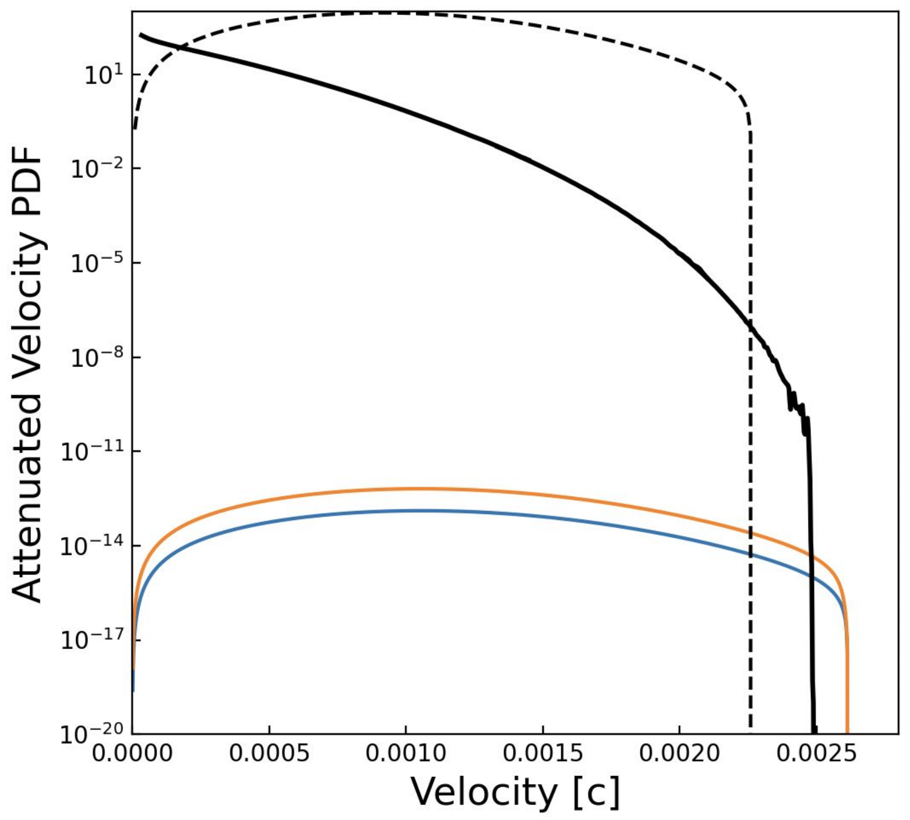


M_X = 100 MeV
σ_{XN} = $5*10^{-30}$ cm^2
Depth = 1400 m

Velocity Distribution



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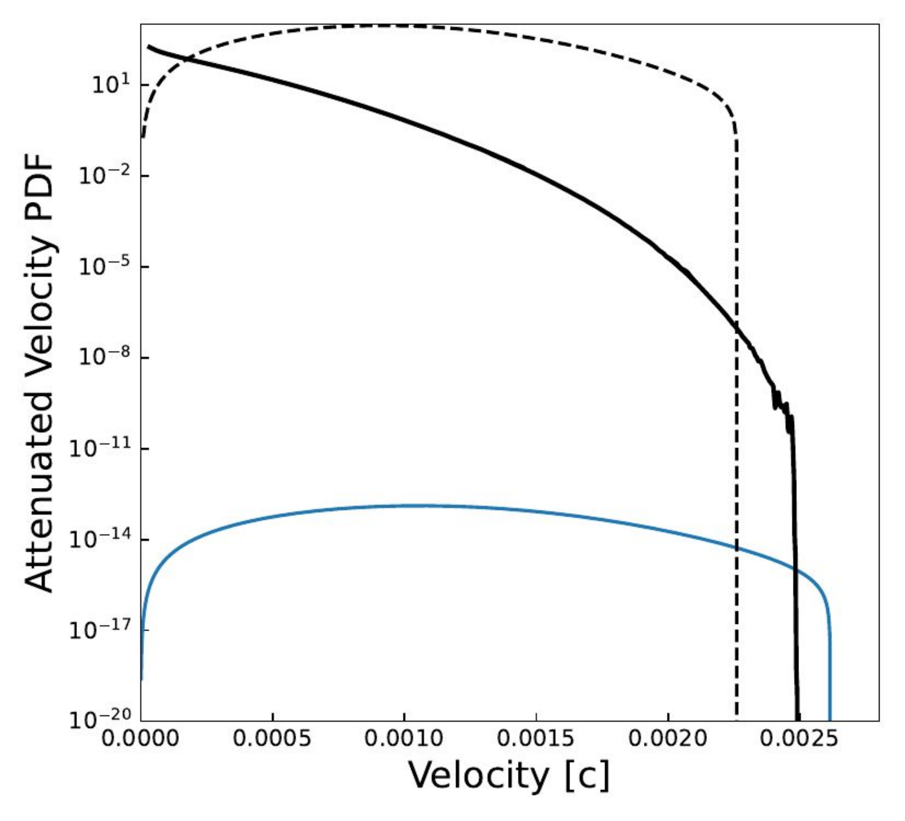

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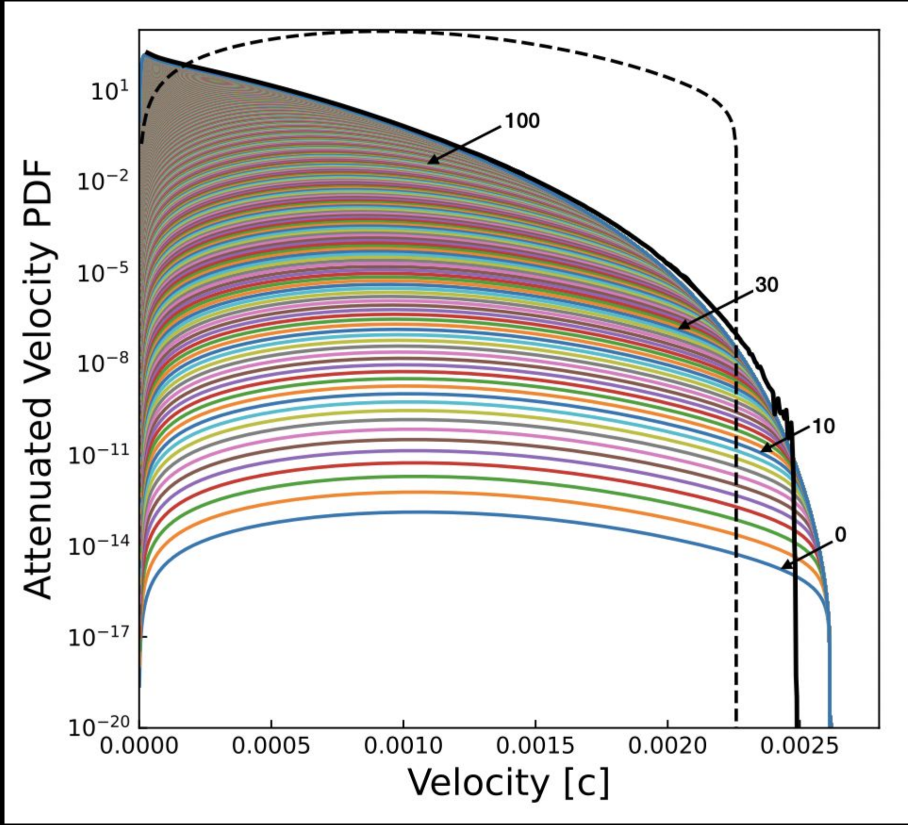

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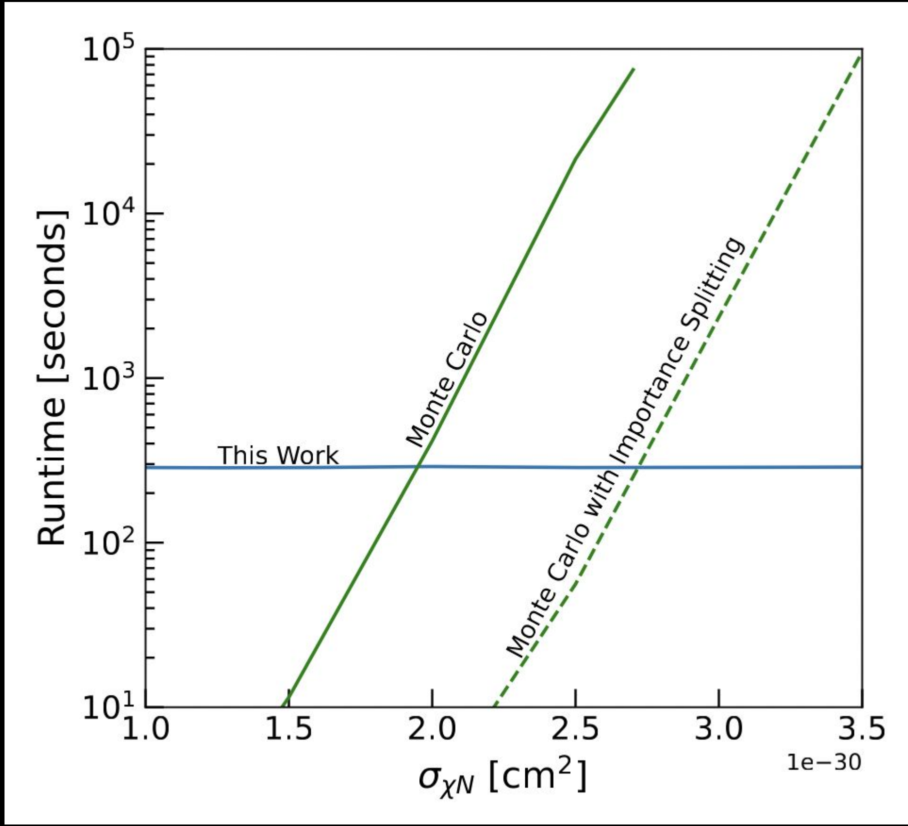
Runtime Comparison



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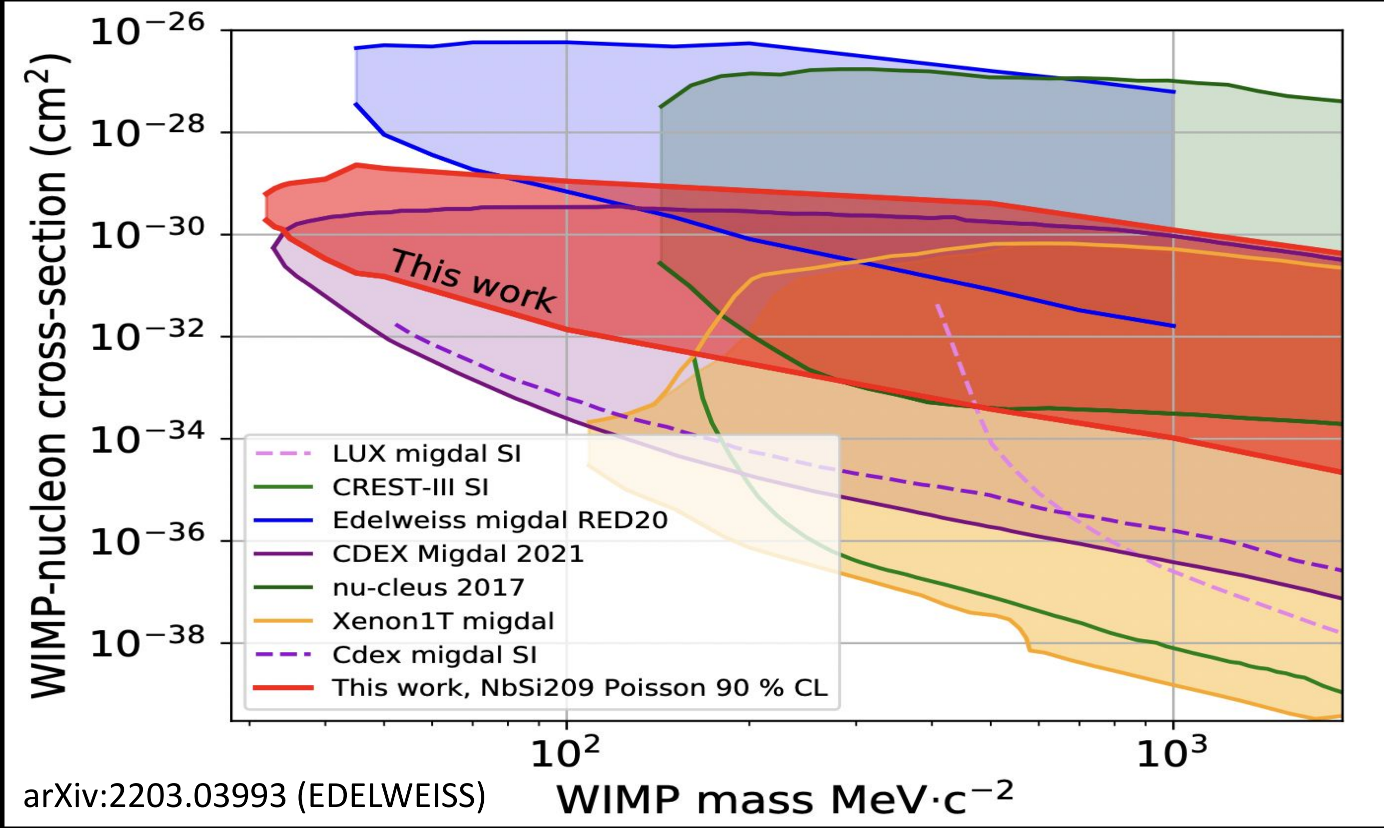
Attenuation: Straight Line Approximation



arXiv:2203.03993 (EDELWEISS)

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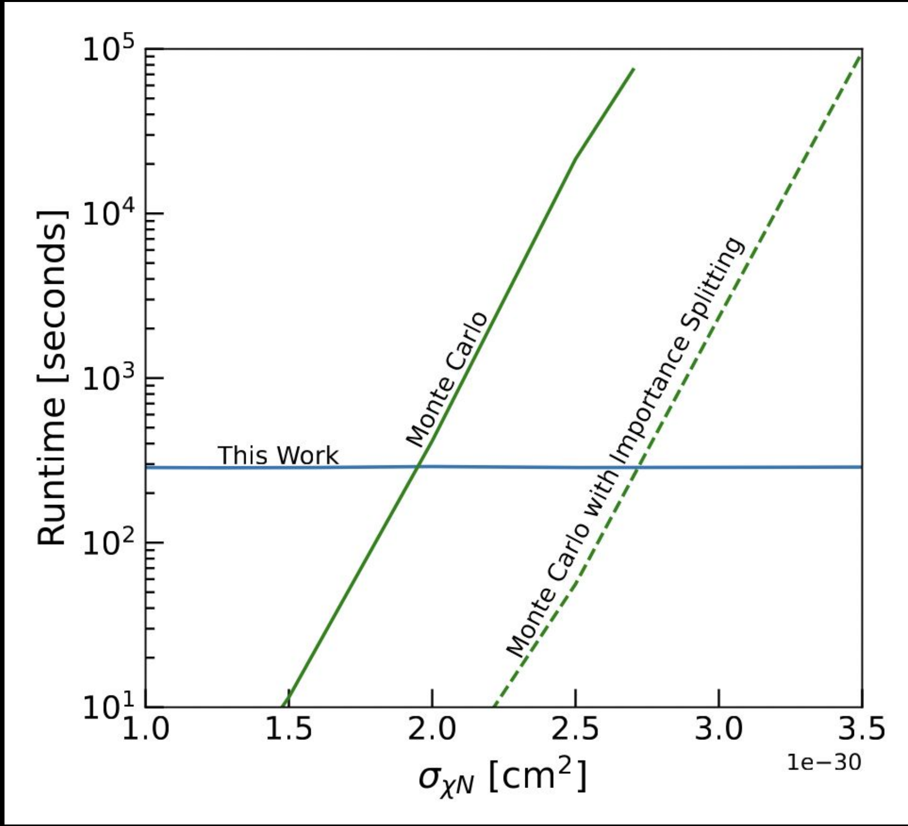

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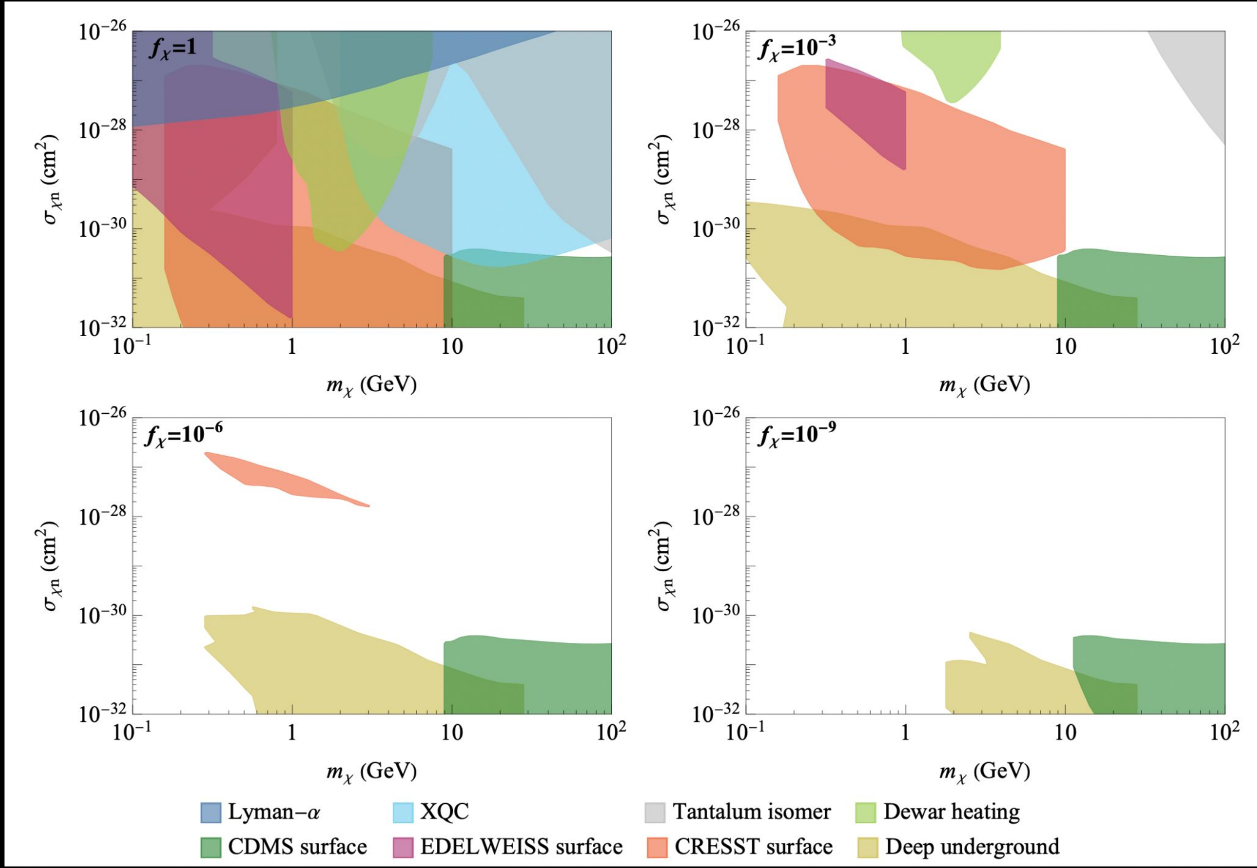
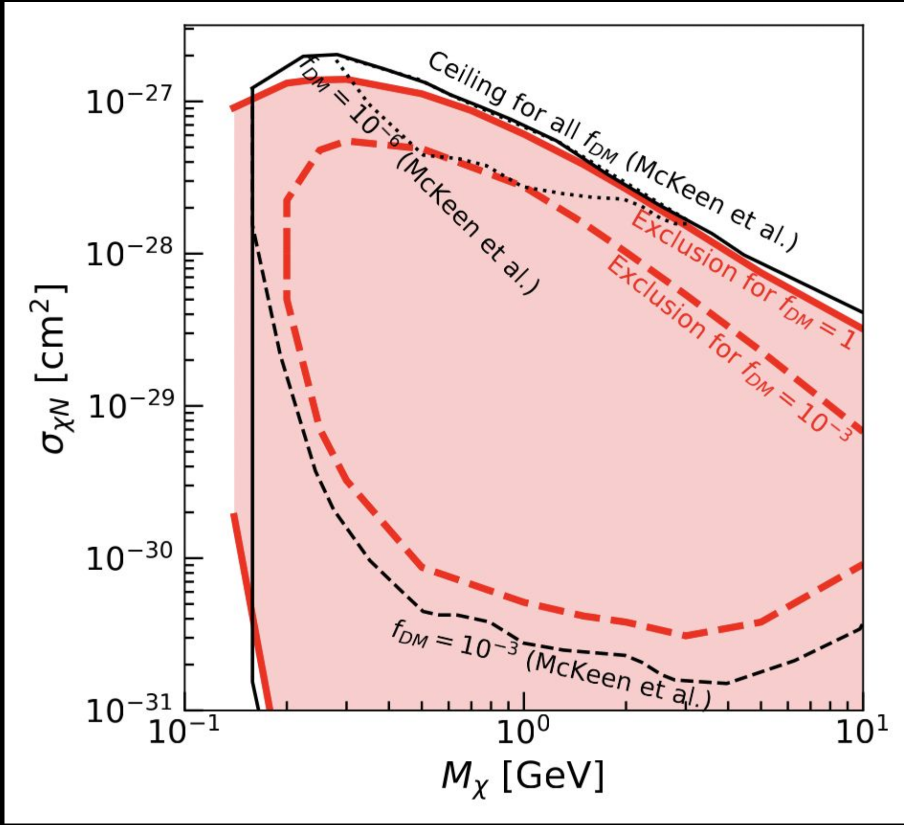
Recomputing Constraints on Sub-Dominant Dark Matter

Sub-Dominant Dark Matter

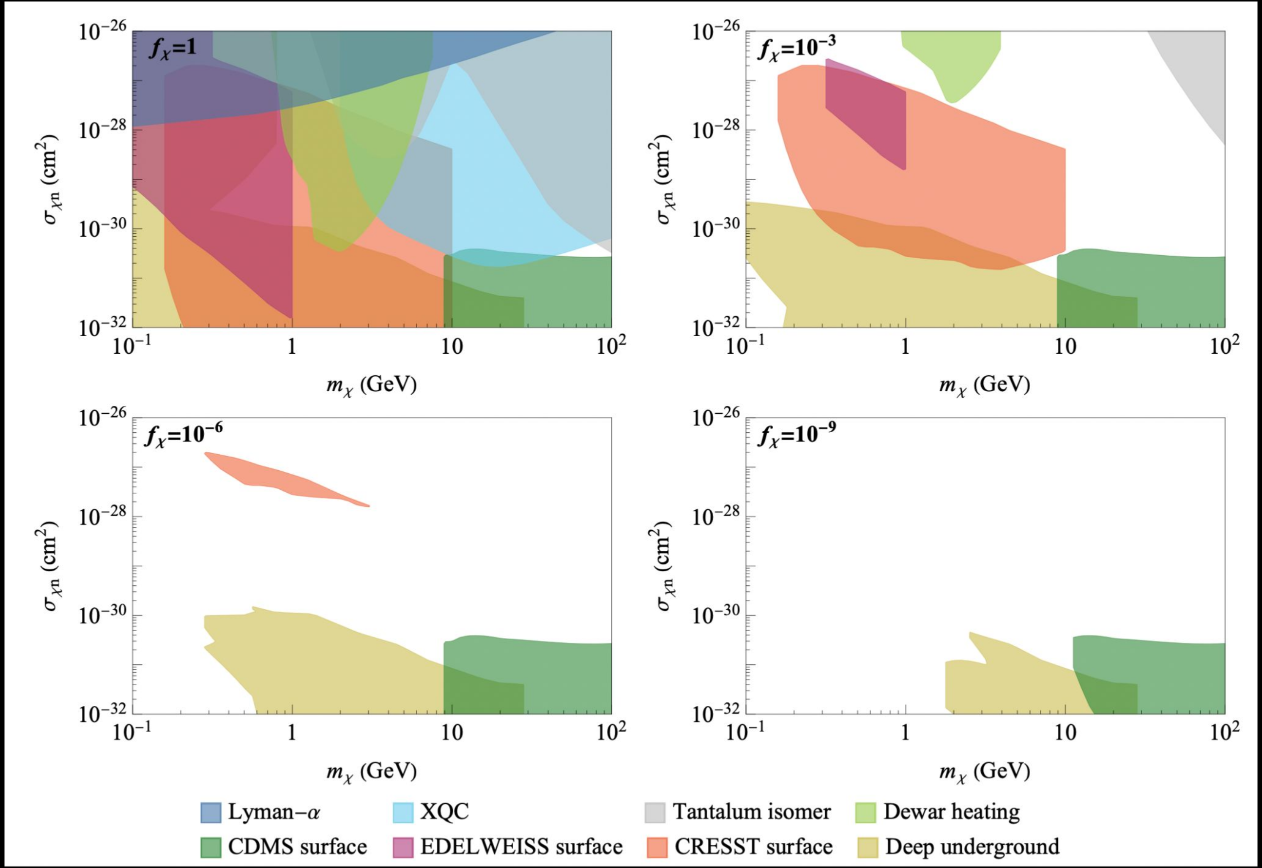


arXiv:2202.08840

Sub-Dominant Dark Matter



Sub-Dominant Dark Matter



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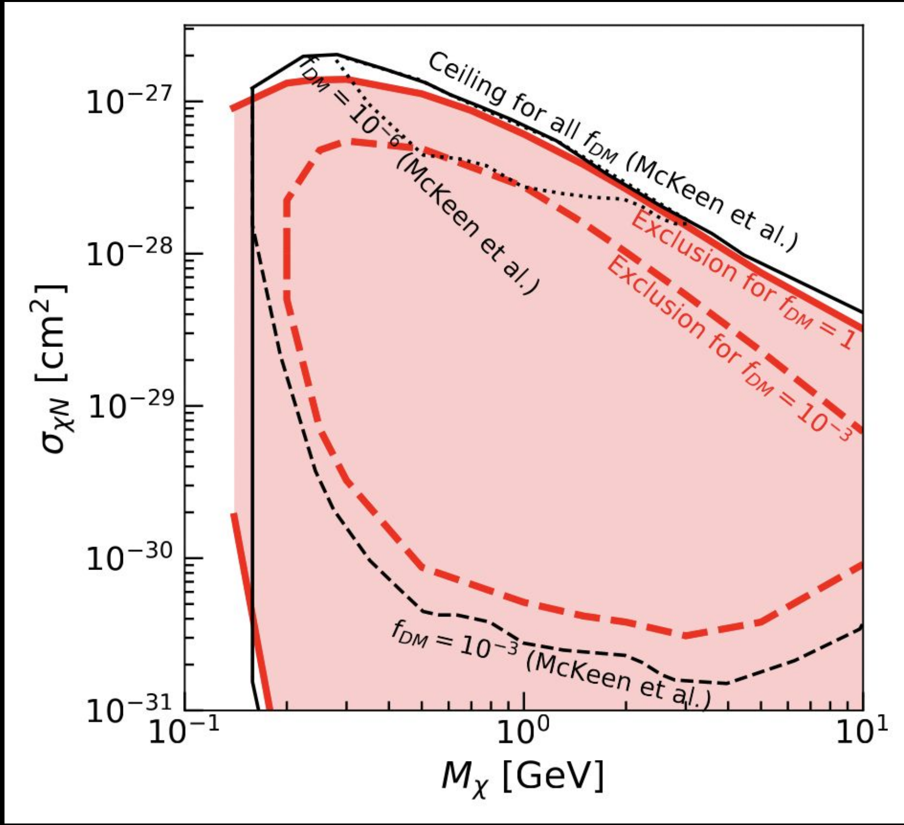

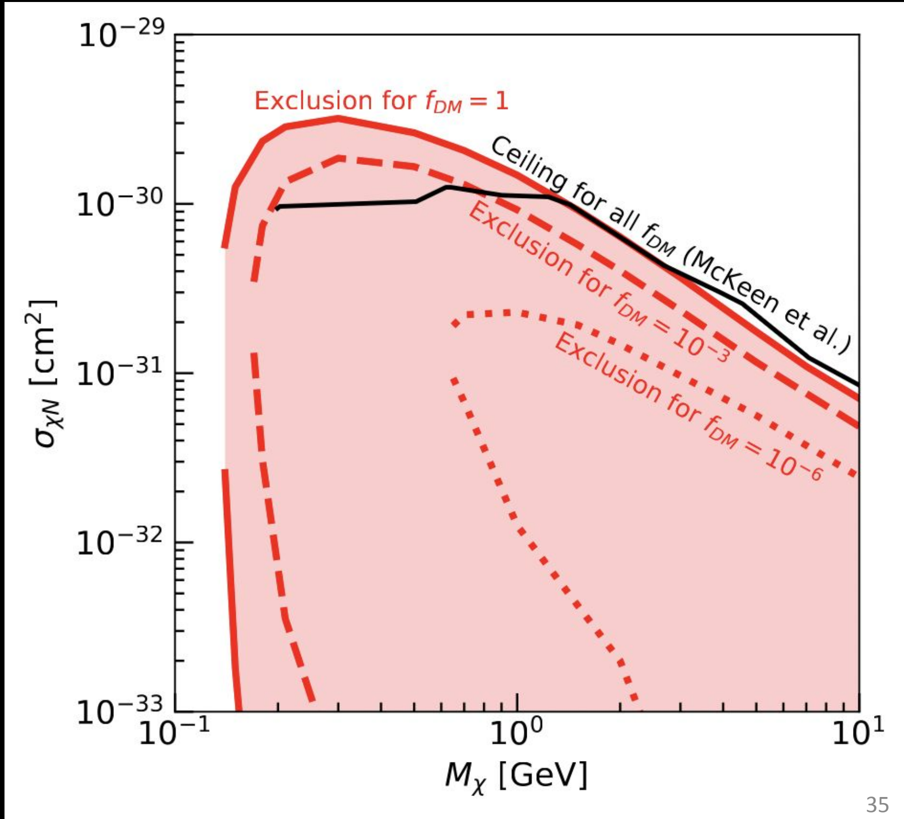

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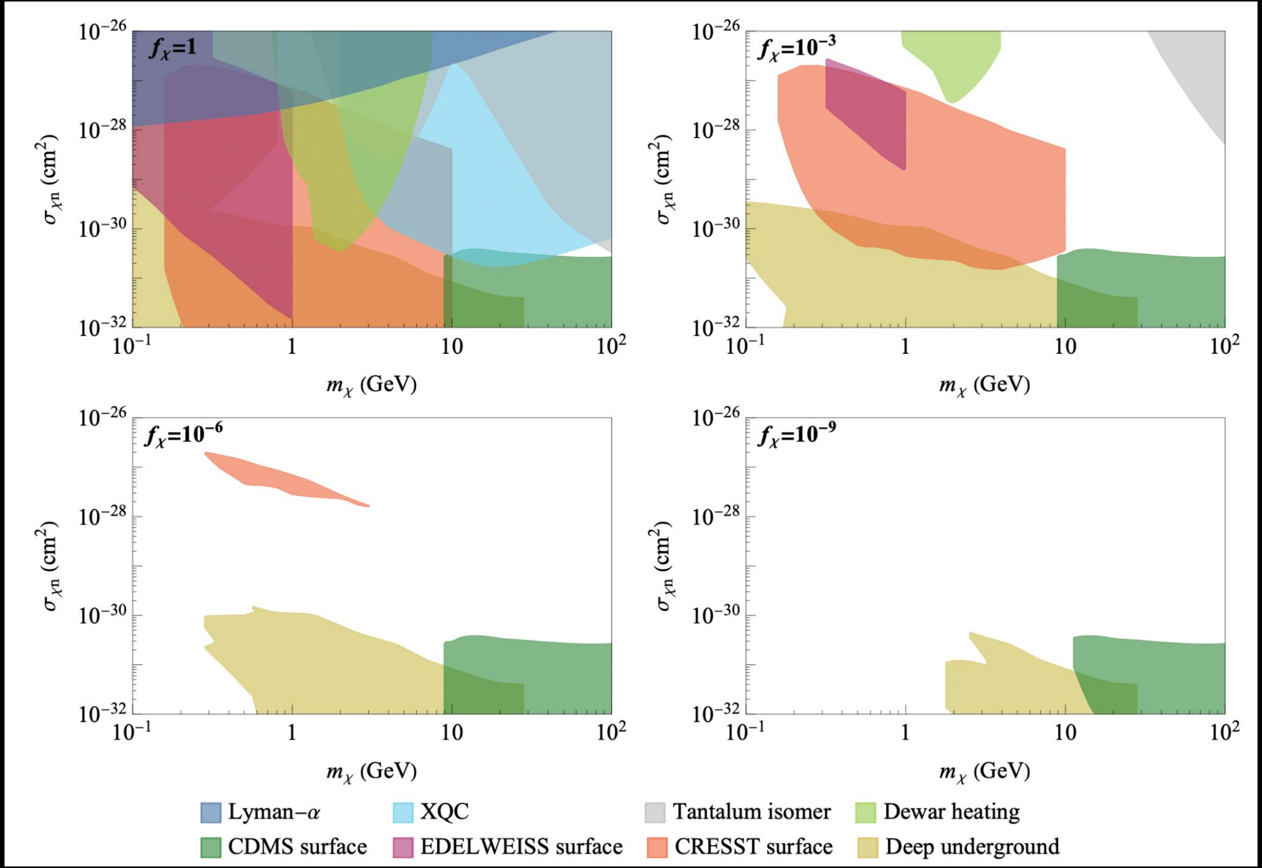

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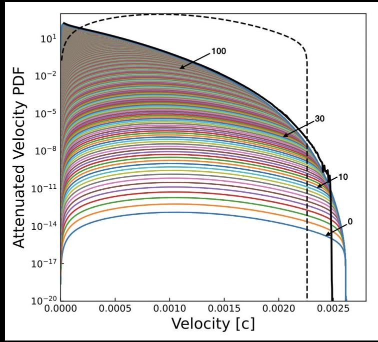
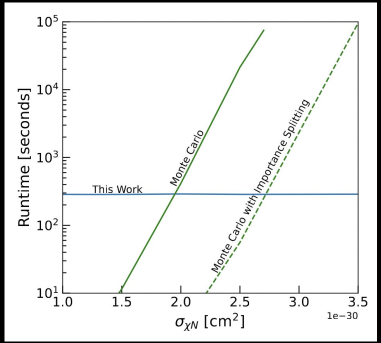
For light DM, this approach approximates Monte Carlo results much better than the straight line approximation

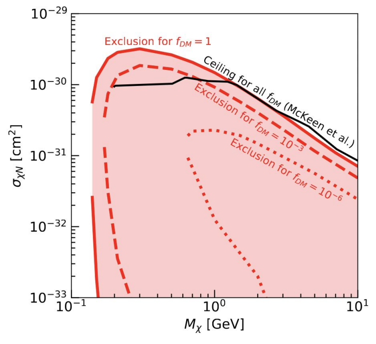




For large cross sections, this approach is much faster than Monte Carlo simulations

Can be used to efficiently recompute/recast exclusion regions from light dark matter searches

- Code: https://github.com/ccapp413/DMpropPublic



Thank you!

Many thanks to Timon Emken for details on the DaMaSCUS Monte Carlo code

Thank you to Ivan Esteban for suggestions on speeding up integration

Backup: 500 MeV, 10^{-30} cm^2